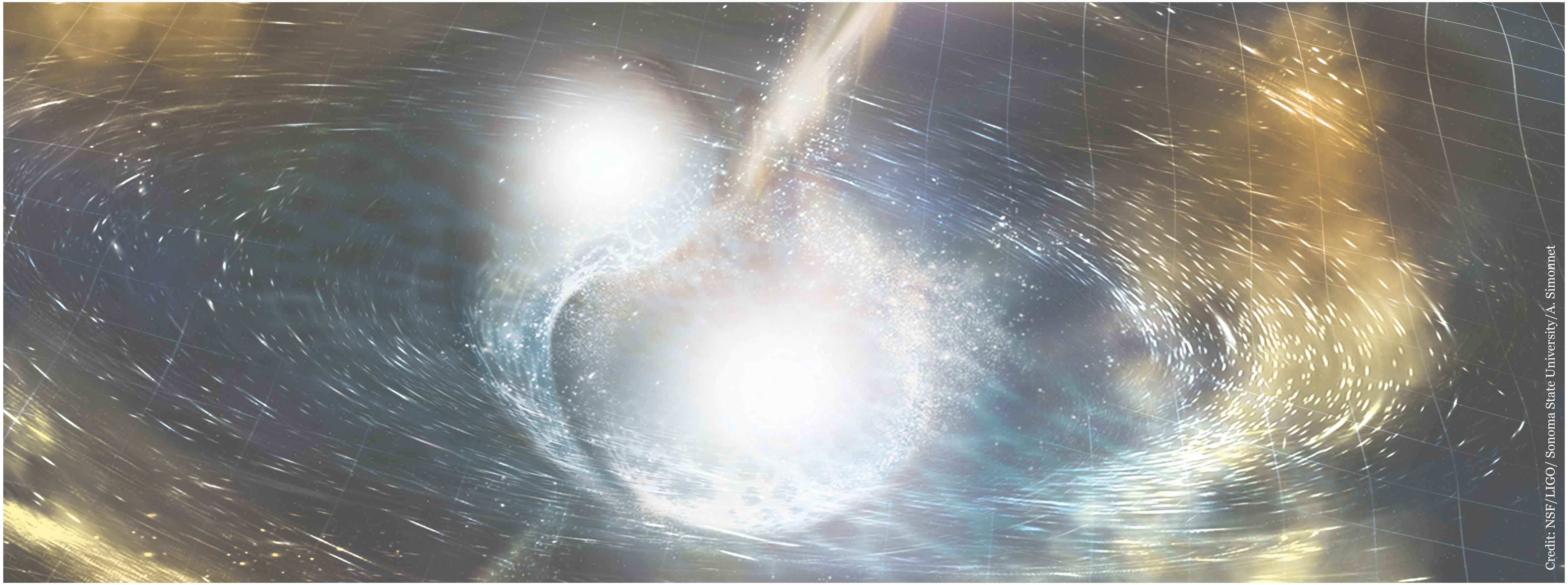
Stockholm University
Università degli Studi di Ferrara

Modelling kilonovae from neutron star mergers with POSSIS

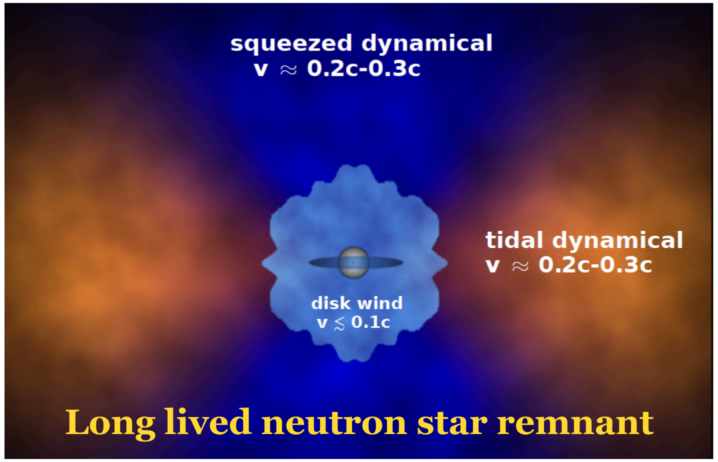
Mattia Bulla

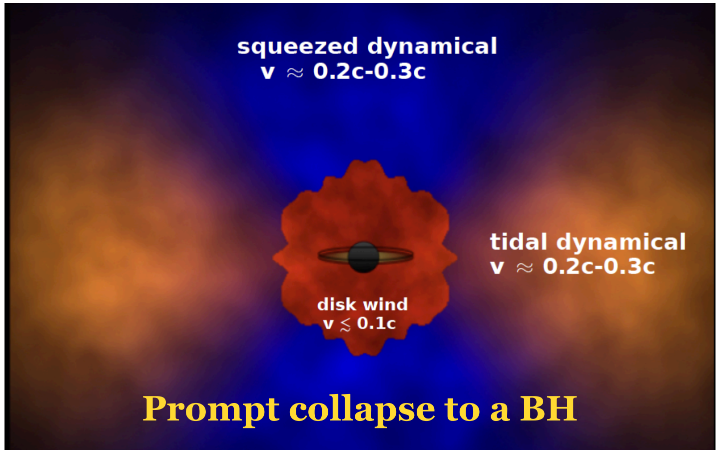
with: A. Sagues-Carracedo, S. Anand, P. T. H. Pang, L. Nativi, M. W. Coughlin, T. Dietrich, I. Tews, M. Shrestha, I. Andreoni, S. Dhawan, K. Mooley, A. Goobar, S. Rosswog, S. Covino, M. Tanaka, K. Kyutoku + many more

Oskar Klein centre
G.R.E.A.T. VR Research Environment
Vetenskapsrådet
Swedish Research Council

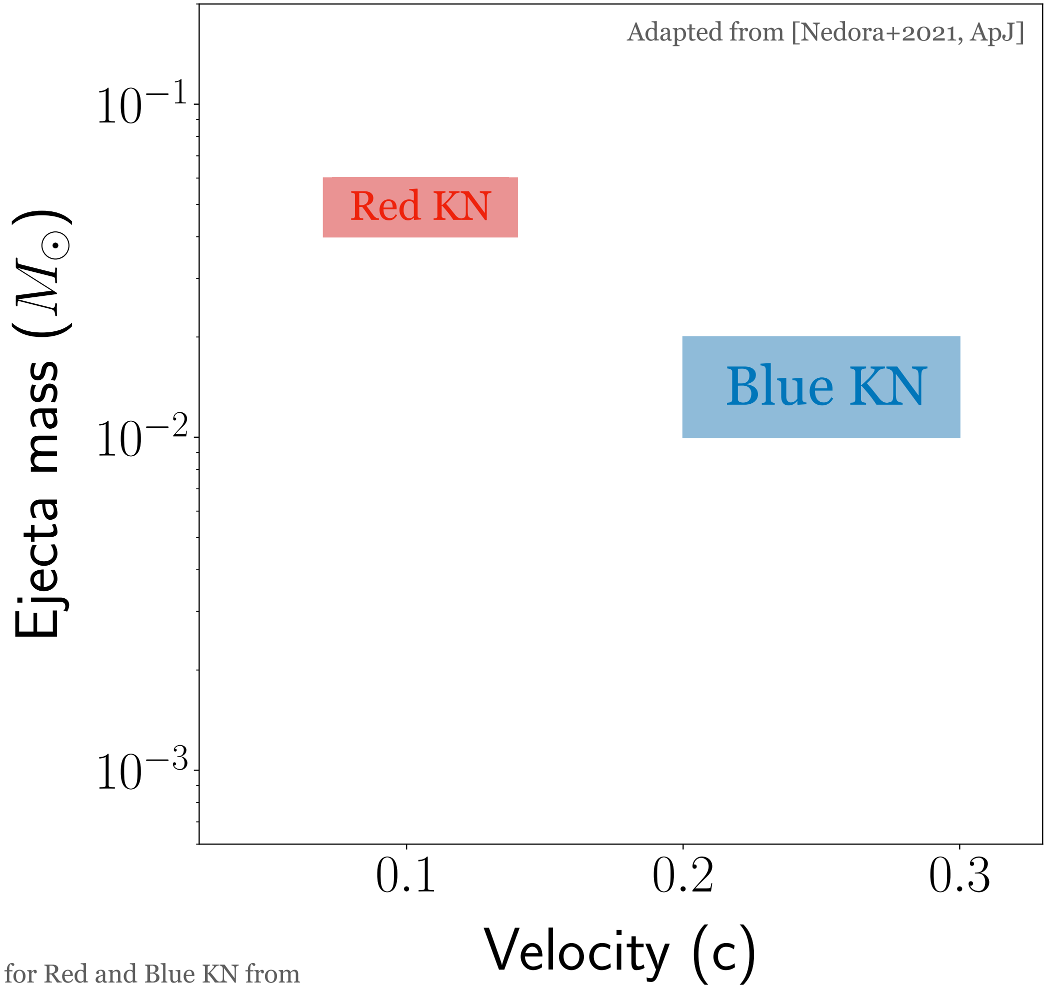


NS mergers and kilonovae



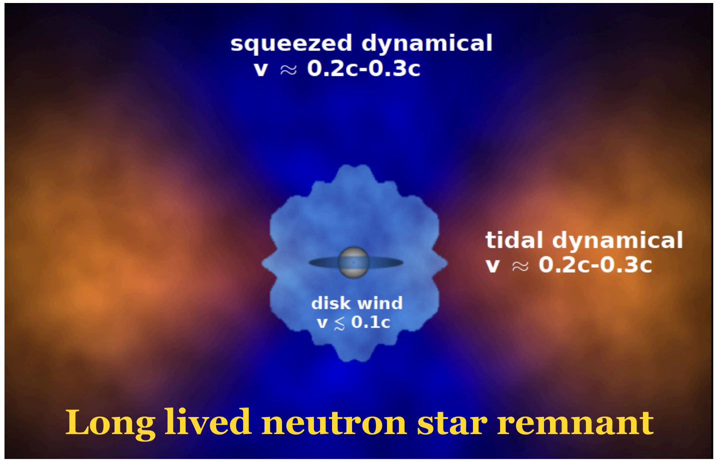


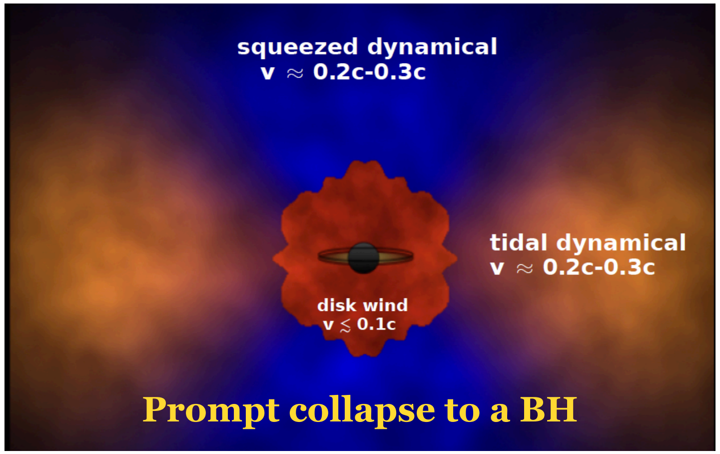
[Kasen+2017, Nature]



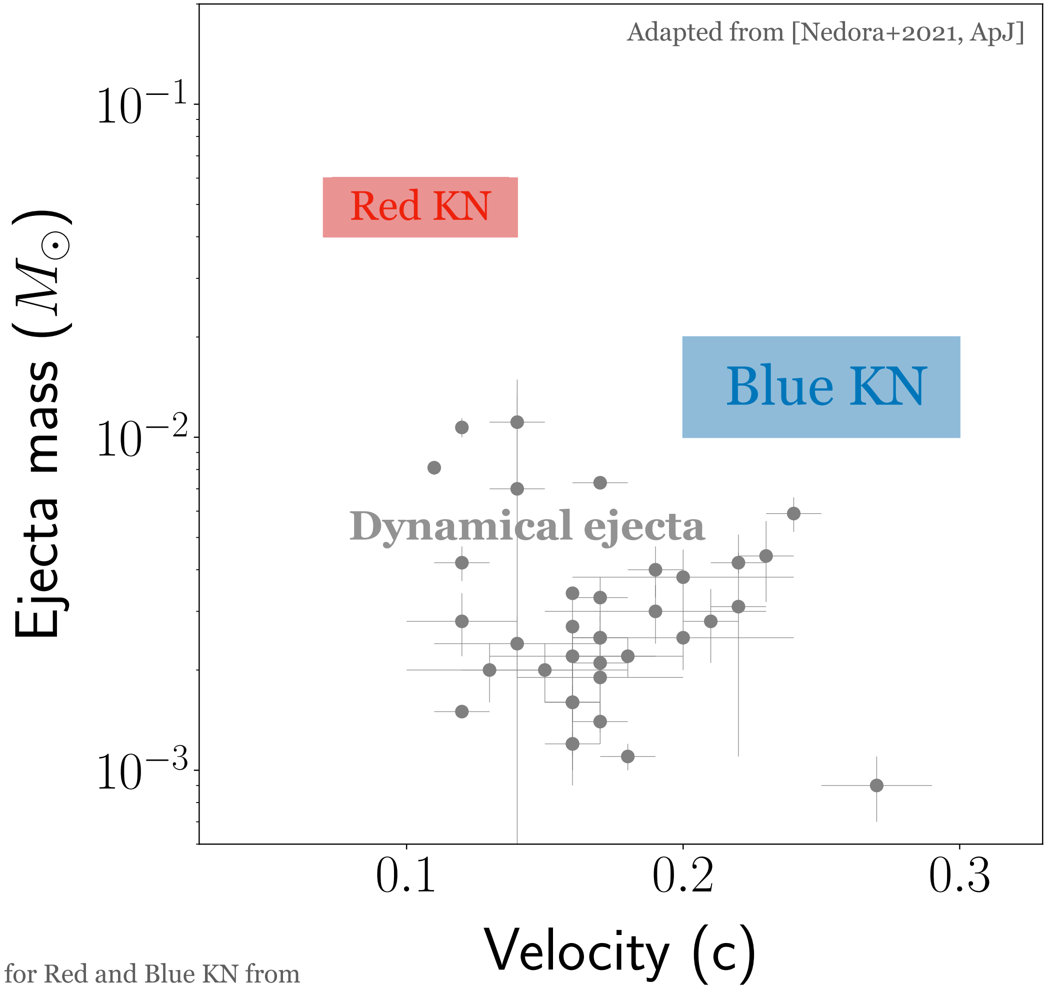
Parameters for Red and Blue KN from
[Siegel 2019, Eur. Phys. J. A.]

NS mergers and kilonovae



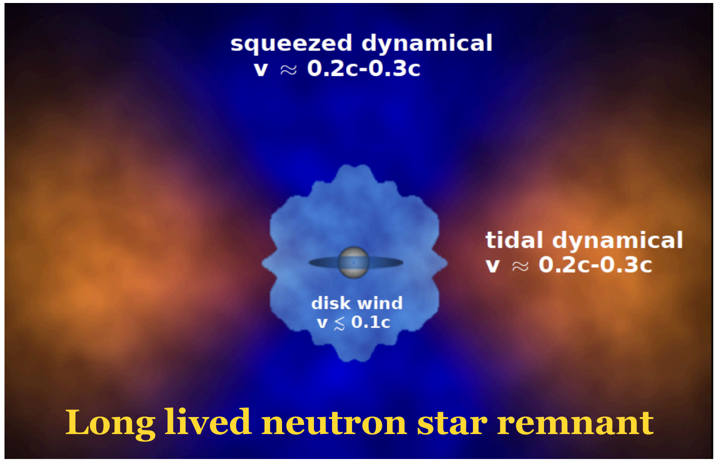


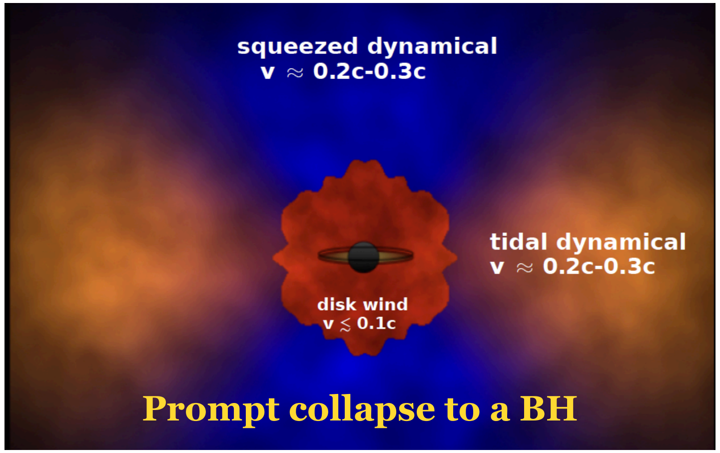
[Kasen+2017, Nature]



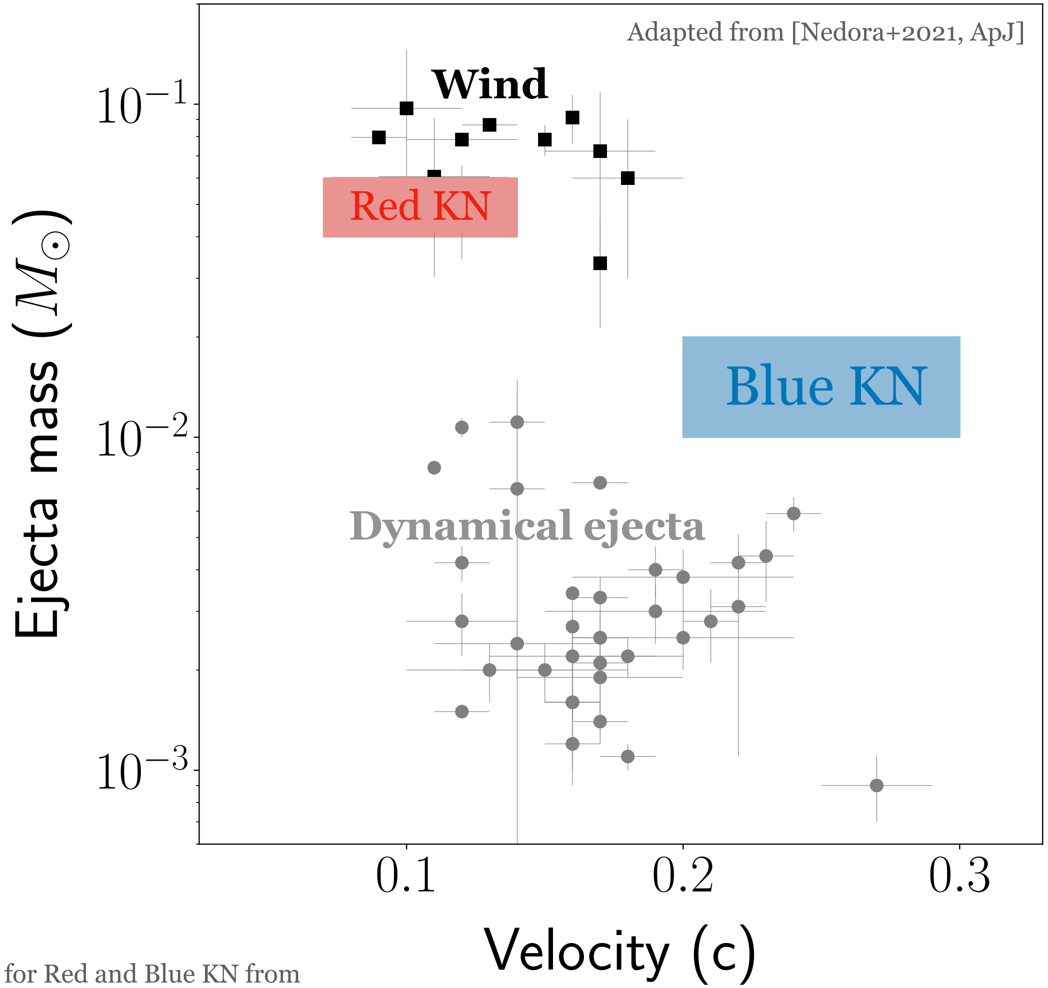
Parameters for Red and Blue KN from
[Siegel 2019, Eur. Phys. J. A.]

NS mergers and kilonovae



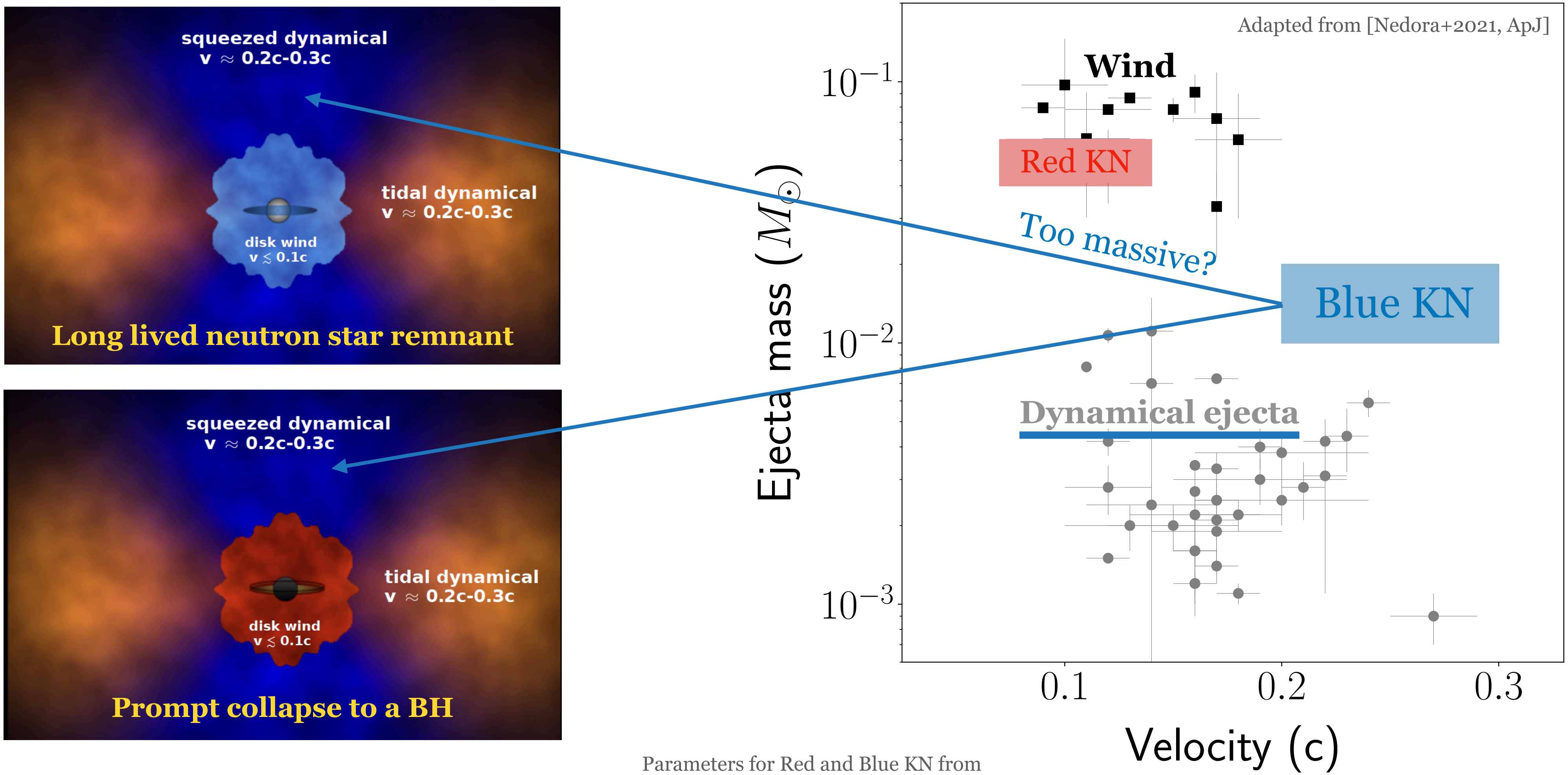


[Kasen+2017, Nature]



Parameters for Red and Blue KN from
[Siegel 2019, Eur. Phys. J. A.]

NS mergers and kilonovae

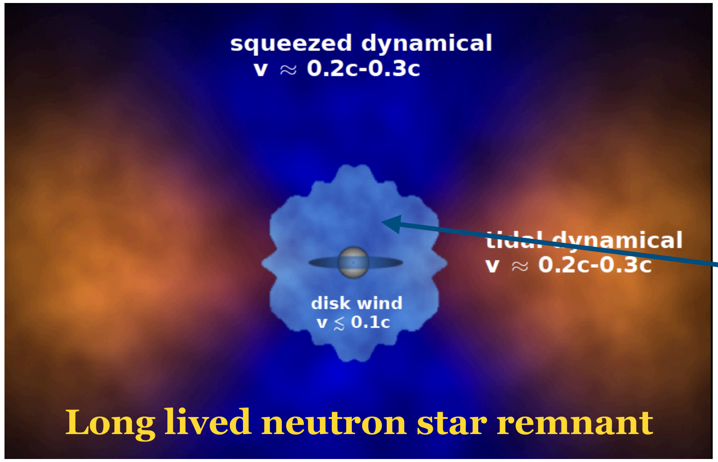


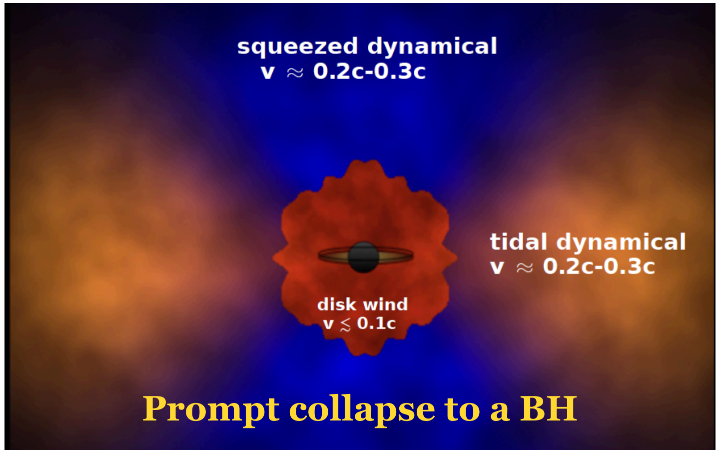
[Kasen+2017, Nature]

Parameters for Red and Blue KN from
[Siegel 2019, Eur. Phys. J. A.]

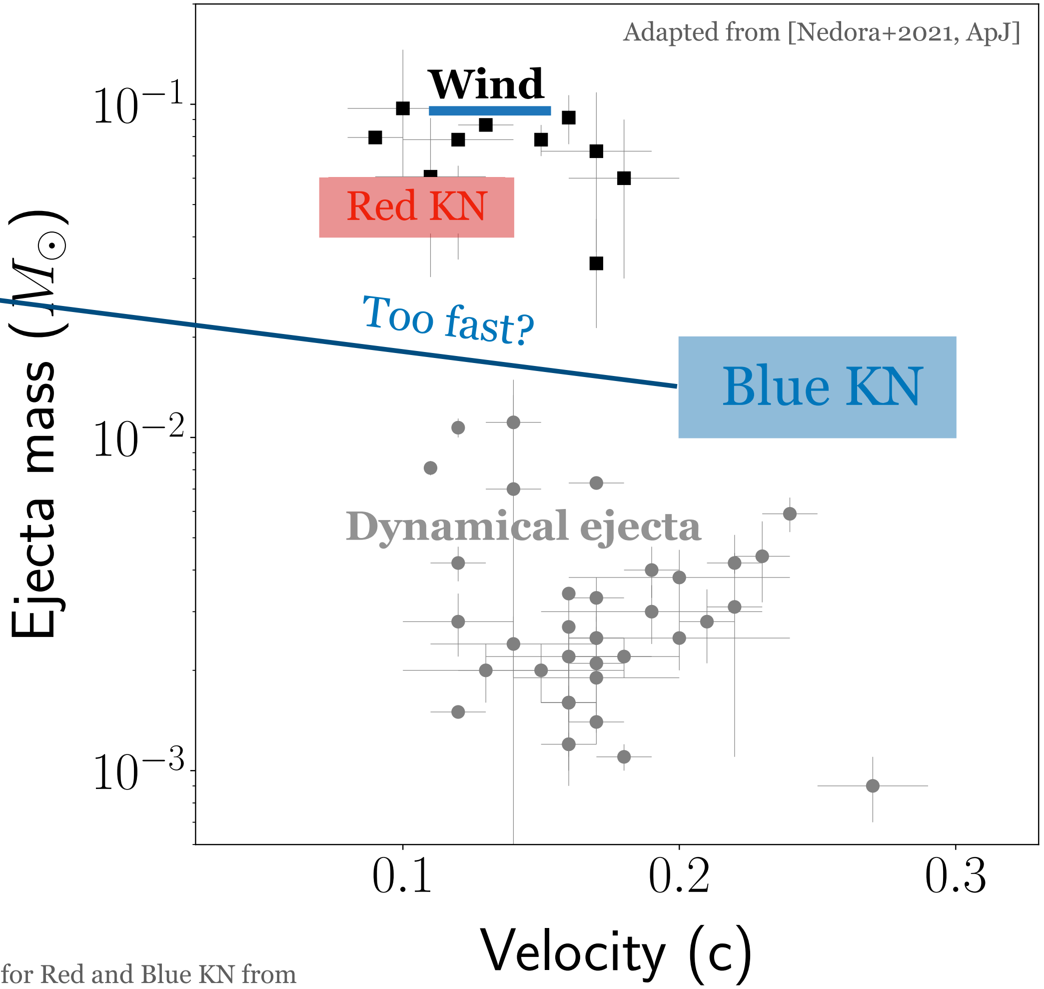
NS mergers and kilonovae

See Steven Fahlman's talk



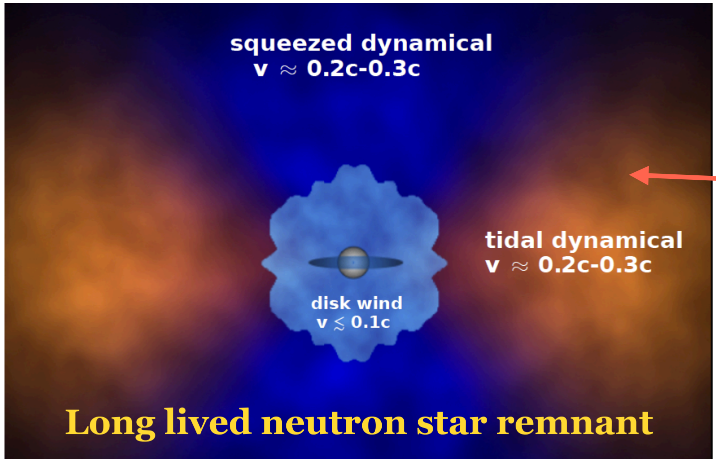


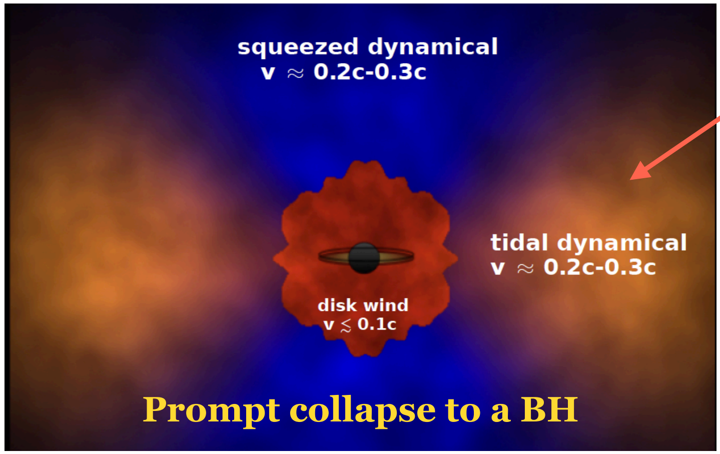
[Kasen+2017, Nature]



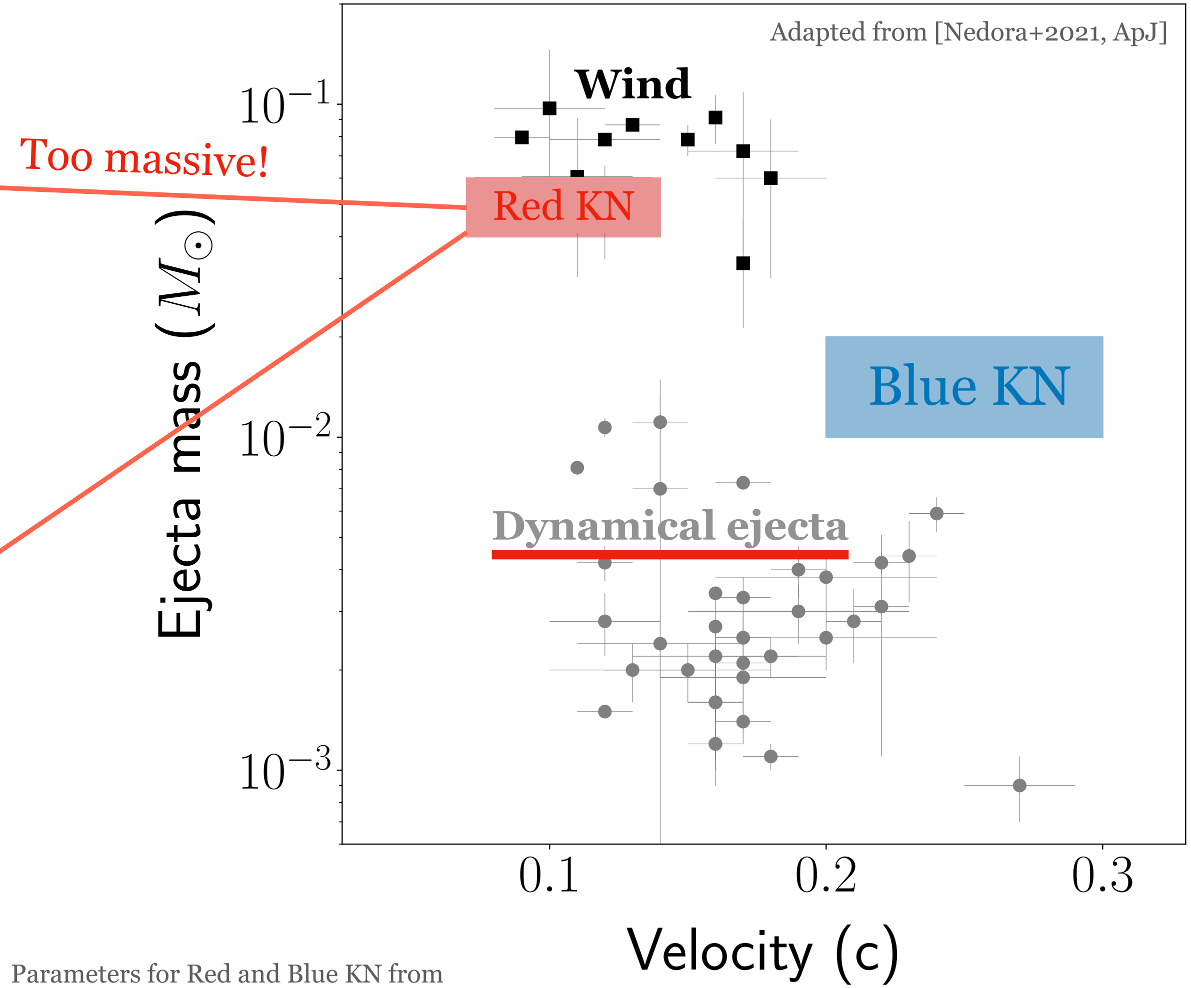
Parameters for Red and Blue KN from
[Siegel 2019, Eur. Phys. J. A.]

NS mergers and kilonovae



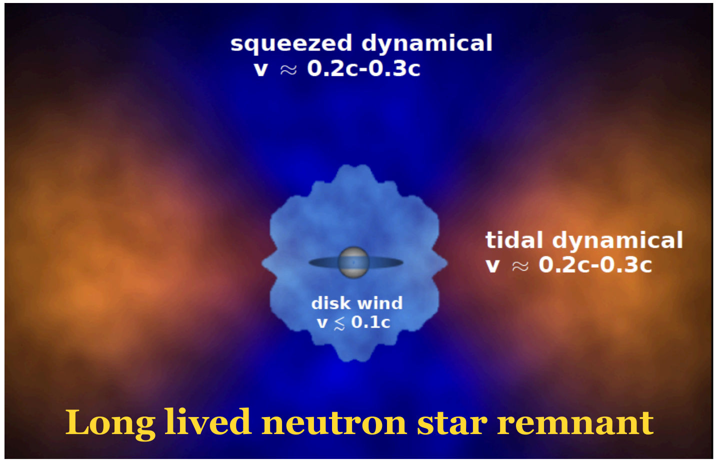


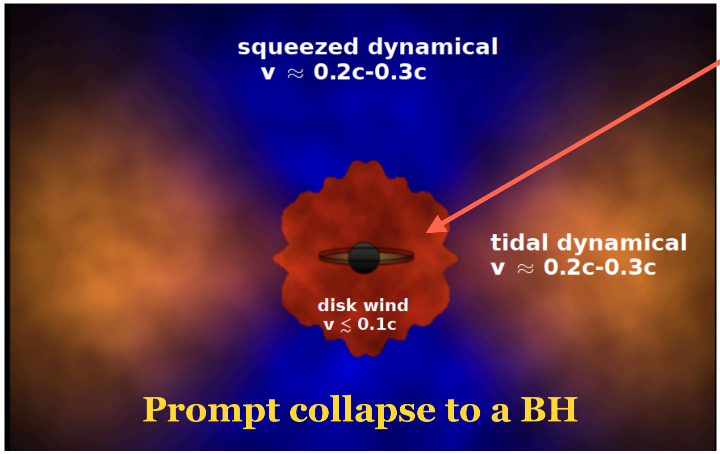
[Kasen+2017, Nature]



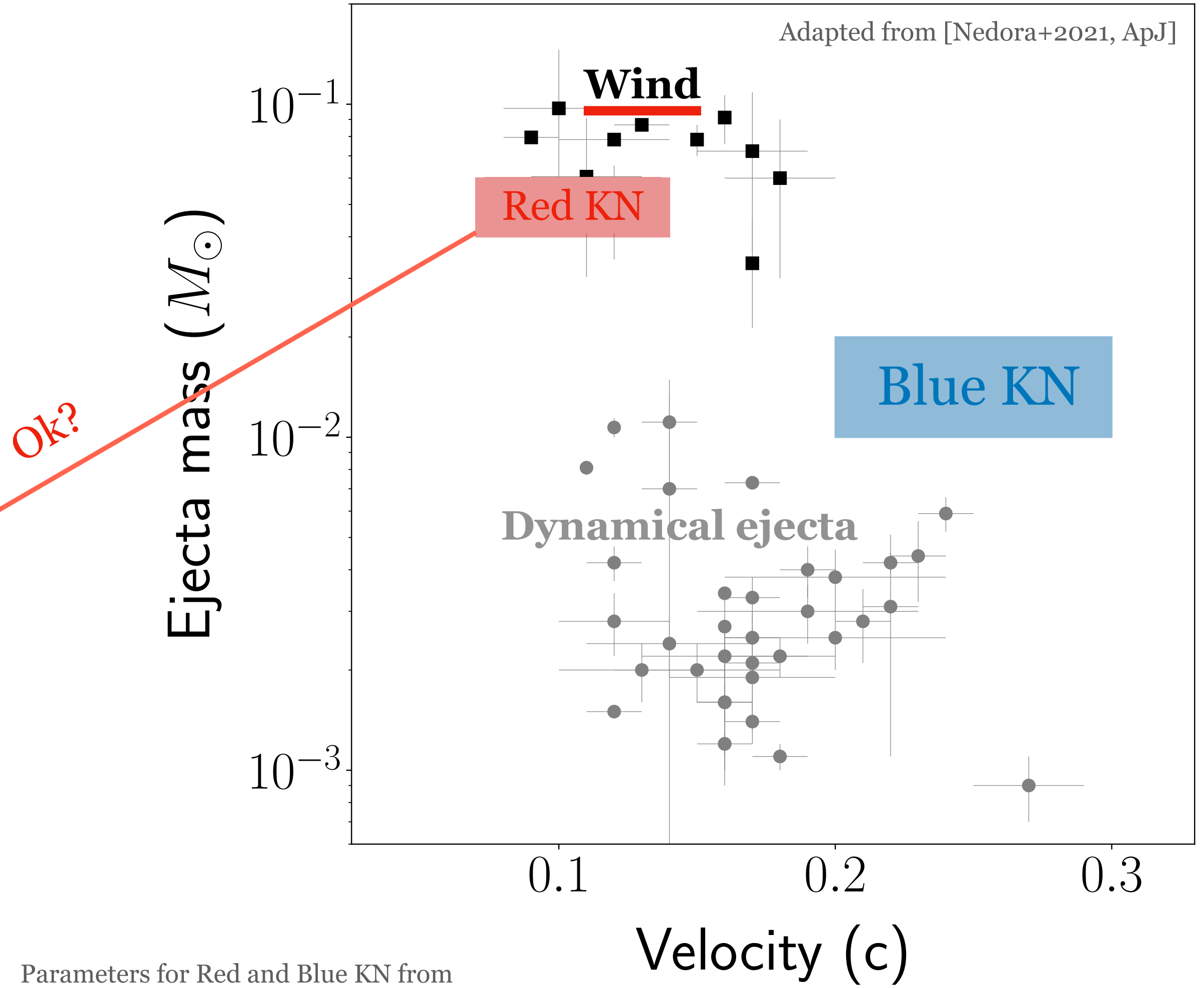
Parameters for Red and Blue KN from
[Siegel 2019, Eur. Phys. J. A.]

NS mergers and kilonovae



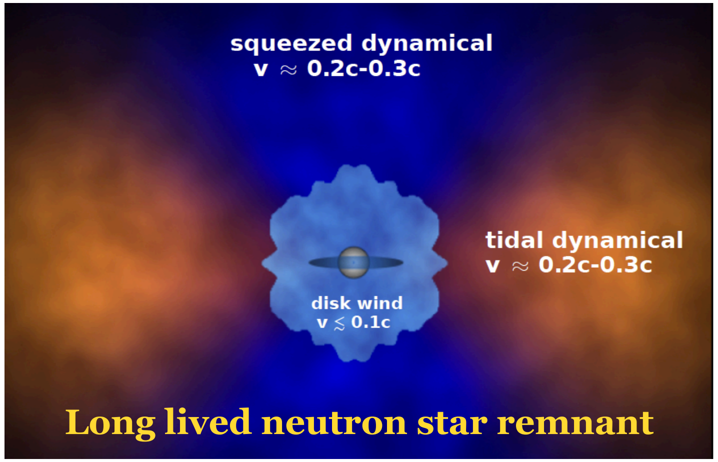


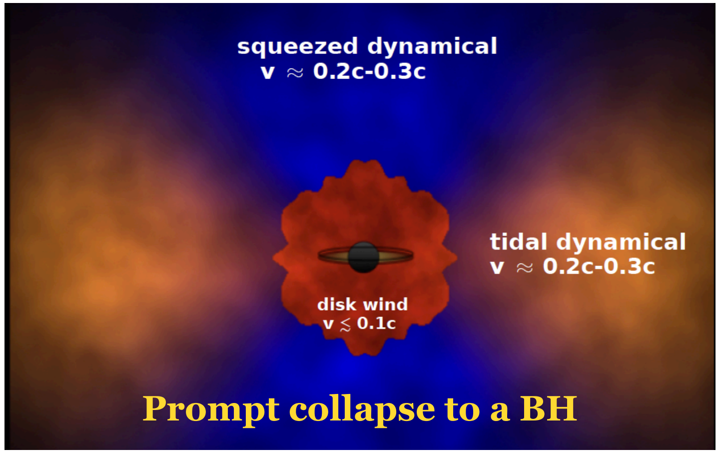
[Kasen+2017, Nature]



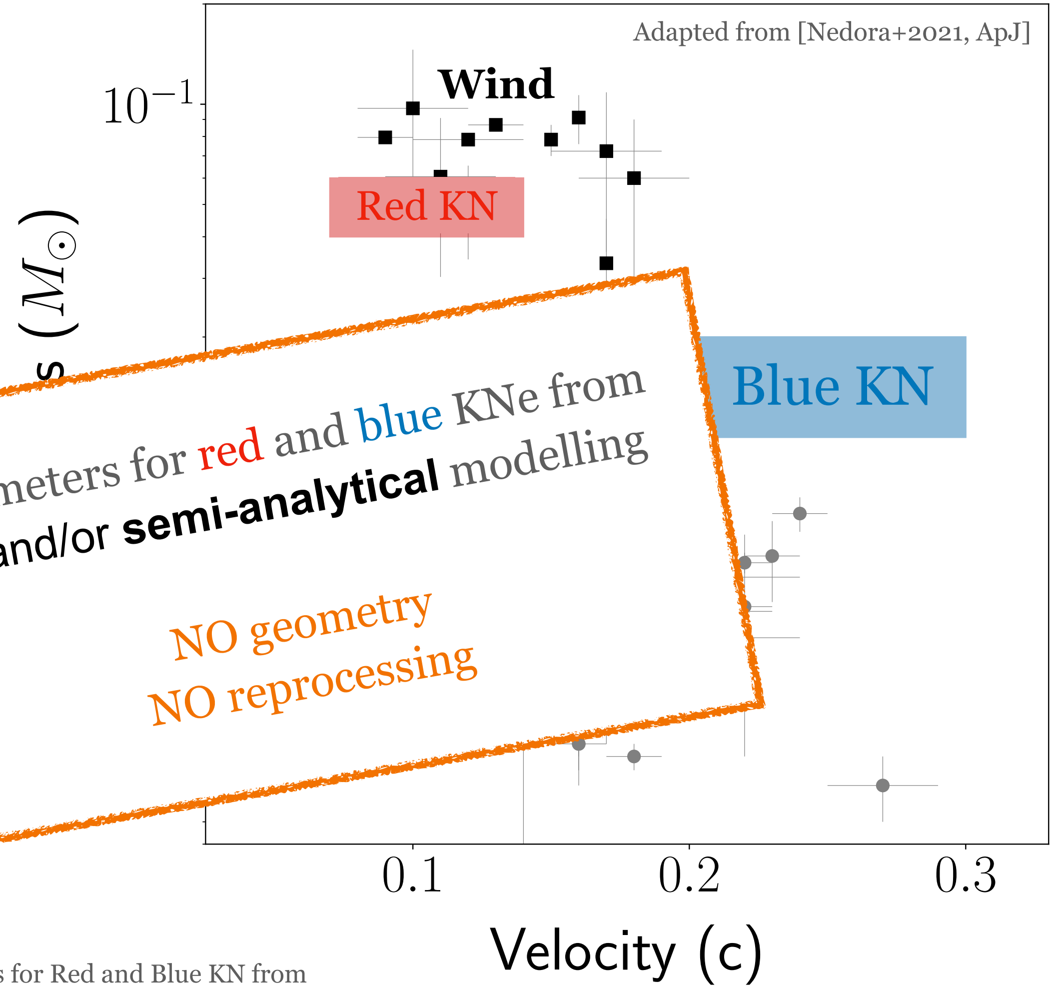
Parameters for Red and Blue KN from
[Siegel 2019, Eur. Phys. J. A.]

NS mergers and kilonovae





[Kasen+2017, Nature]



Parameters for red and blue KNe from **1D** and/or **semi-analytical** modelling

NO geometry
NO reprocessing

Parameters for Red and Blue KN from
[Siegel 2019, Eur. Phys. J. A.]

POSSIS

A 3D Monte Carlo radiative transfer code to model kilonovae

[**MB**+2015, MNRAS; **MB** 2019, MNRAS]

POSSIS

A 3D Monte Carlo radiative transfer code to model kilonovae

[**MB**+2015, MNRAS; **MB** 2019, MNRAS]

Creating photons

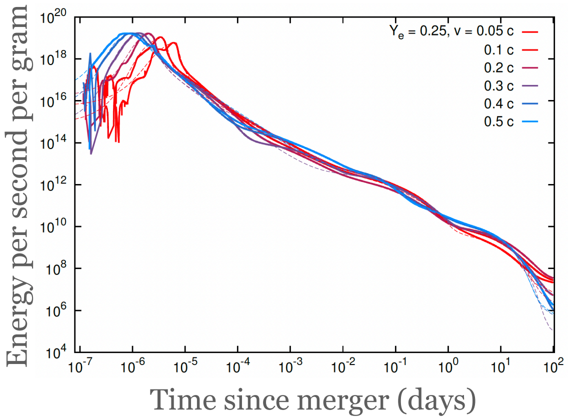
Frequency

From temperature + opacity

Energy

Nuclear heating rates

Thermalisation efficiencies



Stokes parameters

POSSIS

A 3D Monte Carlo radiative transfer code to model kilonovae

[**MB**+2015, MNRAS; **MB** 2019, MNRAS]

Creating photons

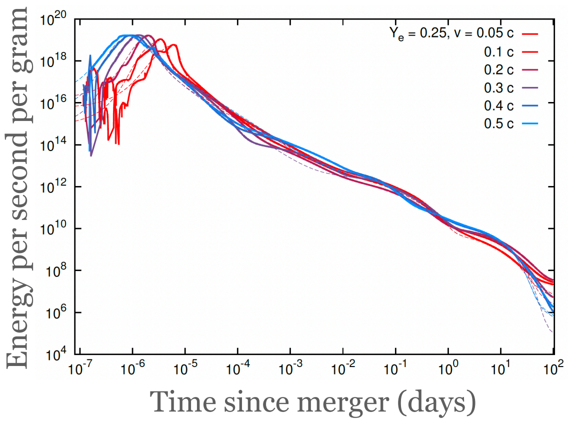
Frequency

From temperature + opacity

Energy

Nuclear heating rates

Thermalisation efficiencies



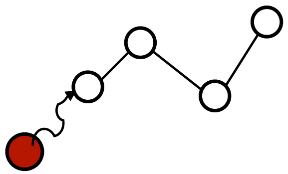
Stokes parameters

Heating rates

from Rosswog & Korobkin 2022

as a function of Y_e and velocity

$$\frac{d\varepsilon}{dt} = \dot{\varepsilon}_0 \left(\frac{1}{2} - \frac{1}{\pi}\arctan\left[\frac{t-t_0}{\sigma}\right]\right)^{\alpha} \left(\frac{1}{2} + \frac{1}{\pi}\arctan\left[\frac{t-t_1}{\sigma_1}\right]\right)^{\alpha_1} + C_1 e^{-t/\tau_1} + C_2 e^{-t/\tau_2} + C_3 e^{-t/\tau_3} \quad (2)$$



POSSIS

A 3D Monte Carlo radiative transfer code to model kilonovae

[**MB**+2015, MNRAS; **MB** 2019, MNRAS]

Creating photons

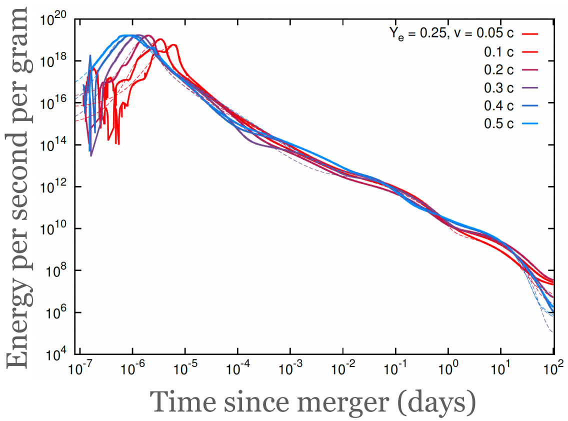
Frequency

From temperature + opacity

Energy

Nuclear heating rates

Thermalisation efficiencies



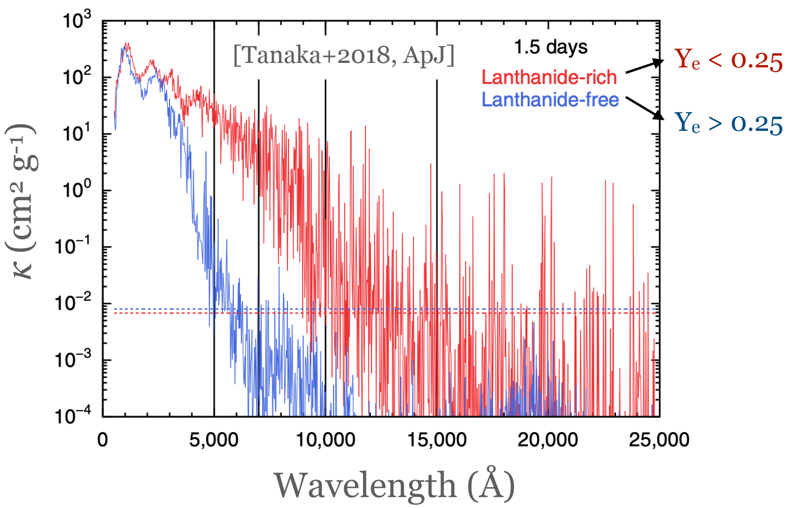
Stokes parameters

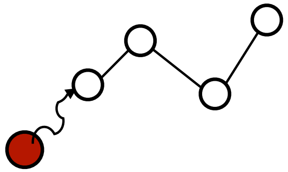
Propagating photons

Opacity

$$\tau = \int \kappa \, \rho \, dr \qquad P_{\mathrm{interaction}} = 1 - e^{-\tau}$$

Main source of opacity in KNe: bound-bound





POSSIS

A 3D Monte Carlo radiative transfer code to model kilonovae

[**MB**+2015, MNRAS; **MB** 2019, MNRAS]

Creating photons

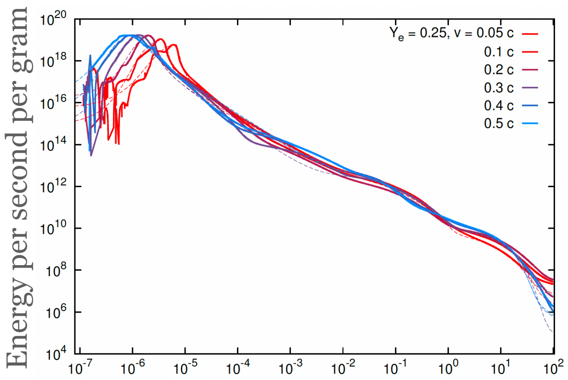
Frequency

From temperature + opacity

Energy

Nuclear heating rates

Thermalisation efficiencies



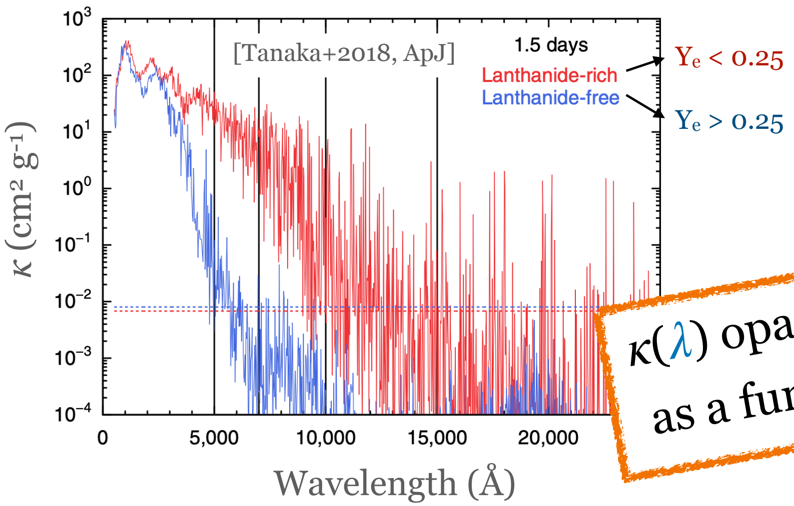
Stokes parameters

Propagating photons

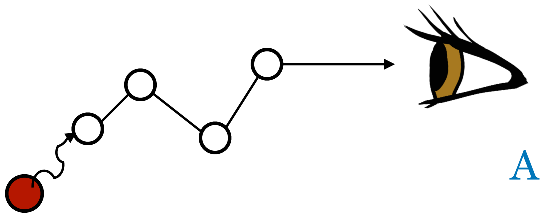
Opacity

$$\tau = \int \kappa \rho \, dr \qquad P_{\text{interaction}} = 1 - e^{-\tau}$$

Main source of opacity in KNe: bound-bound



$\kappa(\lambda)$ opacities from Tanaka+2020 as a function of ρ, T, Y_e and time



POSSIS

A 3D Monte Carlo radiative transfer code to model kilonovae

[**MB**+2015, MNRAS; **MB** 2019, MNRAS]

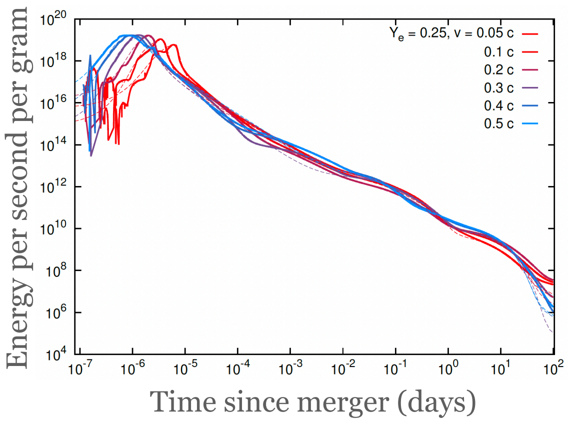
Creating photons

Frequency

From temperature + opacity

Energy

Nuclear heating rates
Thermalisation efficiencies



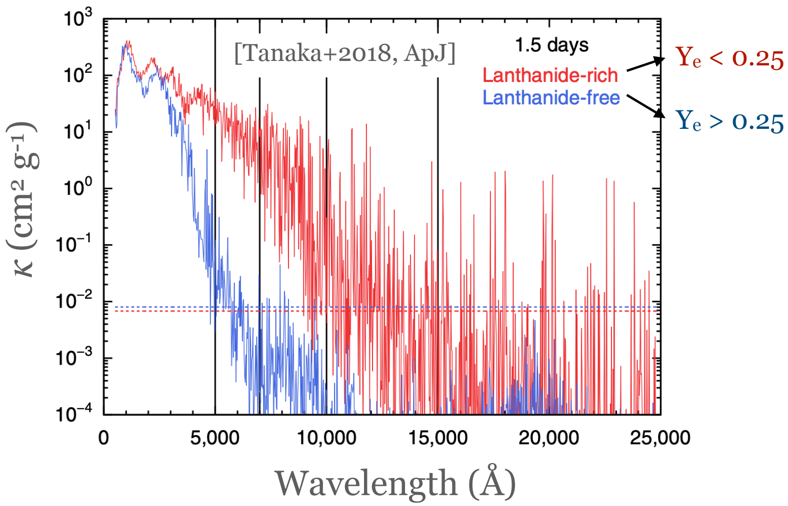
Stokes parameters

Propagating photons

Opacity

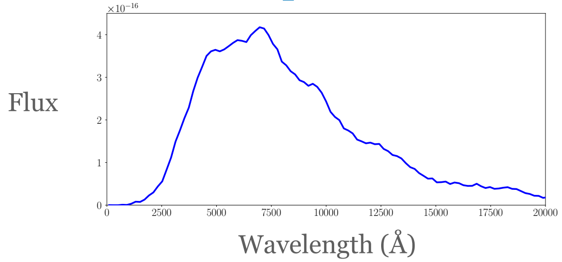
$$\tau = \int \kappa \rho \, dr \qquad P_{\mathrm{interaction}} = 1 - e^{-\tau}$$

Main source of opacity in KNe: bound-bound

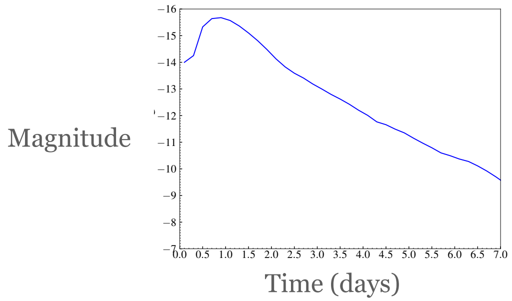


Collecting photons

Spectra



Light curves



Polarization

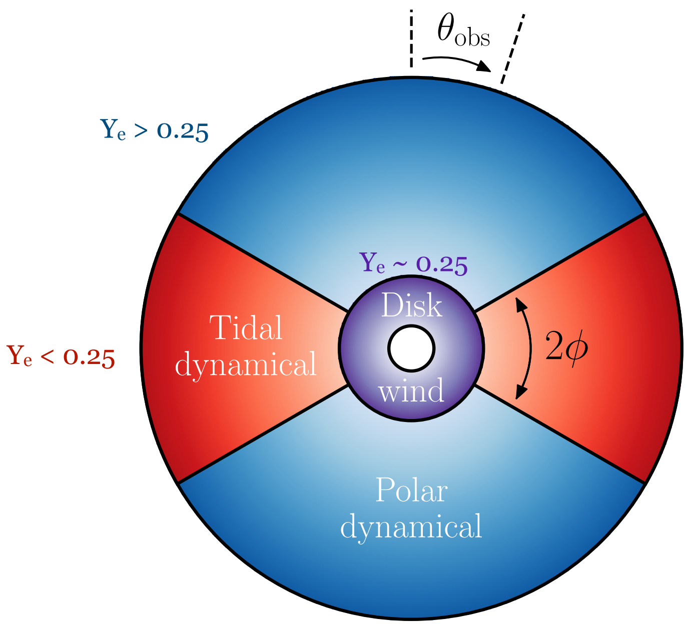
POSSIS

A 3D Monte Carlo radiative transfer code to model kilonovae

[**MB**+2015; **MB** 2019]

Neutron Star - Neutron Star

[Dietrich, Coughlin, Pang, **MB**+2020, Science]

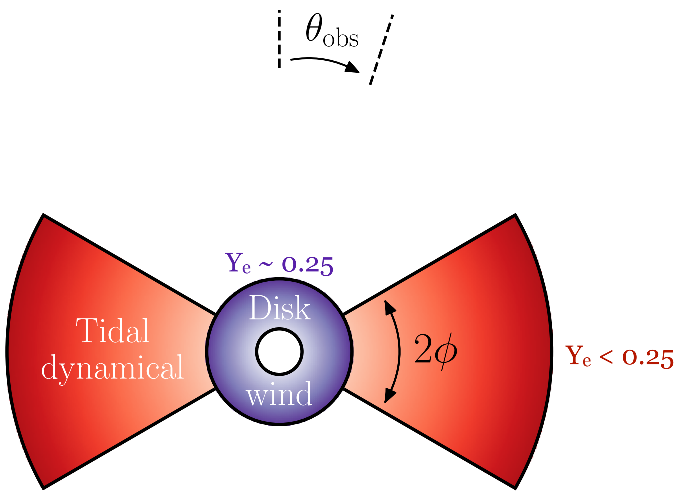


1540 models

varying ejecta masses ($M_{ej,dyn}$, $M_{ej,wind}$),
half-opening angles (ϕ) and viewing angle (θ_{obs})

Black Hole - Neutron Star

[Anand, Coughlin, Kasliwal, **MB**+2020, Nature Astronomy]



891 models

varying ejecta masses ($M_{ej,dyn}$, $M_{ej,wind}$),
and viewing angle (θ_{obs})

POSSIS

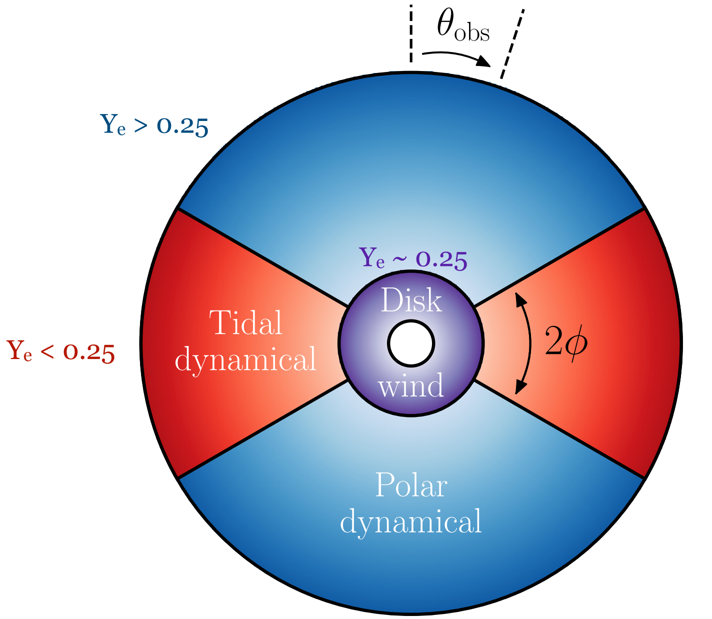
Help yourself! Modelled grids available at
https://github.com/mbulla/kilonova_models GitHub

A 3D Monte Carlo radiative transfer code to model kilonovae

[**MB**+2015; **MB** 2019]

Neutron Star - Neutron Star

[Dietrich, Coughlin, Pang, **MB**+2020, Science]

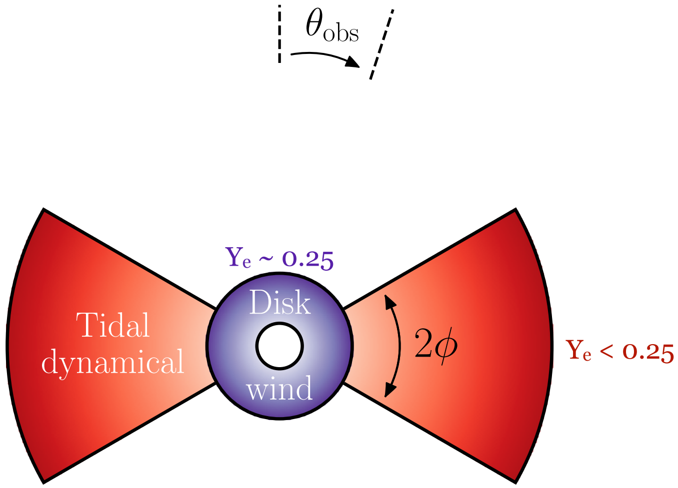


1540 models

varying ejecta masses ($M_{ej,dyn}$, $M_{ej,wind}$),
half-opening angles (ϕ) and viewing angle (θ_{obs})

Black Hole - Neutron Star

[Anand, Coughlin, Kasliwal, **MB**+2020, Nature Astronomy]



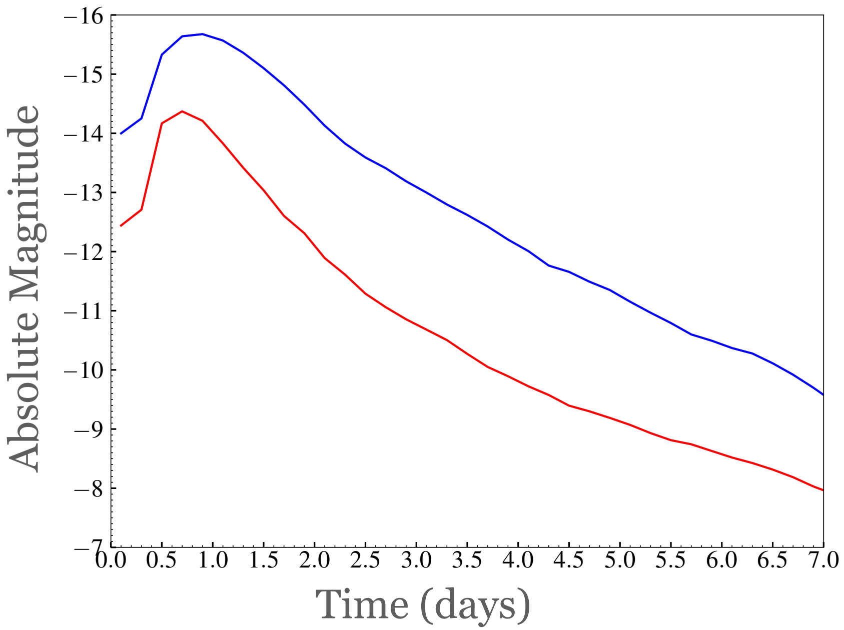
891 models

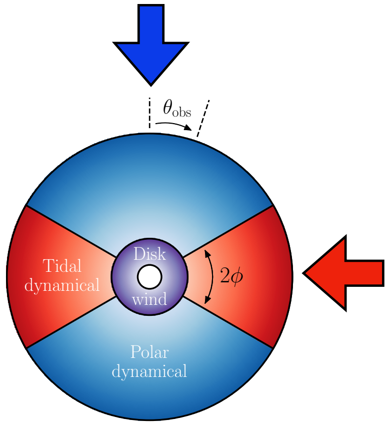
varying ejecta masses ($M_{ej,dyn}$, $M_{ej,wind}$),
and viewing angle (θ_{obs})

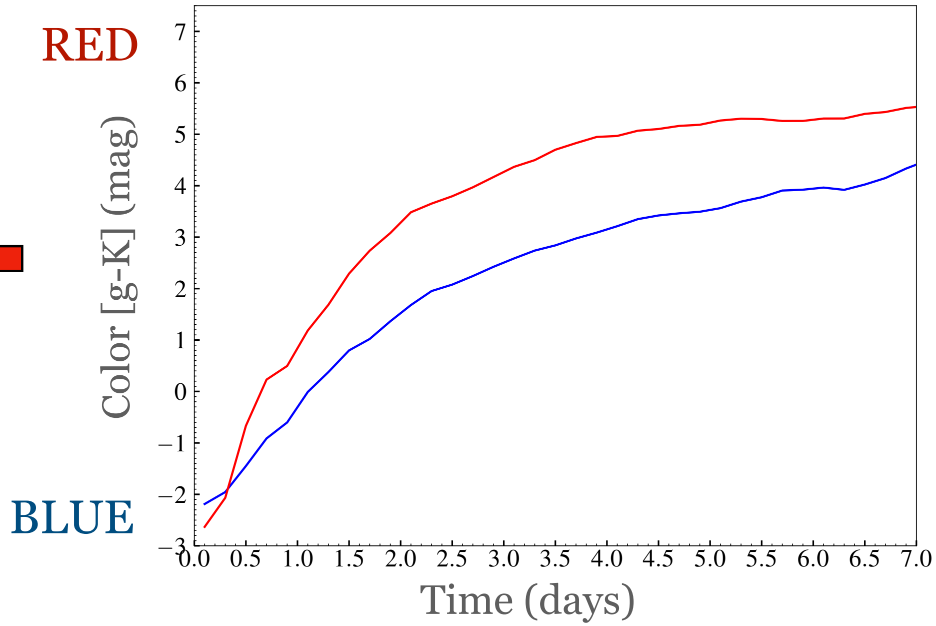
POSSIS

Viewing-angle dependence

Kilonovae viewed face-on ($\theta_{obs} = 0°$, jet axis) are brighter and bluer compared to kilonovae viewed edge-on ($\theta_{obs} = 90°$, merger plane)







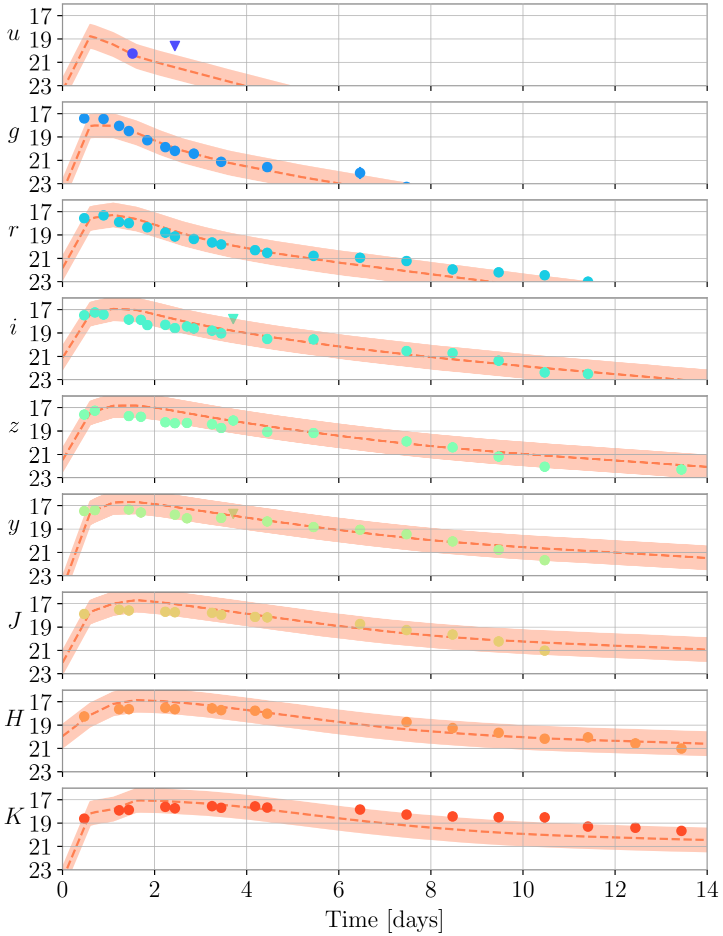
A fit to the kilonova AT2017gfo

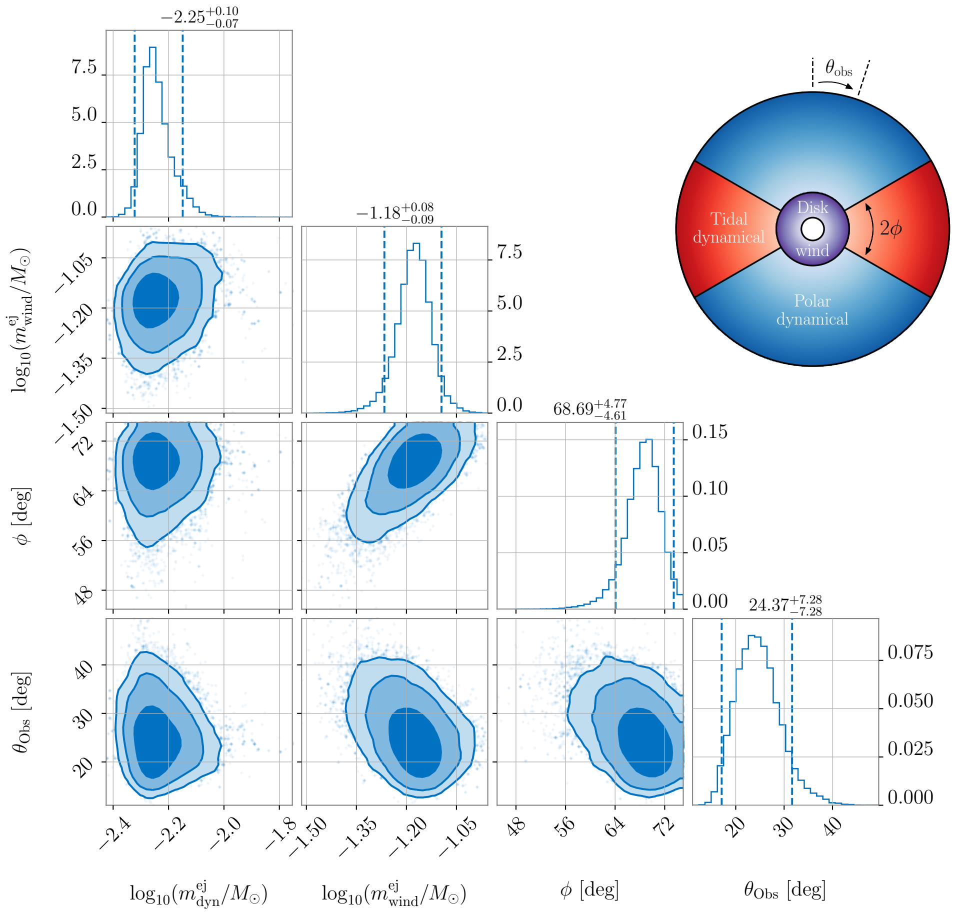
Interpolation scheme using Gaussian Process Regression or Neural Networks

[Pang, Dietrich, Coughlin, **MB**+, arXiv:2205.08513]

Gaussian Process Regression
[Coughlin..**MB**..+2020, PRR]

Neural Networks
[Almualla, Ning, **MB**+2021, arXiv:2112.15470]





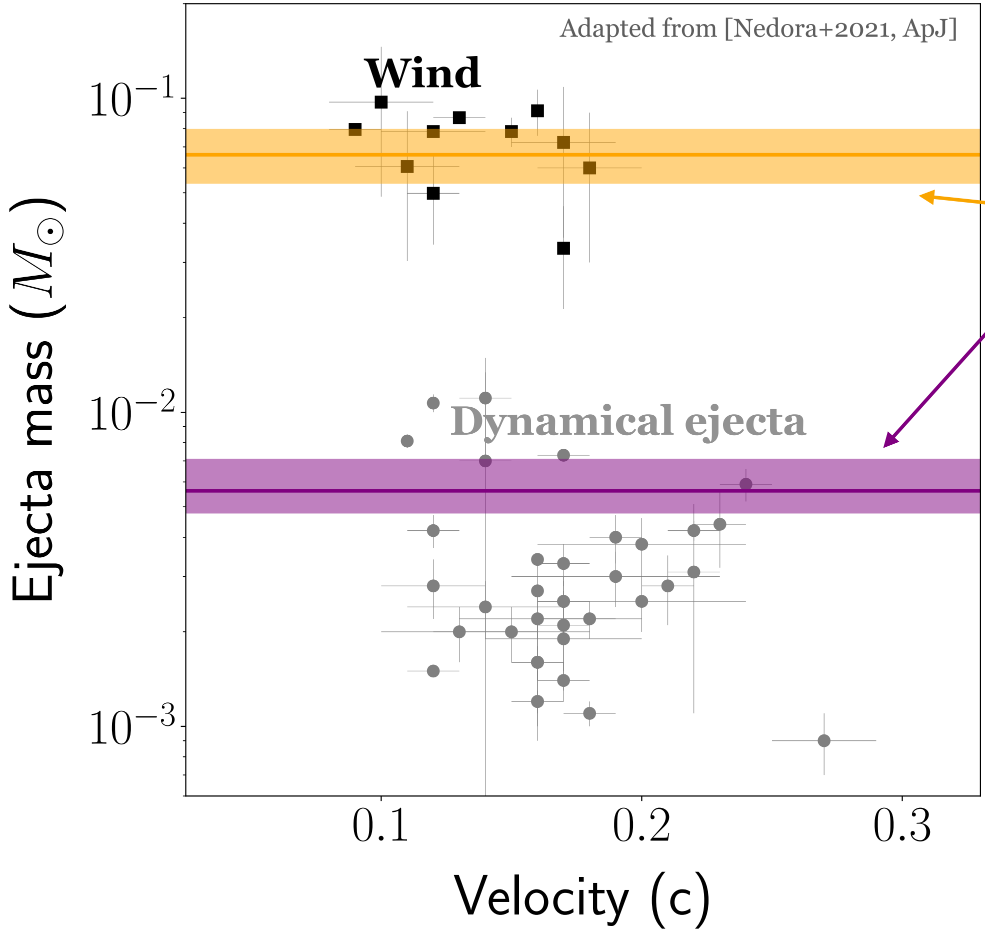
A fit to the kilonova AT2017gfo

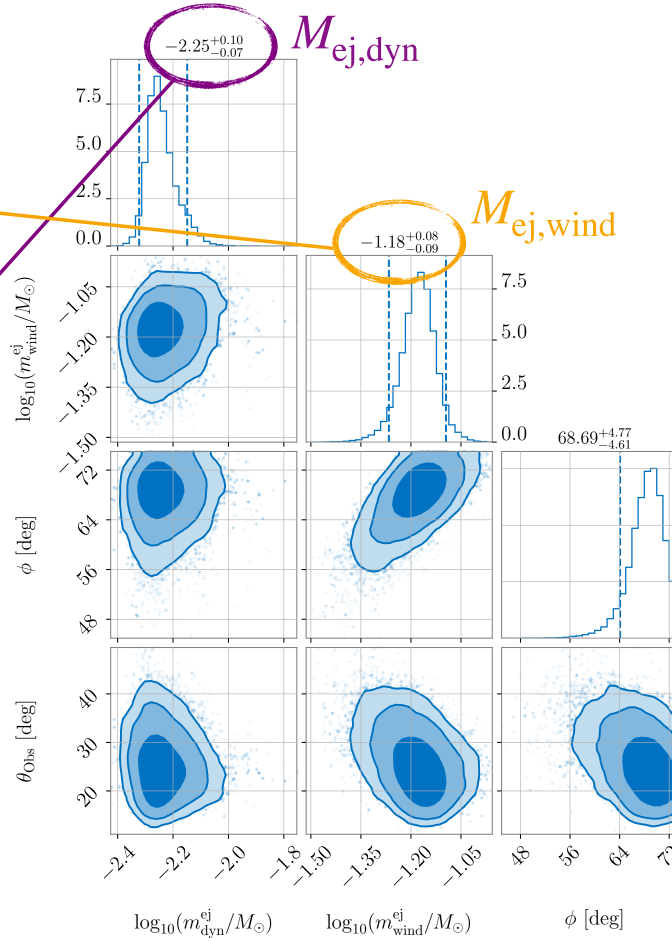
Interpolation scheme using Gaussian Process Regression or Neural Networks

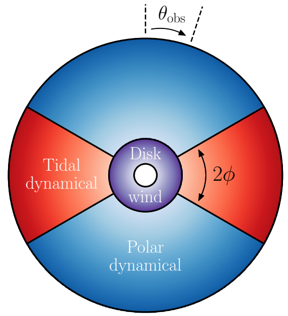
Gaussian Process Regression
[Coughlin..**MB**..+2020, PRR]

[Pang, Dietrich, Coughlin, **MB**+, arXiv:2205.08513]

Neural Networks
[Almualla, Ning, **MB**+2021, arXiv:2112.15470]







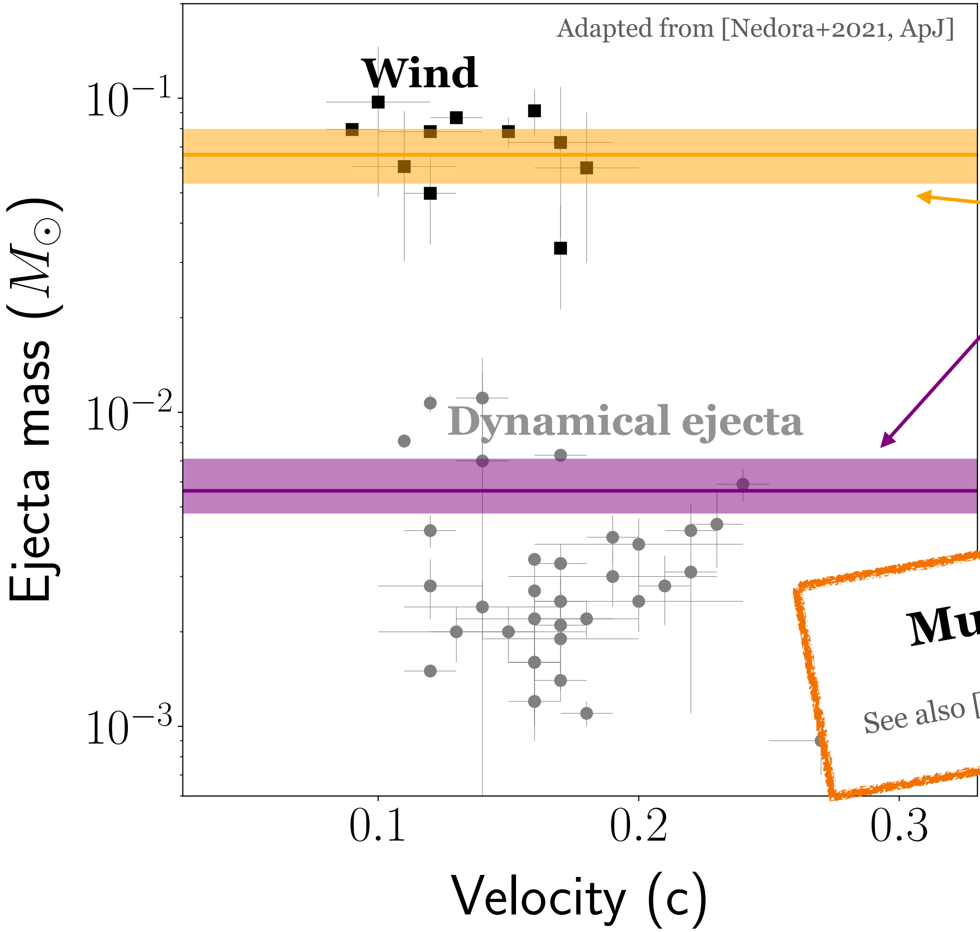
A fit to the kilonova AT2017gfo

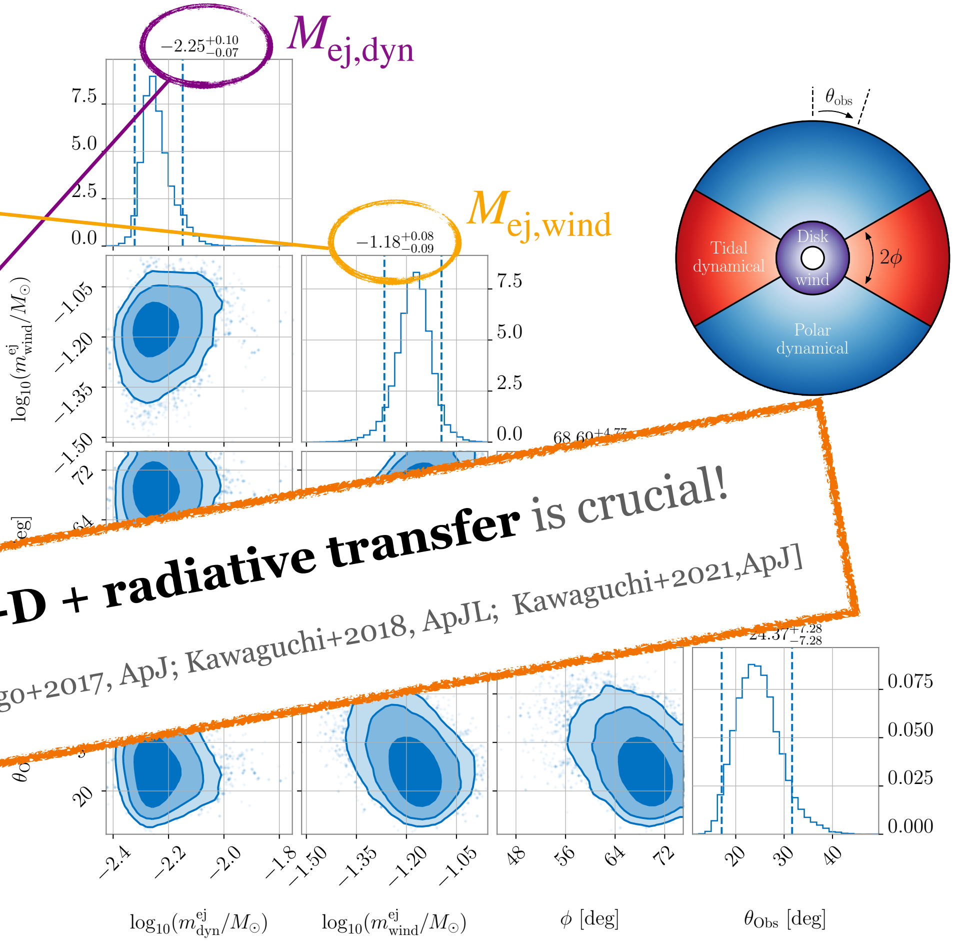
Interpolation scheme using Gaussian Process Regression or Neural Networks

[Pang, Dietrich, Coughlin, **MB**+, arXiv:2205.08513]

Gaussian Process Regression
[Coughlin..**MB**..+2020, PRR]

Neural Networks
[Almualla, Ning, **MB**+2021, arXiv:2112.15470]





Multi-D + radiative transfer is crucial!

See also [Perego+2017, ApJ; Kawaguchi+2018, ApJL; Kawaguchi+2021,ApJ]

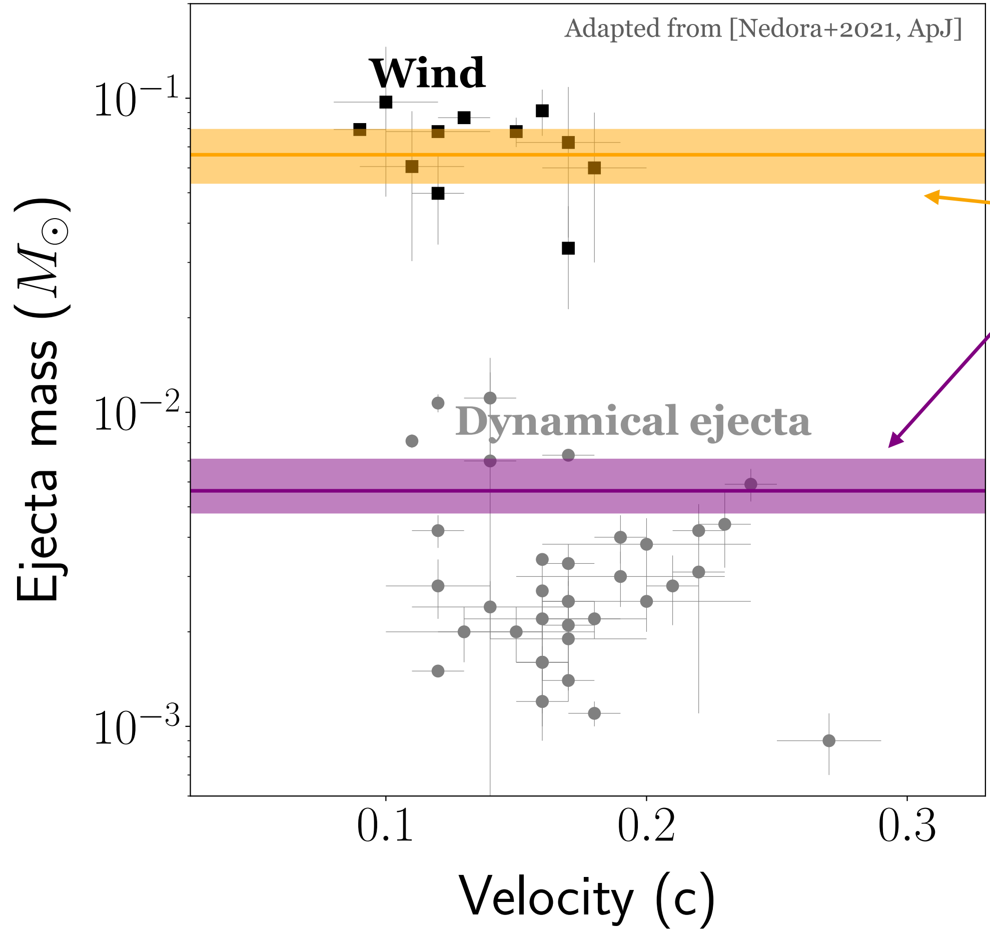
A fit to the kilonova AT2017gfo

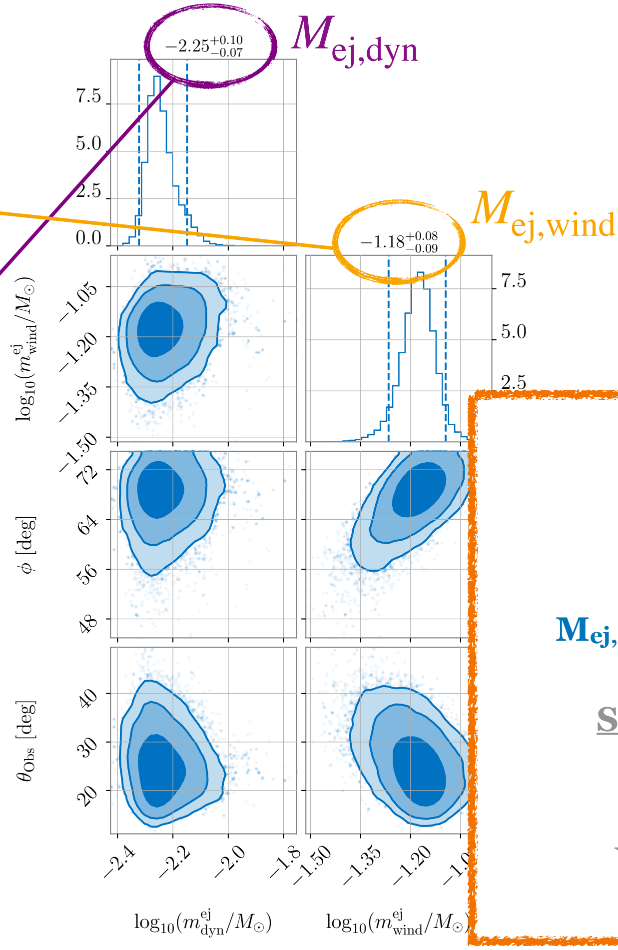
Interpolation scheme using Gaussian Process Regression or Neural Networks

[Pang, Dietrich, Coughlin, **MB**+, arXiv:2205.08513]

Gaussian Process Regression
[Coughlin..**MB**..+2020, PRR]

Neural Networks
[Almualla, Ning, **MB**+2021, arXiv:2112.15470]





[Anand, Mooley+, in prep]

6D GRID

Dynamical ejecta
$M_{ej,dyn}$, $<v_{ej,dyn}>$, $<Y_{e,dyn}>$

Secular/wind ejecta
$M_{ej,wind}$, $<v_{ej,wind}>$

Viewing angle θ_{obs}

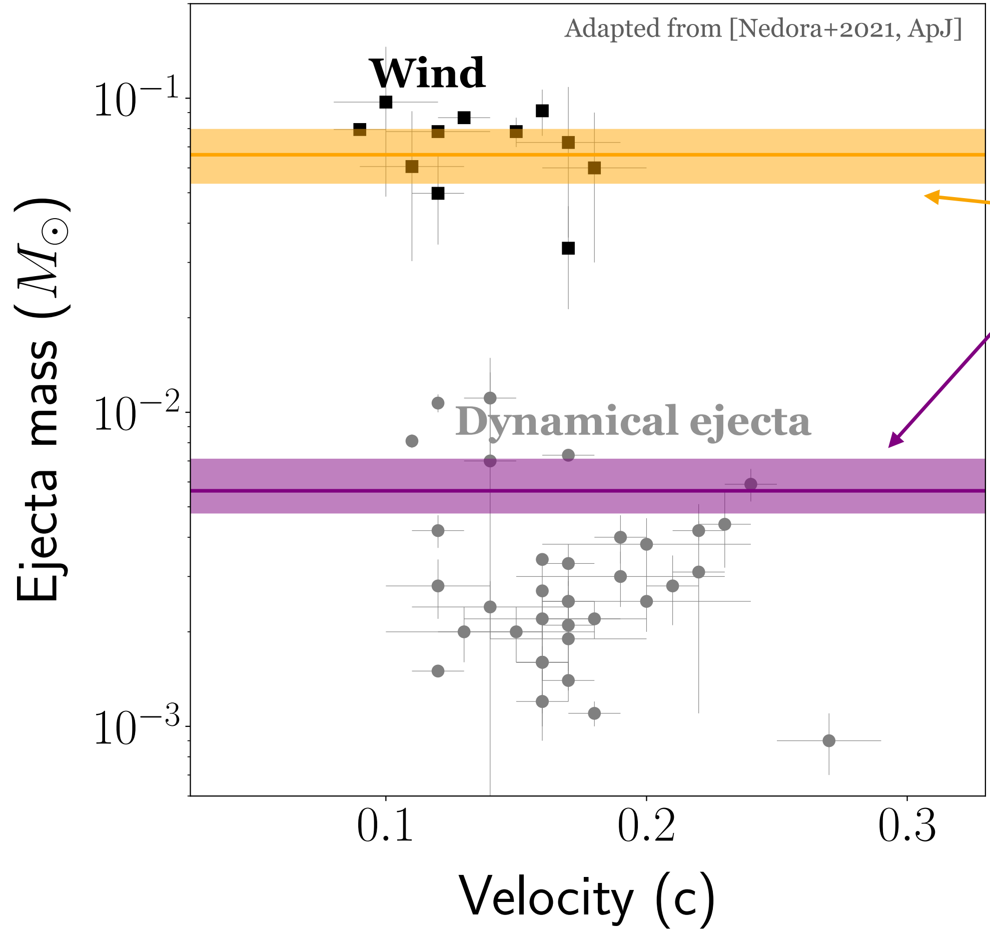
A fit to the kilonova AT2017gfo

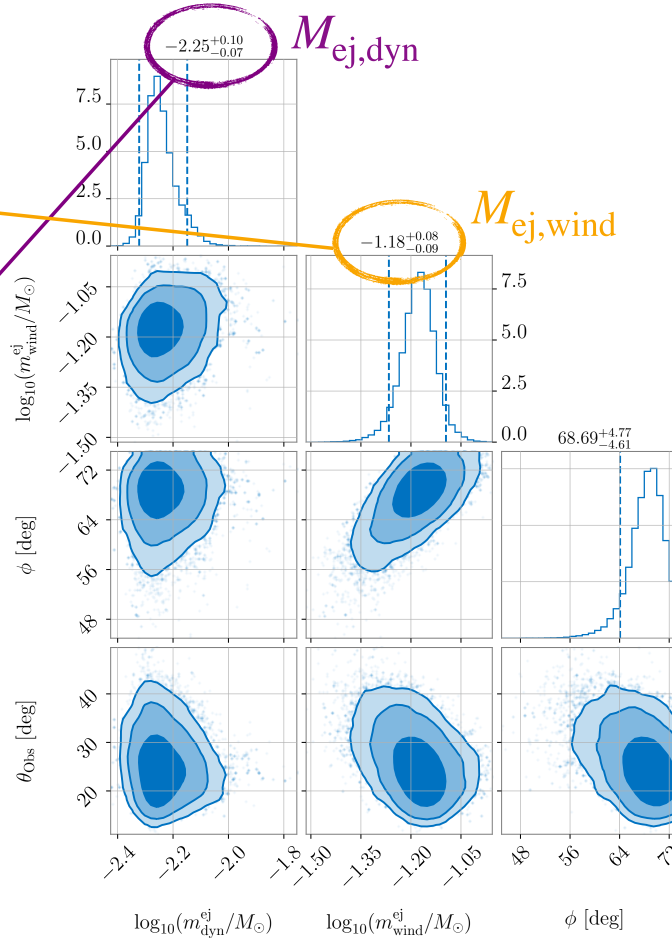
Gaussian Process Regression
[Coughlin..**MB**..+2020, PRR]

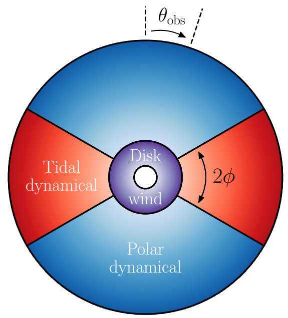
Interpolation scheme using Gaussian Process Regression or Neural Networks

[Pang, Dietrich, Coughlin, **MB**+, arXiv:2205.08513]

Neural Networks
[Almualla, Ning, **MB**+2021, arXiv:2112.15470]





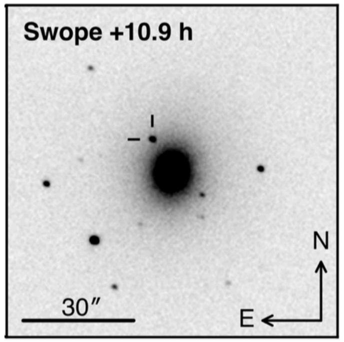


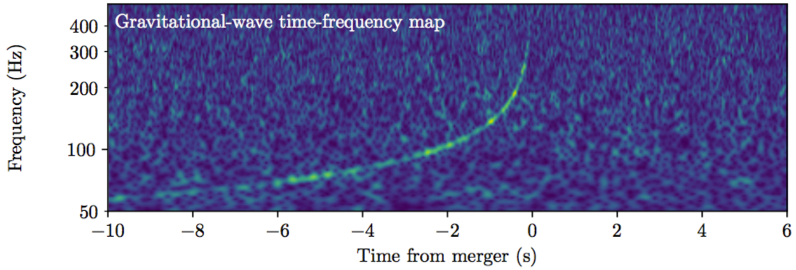
The Hubble constant H₀

Gravitational Waves as Standard Sirens

[Schutz 1986, Nature; Holz & Hughes 2005, ApJ]

$$H_0 = \frac{\text{Velocity}}{\text{Distance}} = \frac{[\text{speed of light}] \cdot \text{Redshift}}{\text{Distance}}$$



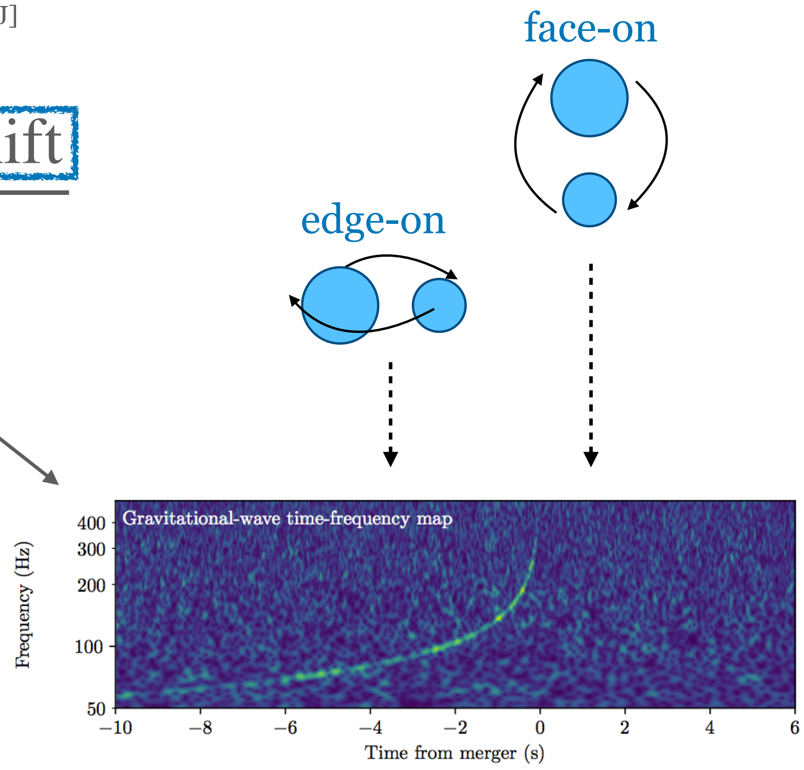


The Hubble constant H_0

Gravitational Waves as Standard Sirens

[Schutz 1986, Nature; Holz & Hughes 2005, ApJ]

$$H_0 = \frac{\text{Velocity}}{\text{Distance}} = \frac{[\text{speed of light}] \cdot \text{Redshift}}{\text{Distance}}$$

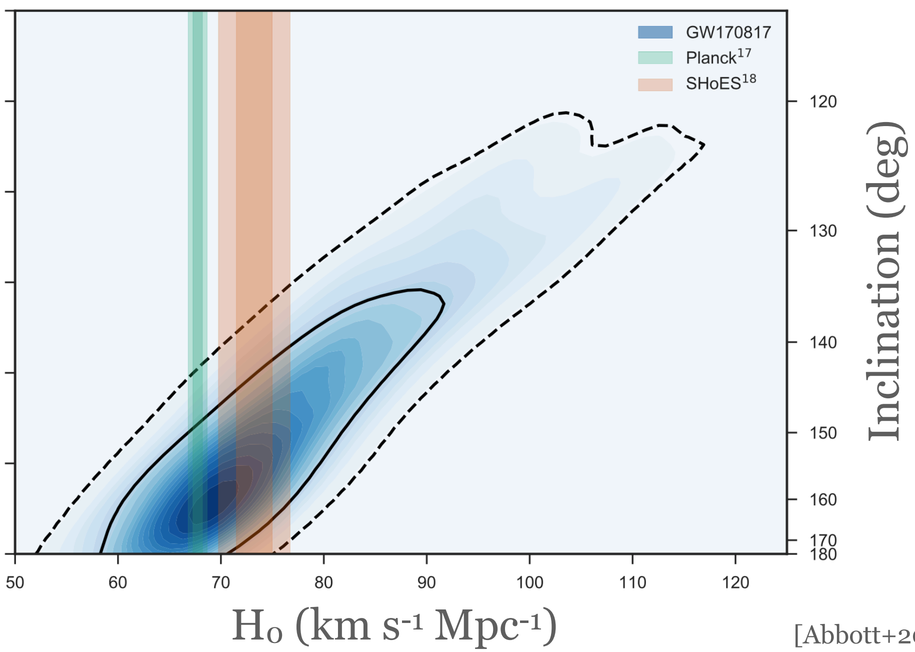


The Hubble constant H_0

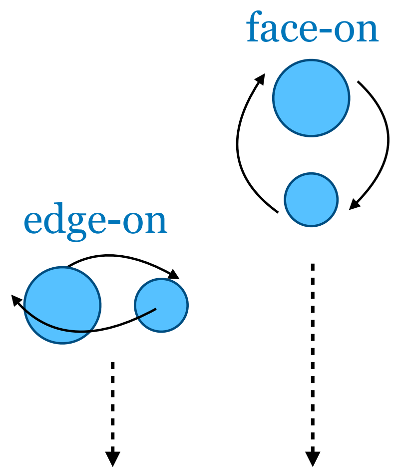
Gravitational Waves as Standard Sirens

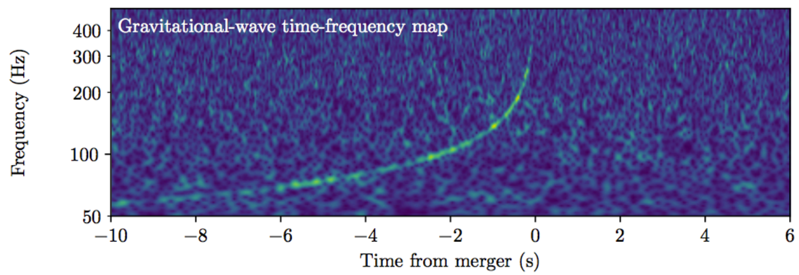
[Schutz 1986, Nature; Holz & Hughes 2005, ApJ]

$$H_0 = \frac{\text{Velocity}}{\text{Distance}} = \frac{[\text{speed of light}] \cdot \boxed{\text{Redshift}}}{\boxed{\text{Distance}}}$$



[Abbott+2017, Nature]





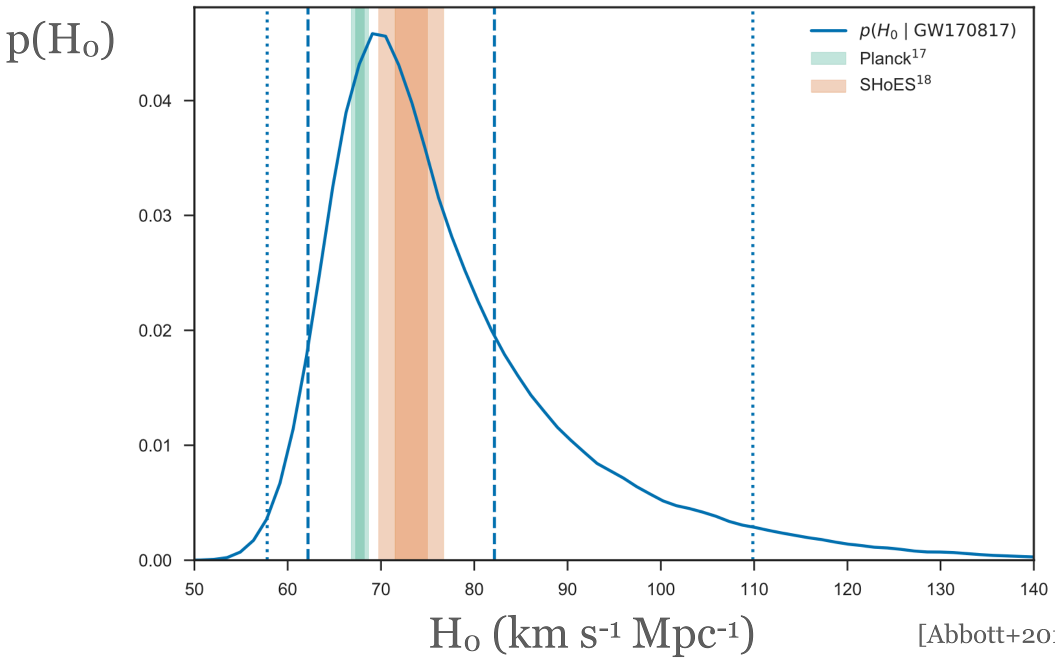
$$H_0 = 70.0^{+12.0}_{-8.0}\,\mathrm{km\,s^{-1}\,Mpc^{-1}}$$

The Hubble constant H_0

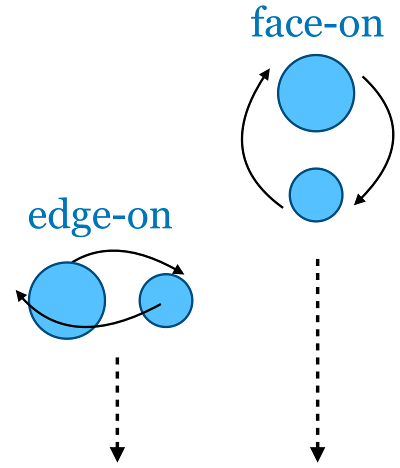
Gravitational Waves as Standard Sirens

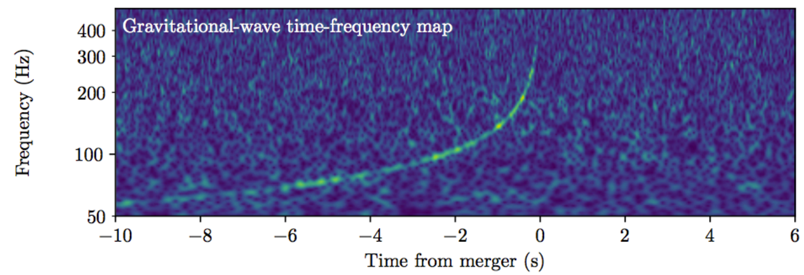
[Schutz 1986, Nature; Holz & Hughes 2005, ApJ]

$$H_0 = \frac{\text{Velocity}}{\text{Distance}} = \frac{[\text{speed of light}] \cdot \text{Redshift}}{\text{Distance}}$$



[Abbott+2017, Nature]





$$H_0 = 70.0^{+12.0}_{-8.0}\,\mathrm{km\,s^{-1}\,Mpc^{-1}}$$

The Hubble constant H_0

24% improvement on H_0

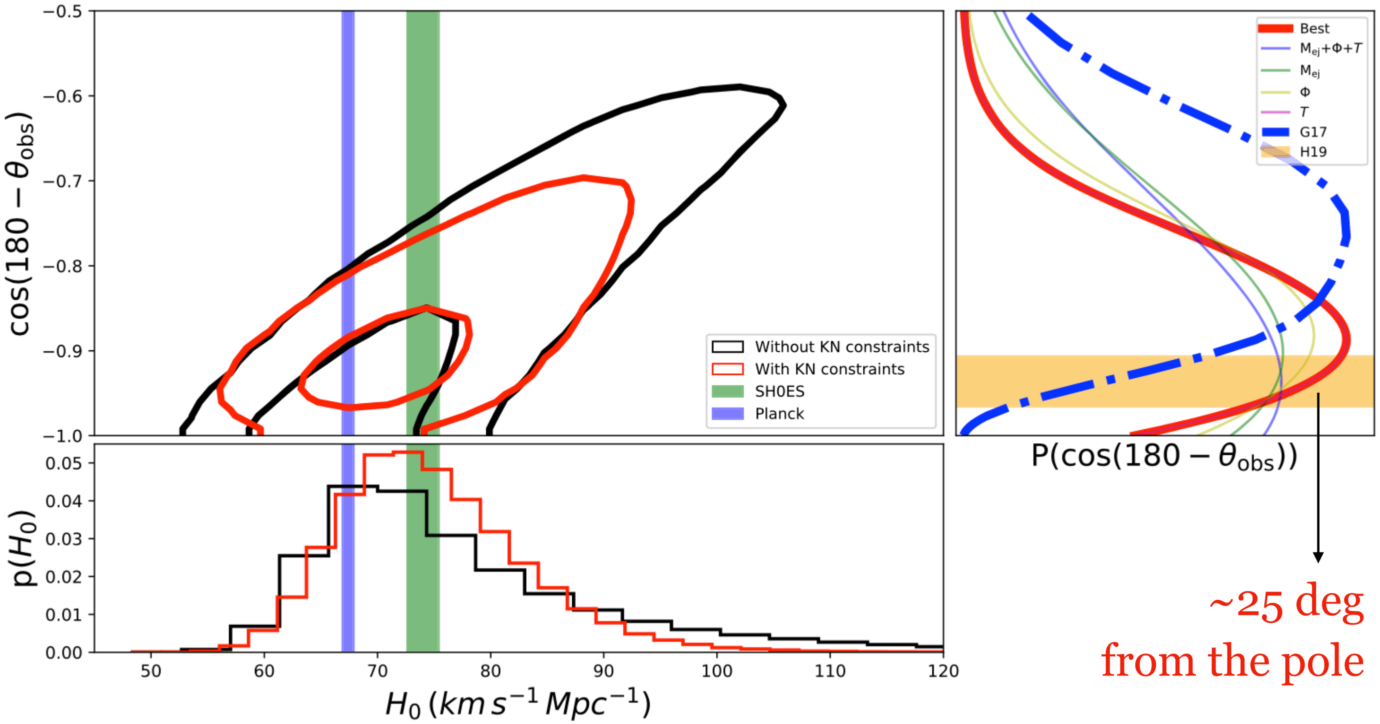
[Dhawan, **MB**+2020, ApJ]

GW only

$$H_0 = 70.0^{+12.0}_{-8.0}\,\mathrm{km\,s^{-1}\,Mpc^{-1}}$$

GW + kilonova

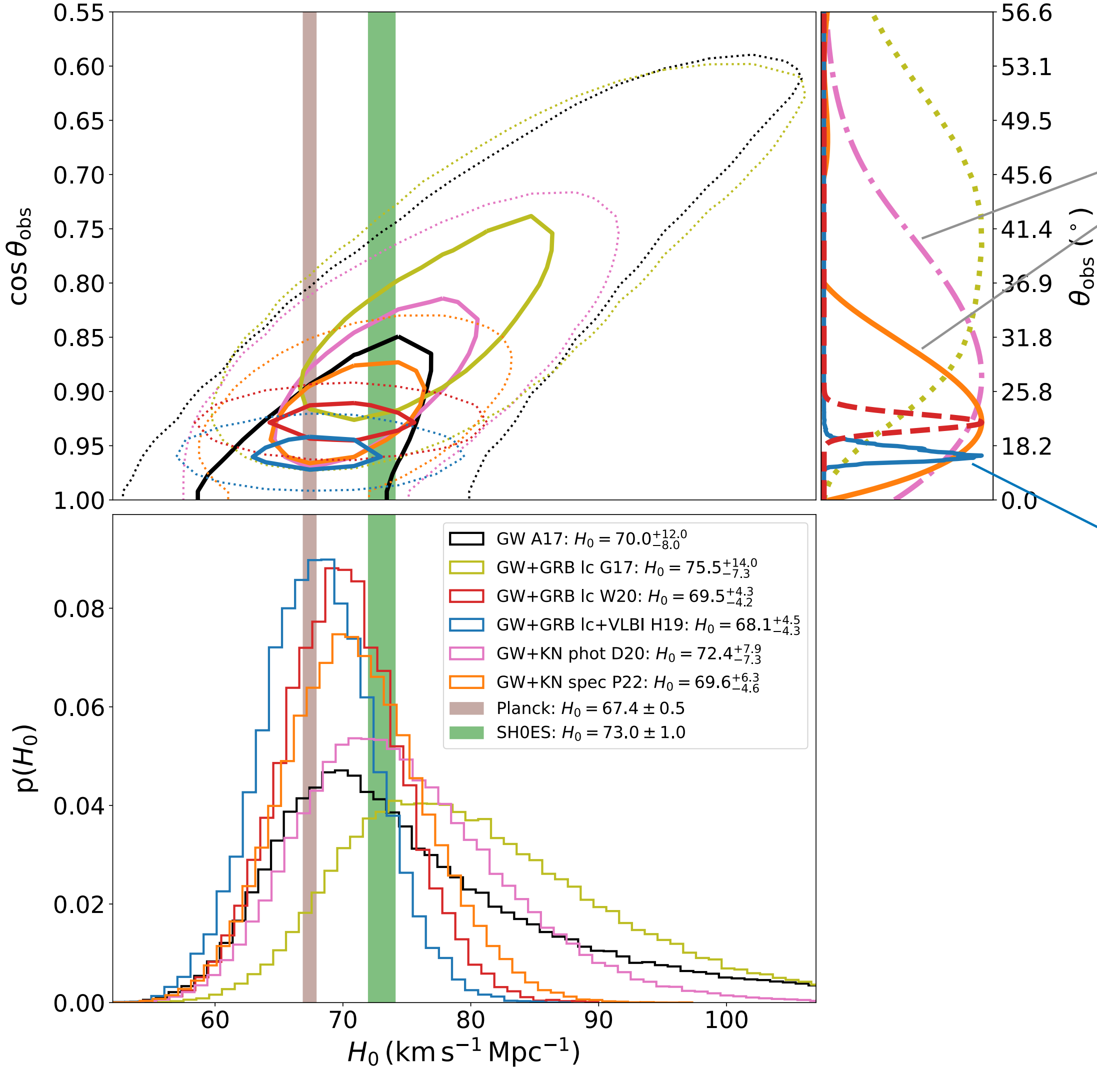
$$H_0 = 72.4^{+7.9}_{-7.3}\,\mathrm{km\,s^{-1}\,Mpc^{-1}}$$



~25 deg
from the pole

see also [Coughlin...**MB**...+2020, Nature Communications]

The Hubble constant H_0

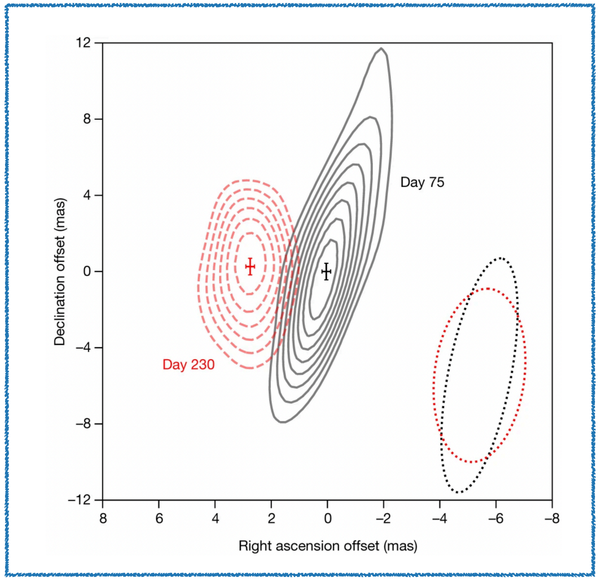


[**MB**, Coughlin, Dhawan & Dietrich 2022, Universe]

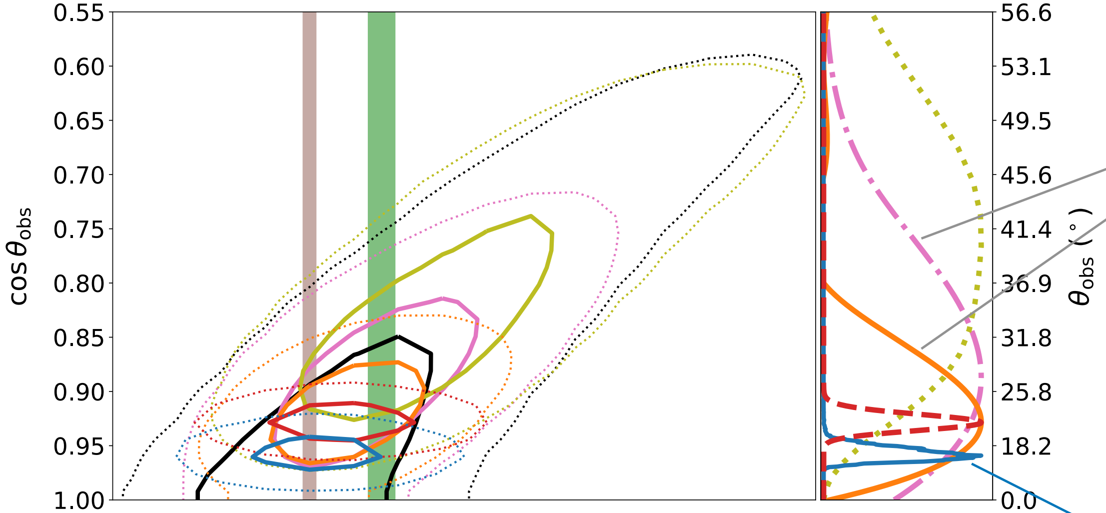
Kilonova

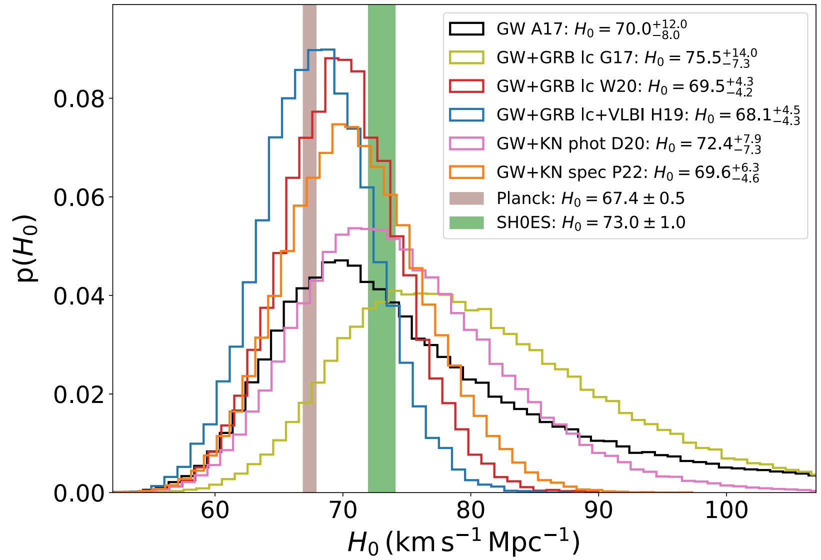
Superluminal motion

[Mooley+2018, Nature]
[Hotokezaka+2019, NatAstro]



The Hubble constant H_0



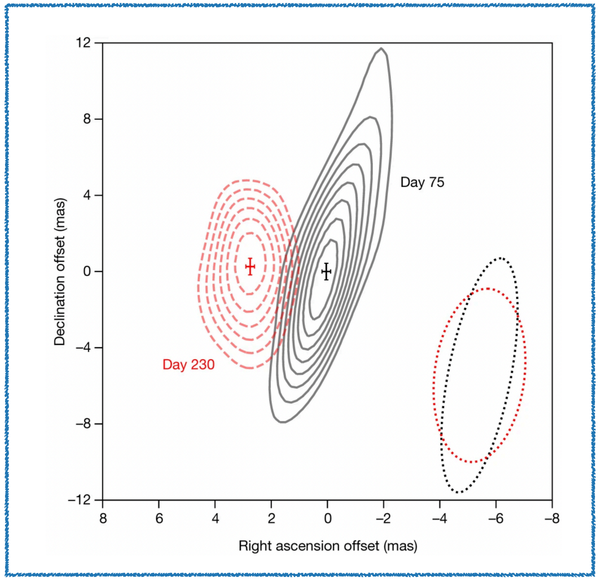


[**MB**, Coughlin, Dhawan & Dietrich 2022, Universe]

Kilonova

Superluminal motion

[Mooley+2018, Nature]
[Hotokezaka+2019, NatAstro]



Systematics?

[Heinzel, Coughlin, Dietrich, **MB**, Antier+ 2021, MNRAS]

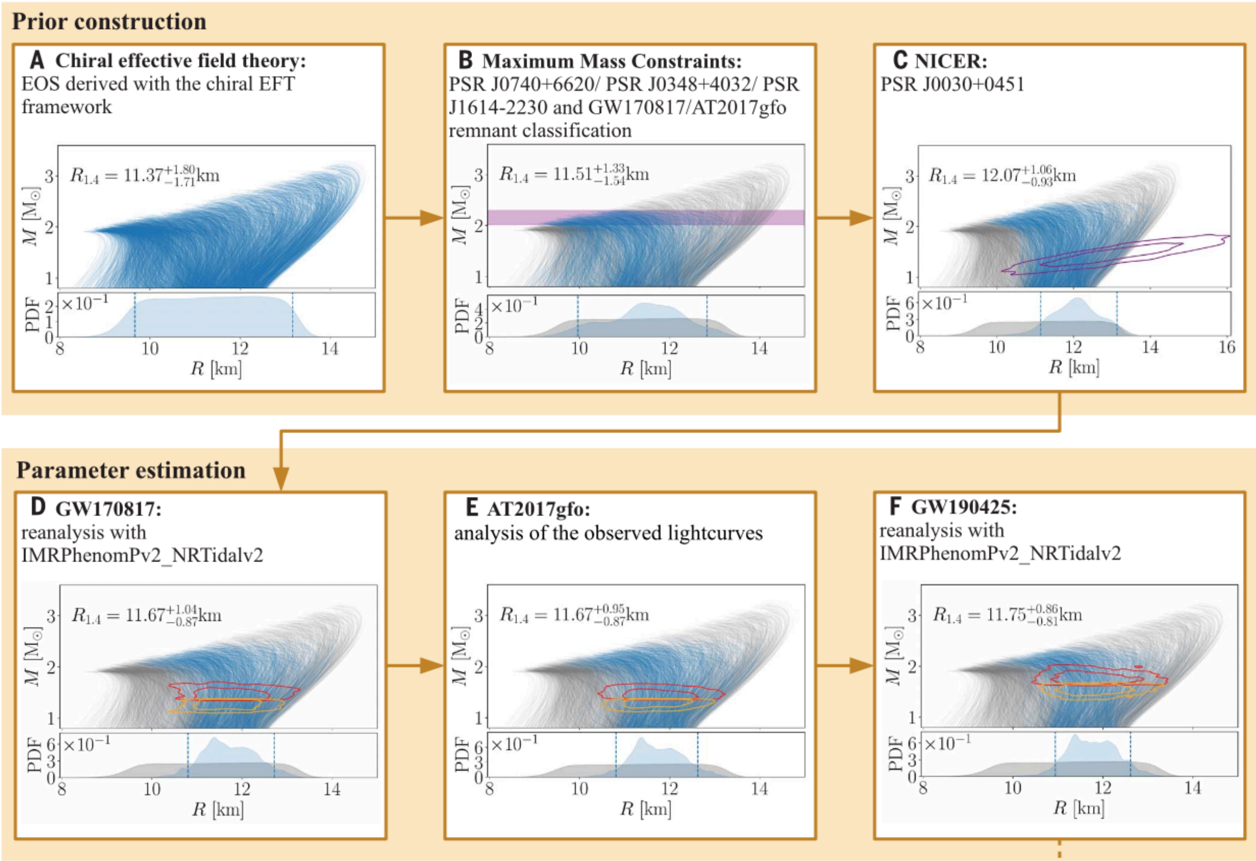
NMMA: A framework to rule them all

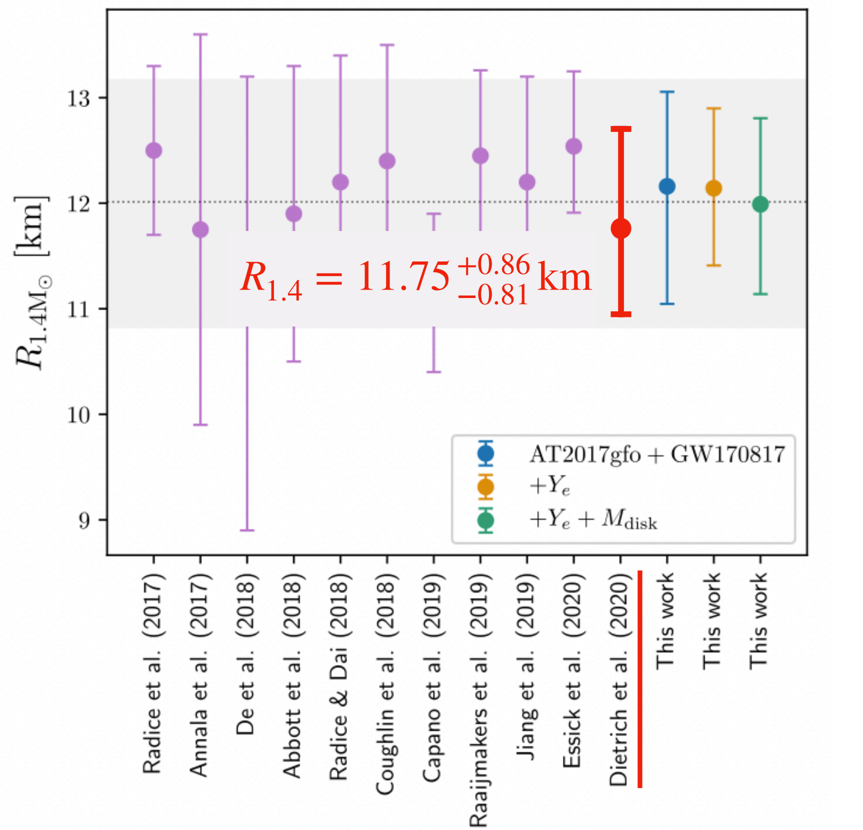
A nuclear-physics multi-messenger bayesian framework

[Dietrich, Coughlin, Pang, **MB**+2020, Science]

GitHub

https://nuclear-multimessenger-astronomy.github.io/nmma/





[Breschi+2021,MNRAS]

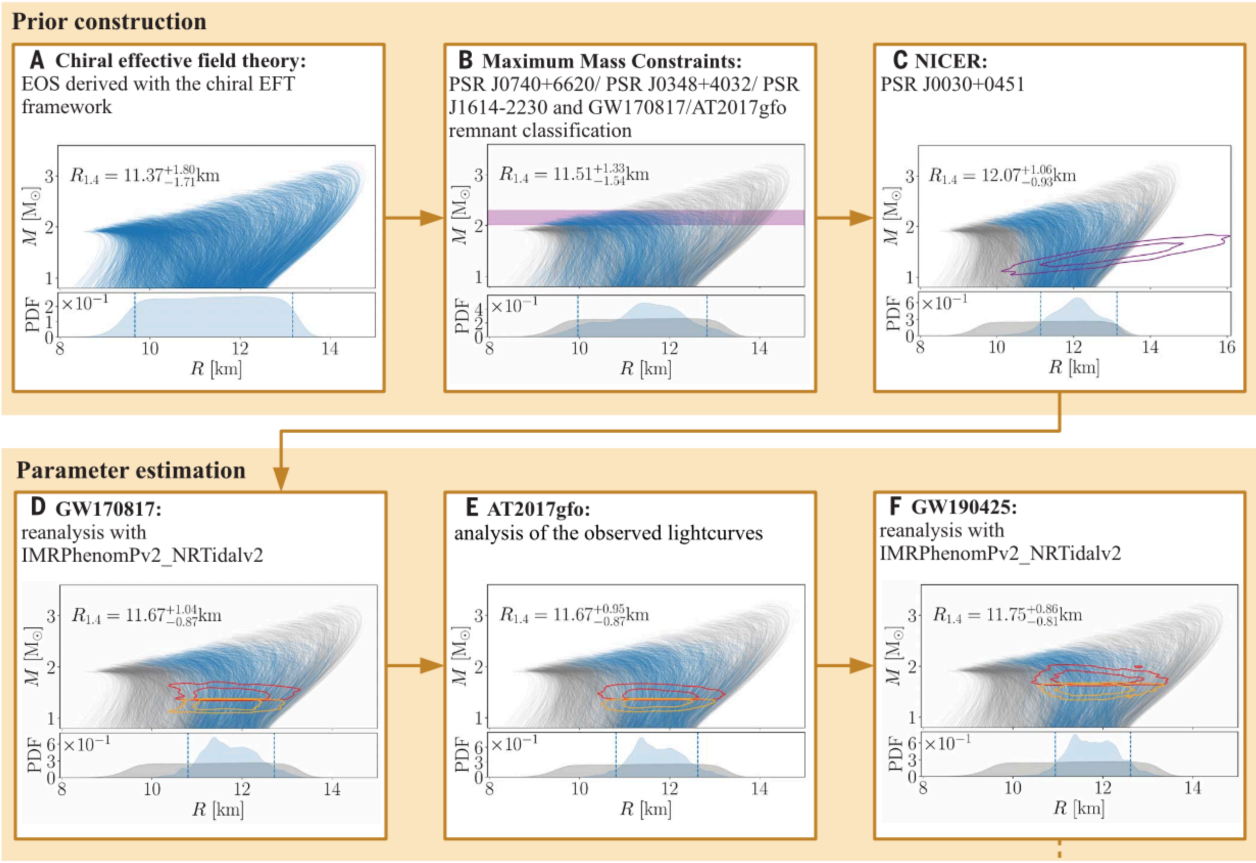
NMMA: A framework to rule them all

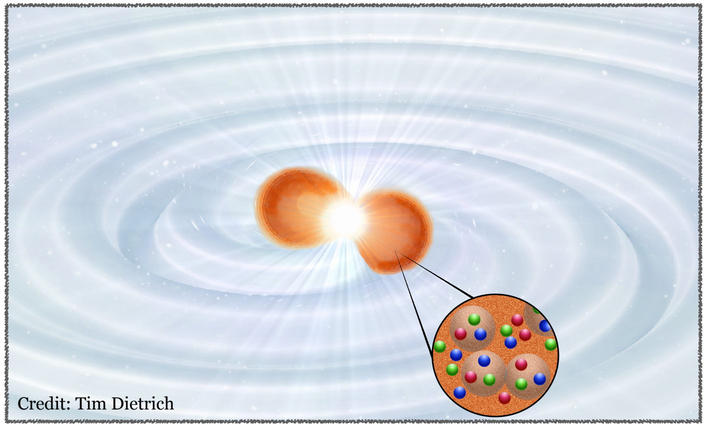
A nuclear-physics multi-messenger bayesian framework

[Dietrich, Coughlin, Pang, **MB**+2020, Science]

GitHub

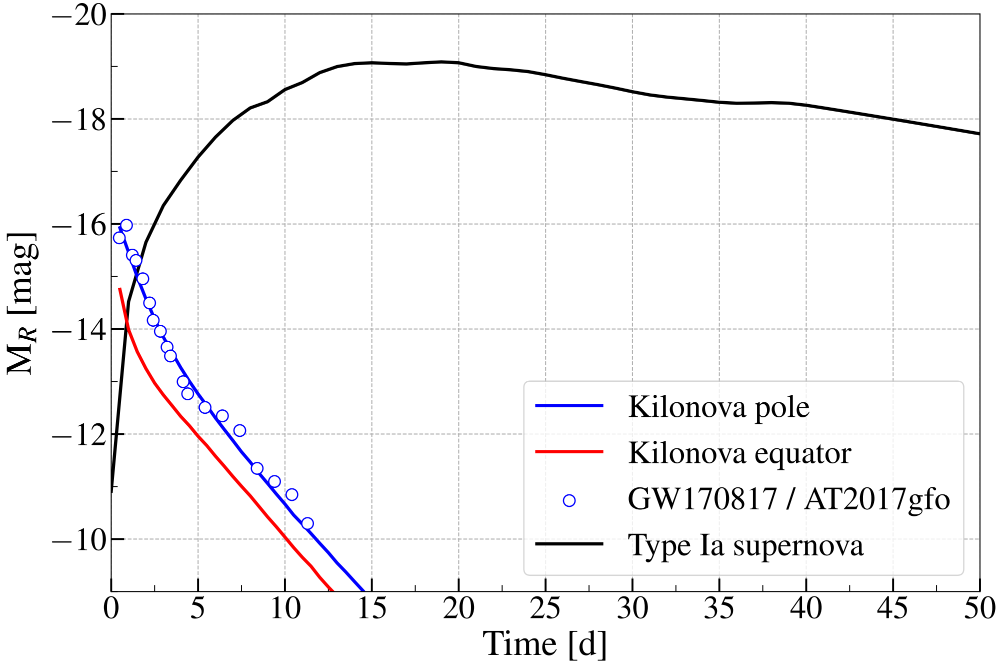
https://nuclear-multimessenger-astronomy.github.io/nmma/





- GW190814 as a BBH [Tews..**MB**..+2021, ApJL]
- Adding PSRJ0740+6620 [Pang, Tews, Coughlin, **MB**+2021, ApJ]
- Kilonova searches [Andreoni...**MB**...+,2021,ApJ]
- MM observations + HIC [Huth...**MB**...+, Nature]
- GRB211211A, in prep.

Hunting for kilonovae in O3

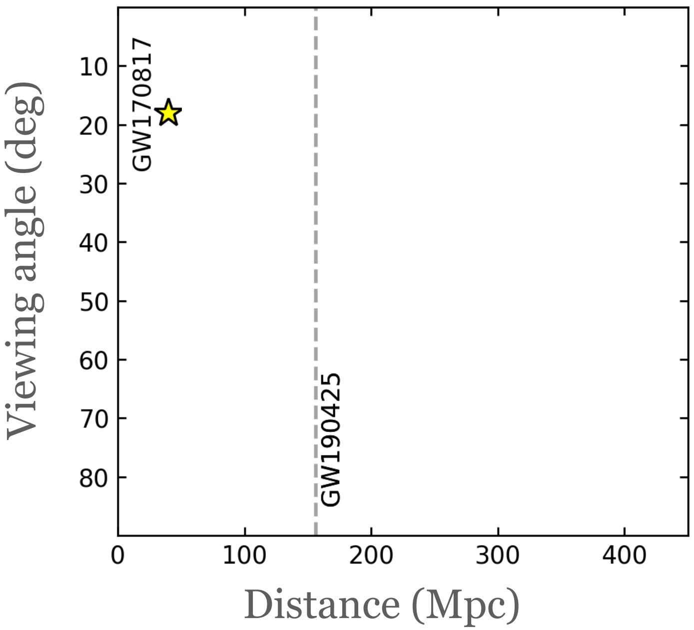


Name	Localization	Distance	Class
GW190425	7461 deg^2	156 ± 41 Mpc	BNS
S190426c	1131 deg^2	377 ± 100 Mpc	NSBH
GW190814	23 deg^2	267 ± 52 Mpc	NSBH
S190901ap	14,753 deg^2	241 ± 79 Mpc	BNS
S190910d	2482 deg^2	632 ± 186 Mpc	NSBH
S190910h	24,264 deg^2	230 ± 88 Mpc	BNS
S190923y	2107 deg^2	438 ± 133 Mpc	NSBH
S190930t	24,220 deg^2	108 ± 38 Mpc	NSBH
S191205ah	6378 deg^2	385 ± 164 Mpc	NSBH
S191213g	4480 deg^2	201 ± 81 Mpc	BNS
S200105ae	7373 deg^2	283 ± 74 Mpc	NSBH
S200115j	765 deg^2	340 ± 79 Mpc	NSBH
S200213t	2326 deg^2	201 ± 80 Mpc	BNS
GW170817	30 deg^2	40 Mpc	BNS

[Kasliwal..**MB**..,+2020, ApJ]

Detectability of kilonovae

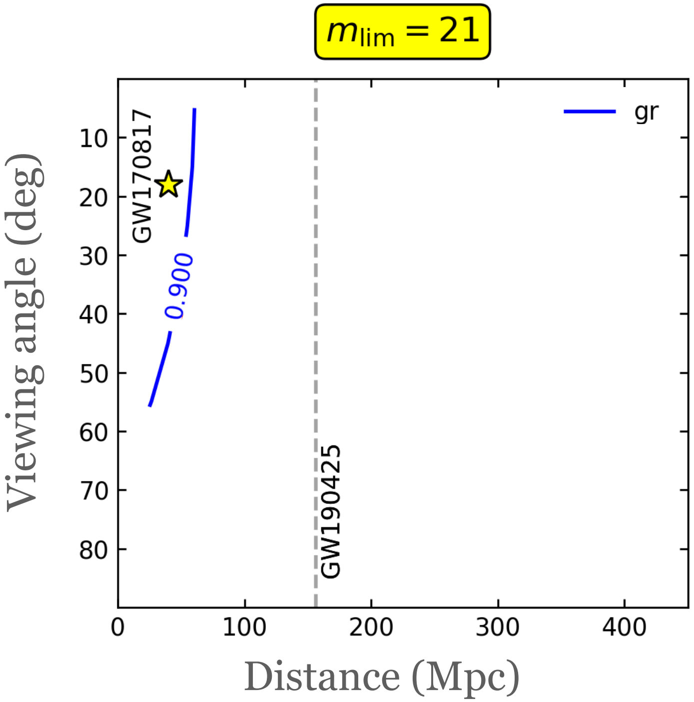
[Sagues-Carracedo, **MB**, Feindt & Goobar 2021, MNRAS]



see also [Coughlin, Dietrich, Antier, **MB**+2020a, MNRAS / Coughlin..**MB**..+2020b, MNRAS / Almualla..**MB**..+2021,MNRAS]

Detectability of kilonovae

[Sagues-Carracedo, **MB**, Feindt & Goobar 2021, MNRAS]

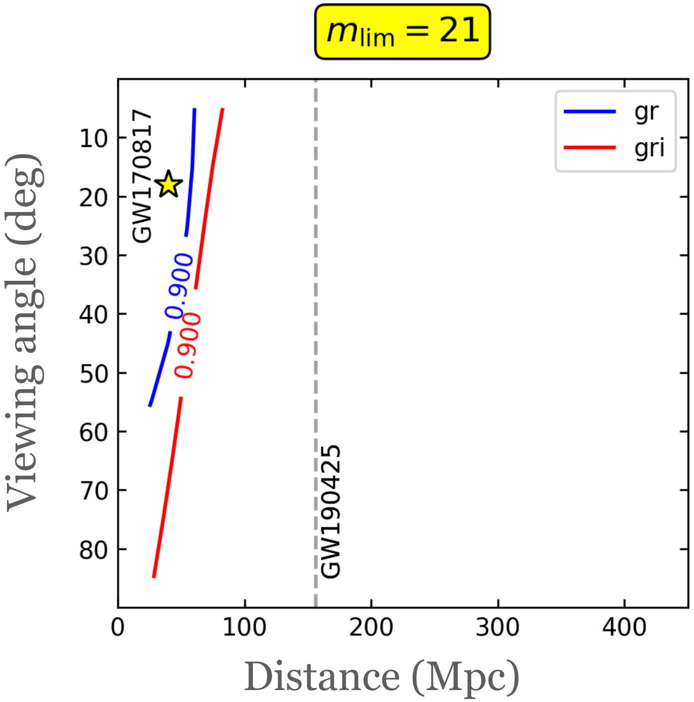


see also [Coughlin, Dietrich, Antier, **MB**+2020a, MNRAS / Coughlin..**MB**..+2020b, MNRAS / Almualla..**MB**..+2021,MNRAS]

Detectability of kilonovae

Go red!

[Sagues-Carracedo, **MB**, Feindt & Goobar 2021, MNRAS]

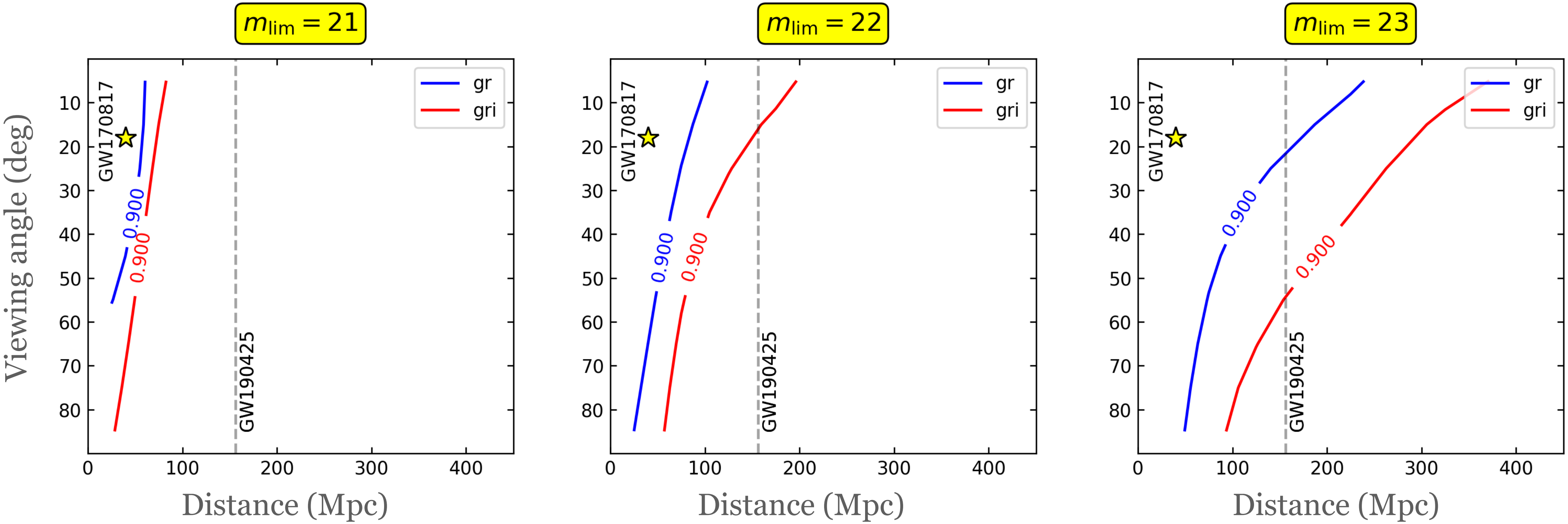


see also [Coughlin, Dietrich, Antier, **MB**+2020a, MNRAS / Coughlin..**MB**..+2020b, MNRAS / Almualla..**MB**..+2021,MNRAS]

Detectability of kilonovae

Go red! Go deep!

[Sagues-Carracedo, **MB**, Feindt & Goobar 2021, MNRAS]

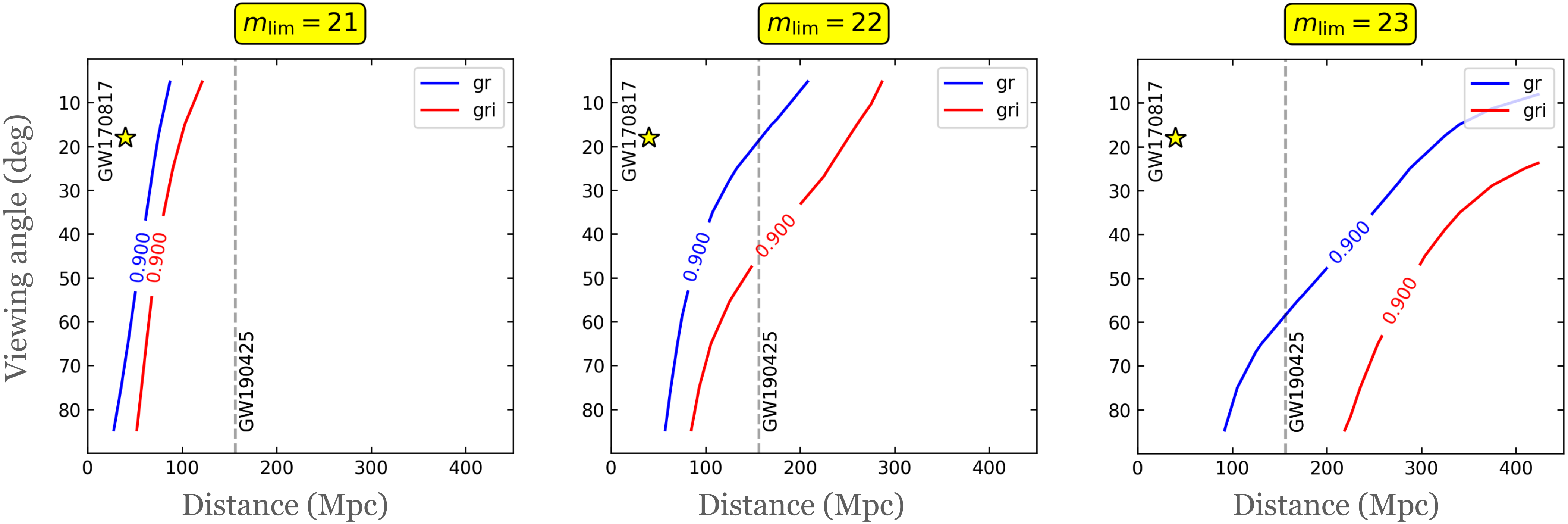


see also [Coughlin, Dietrich, Antier, **MB**+2020a, MNRAS / Coughlin..**MB**..+2020b, MNRAS / Almualla..**MB**..+2021,MNRAS]

Detectability of kilonovae

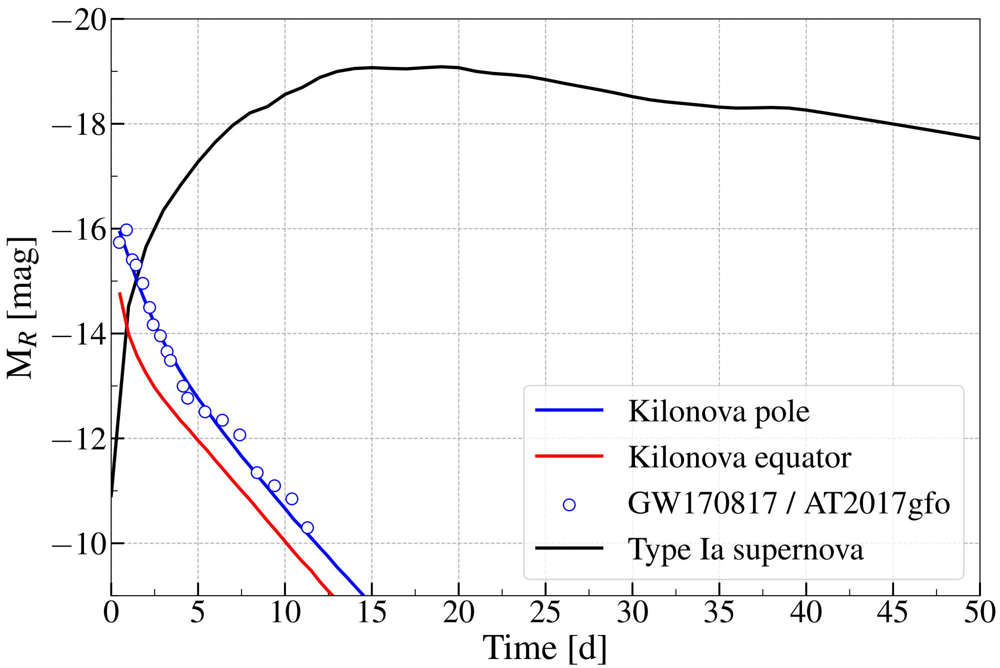
Go red! Go deep! Be quick!

[Sagues-Carracedo, **MB**, Feindt & Goobar 2021, MNRAS]



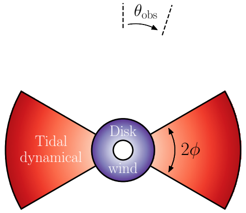
see also [Coughlin, Dietrich, Antier, **MB**+2020a, MNRAS / Coughlin..**MB**..+2020b, MNRAS / Almualla..**MB**..+2021,MNRAS]

Hunting for kilonovae in O3



Name	Localization	Distance	Class
GW190425	7461 deg^2	156 ± 41 Mpc	BNS
S190426c	1131 deg^2	377 ± 100 Mpc	NSBH
GW190814	23 deg^2	267 ± 52 Mpc	NSBH
S190901ap	14,753 deg^2	241 ± 79 Mpc	BNS
S190910d	2482 deg^2	632 ± 186 Mpc	NSBH
S190910h	24,264 deg^2	230 ± 88 Mpc	BNS
S190923y	2107 deg^2	438 ± 133 Mpc	NSBH
S190930t	24,220 deg^2	108 ± 38 Mpc	NSBH
S191205ah	6378 deg^2	385 ± 164 Mpc	NSBH
S191213g	4480 deg^2	201 ± 81 Mpc	BNS
S200105ae	7373 deg^2	283 ± 74 Mpc	NSBH
S200115j	765 deg^2	340 ± 79 Mpc	NSBH
S200213t	2326 deg^2	201 ± 80 Mpc	BNS
GW170817	30 deg^2	40 Mpc	BNS

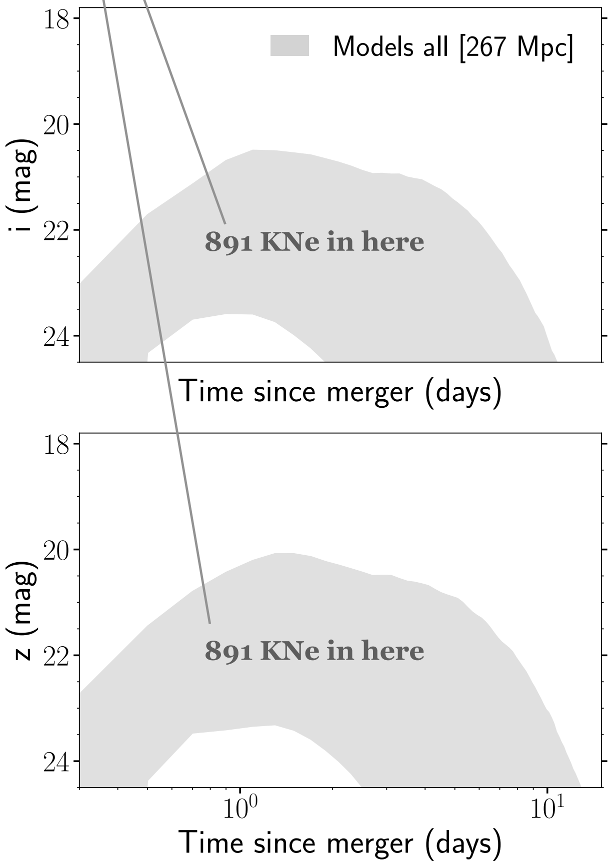
[Kasliwal..**MB**..,+2020, ApJ]

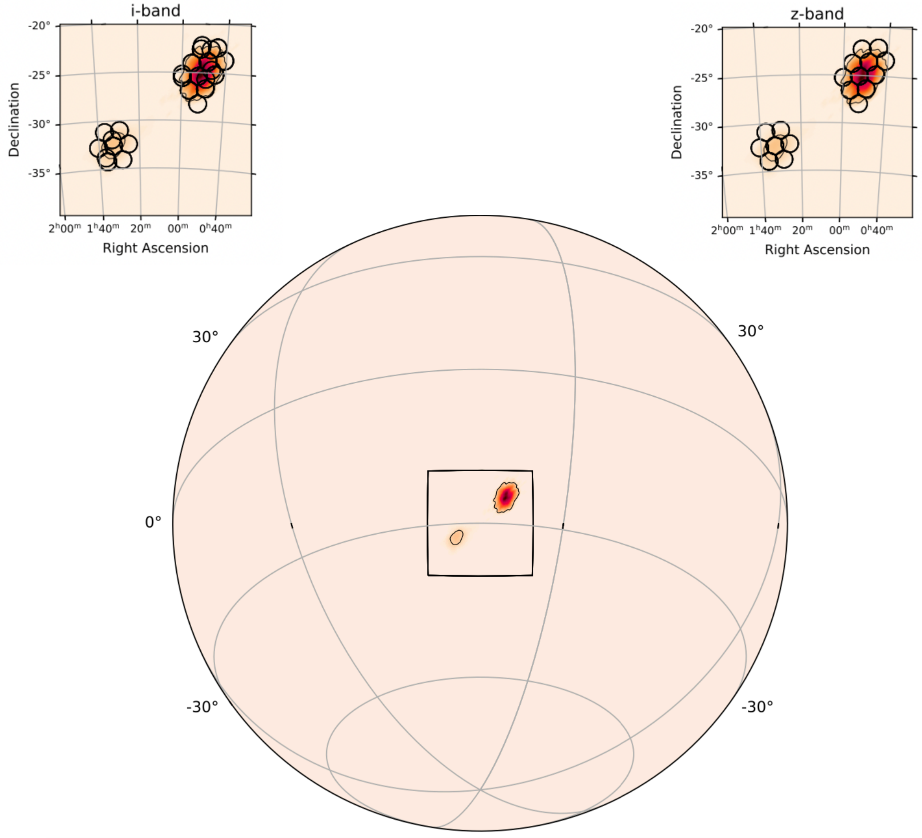


Hunting for kilonovae in O3

Constraining the parameter space of models from non-detections

[Anand, Coughlin, Kasliwal, **MB** +2020, Nature Astronomy] [Andreoni..**MB**..+2020a, ApJ]





S190814bv

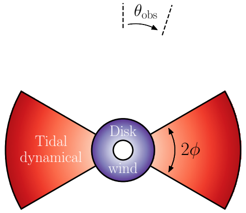
(now GW190814)

23 deg^2

267 Mpc

Possible NS-BH

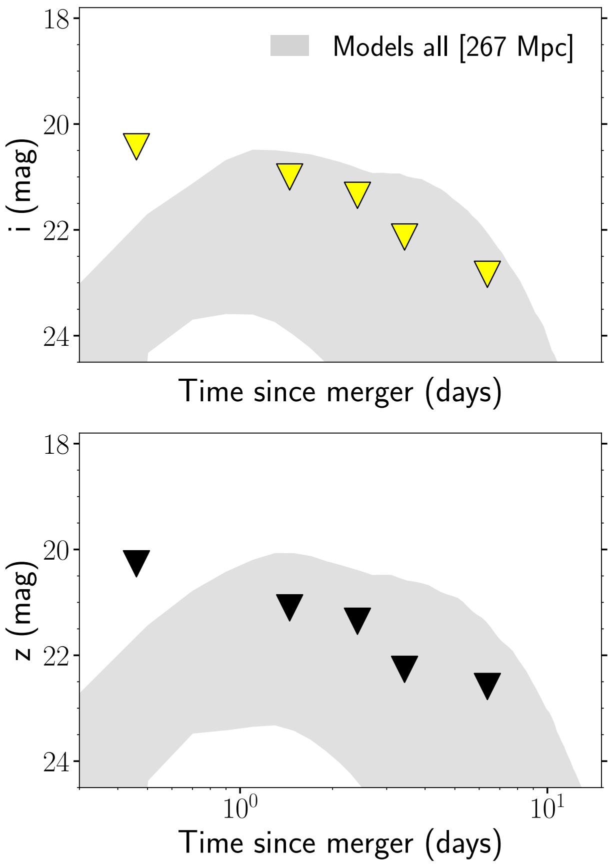
see also [Coughlin..**MB**..+2020ab, MNRAS / Ackley..**MB**..+2020, A&A / Andreoni..**MB**..+2020b, ApJ / Kasliwal..**MB**..,+2020, ApJ / Andreoni..**MB**..,+2021, ApJ]

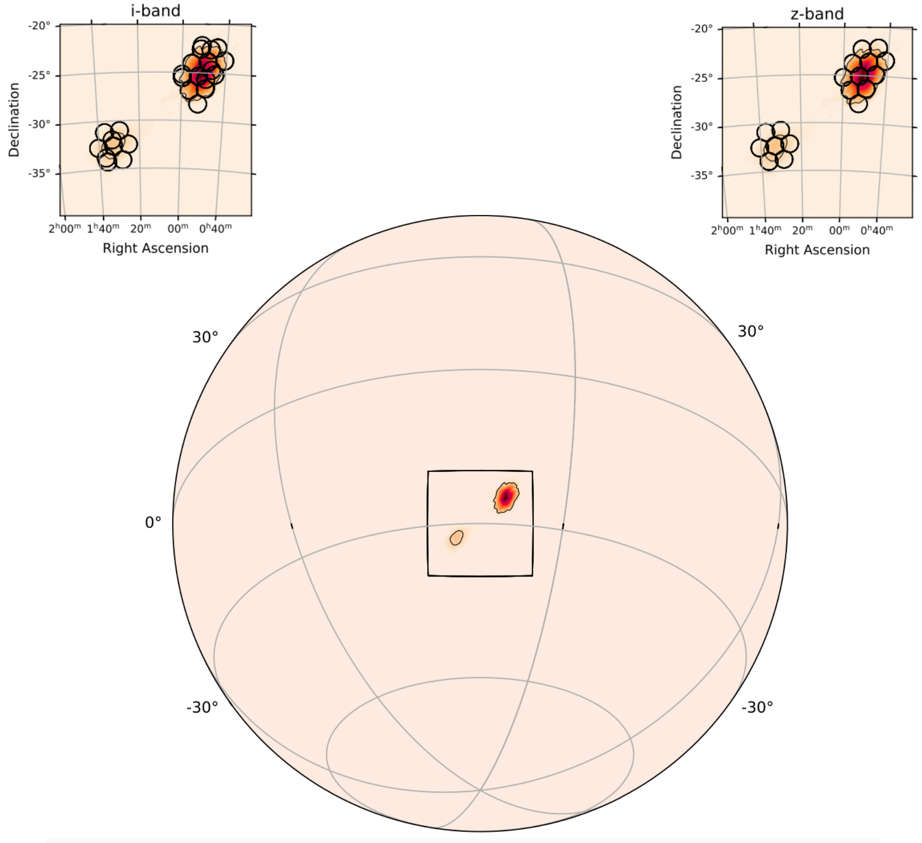


Hunting for kilonovae in O3

Constraining the parameter space of models from non-detections

[Anand, Coughlin, Kasliwal, **MB** +2020, Nature Astronomy] [Andreoni..**MB**..+2020a, ApJ]





S190814bv

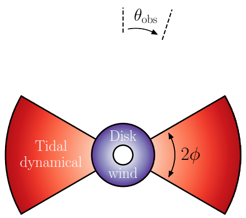
(now GW190814)

23 deg^2

267 Mpc

Possible NS-BH

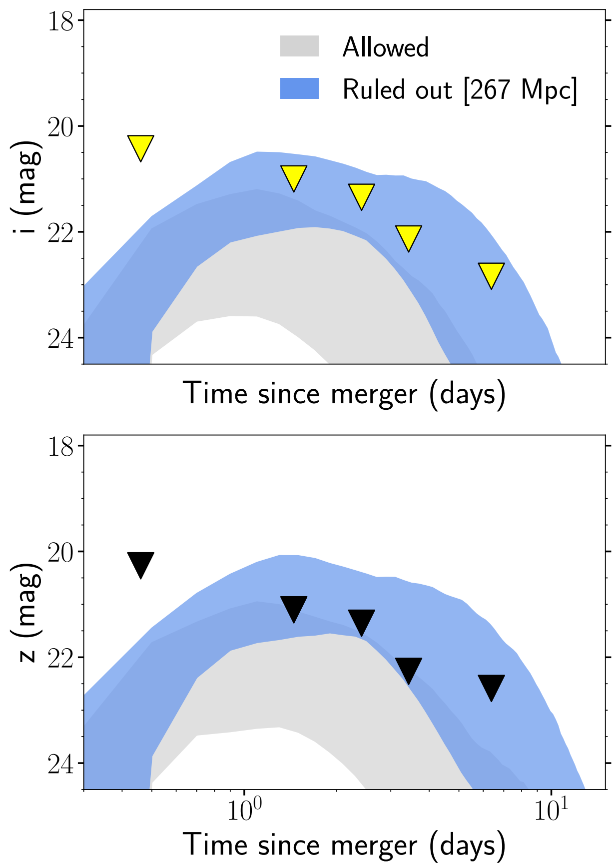
see also [Coughlin..**MB**..+2020ab, MNRAS / Ackley..**MB**..+2020, A&A / Andreoni..**MB**..+2020b, ApJ / Kasliwal..**MB**..,+2020, ApJ / Andreoni..**MB**..,+2021, ApJ]

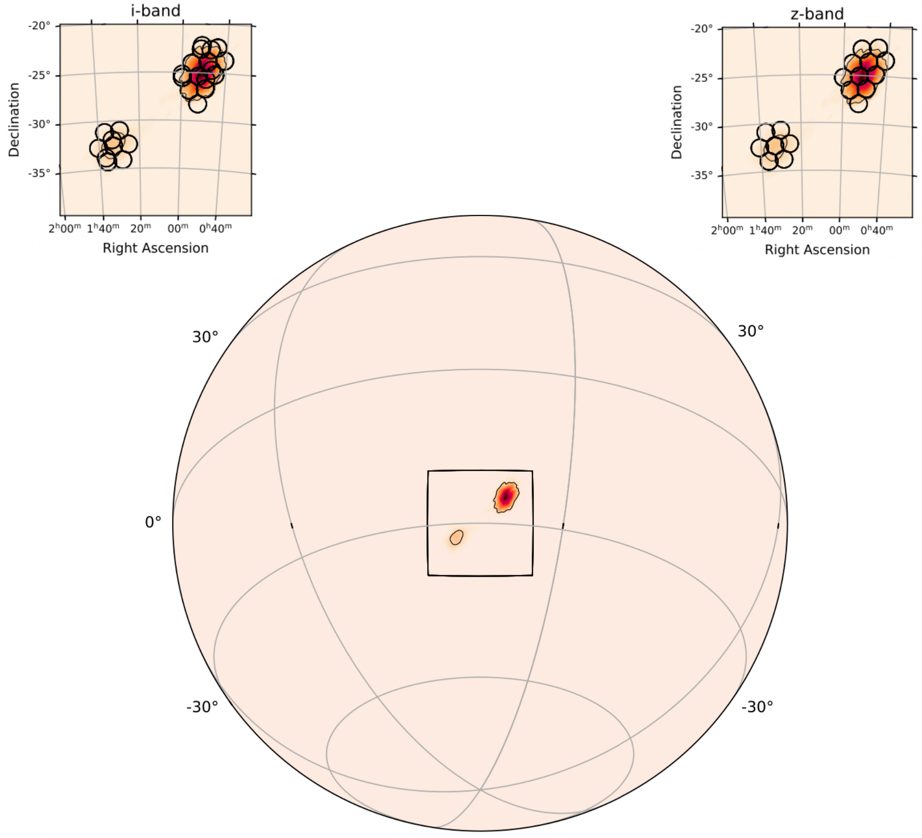


Hunting for kilonovae in O3

Constraining the parameter space of models from non-detections

[Anand, Coughlin, Kasliwal, **MB** +2020, Nature Astronomy] [Andreoni..**MB**..+2020a, ApJ]





S190814bv
(now GW190814)

23 deg^2
267 Mpc

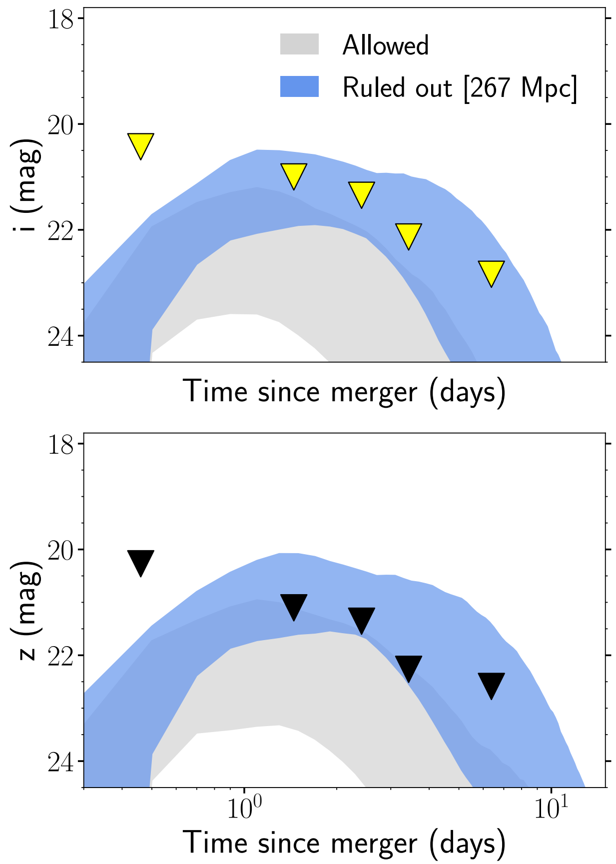
Possible NS-BH

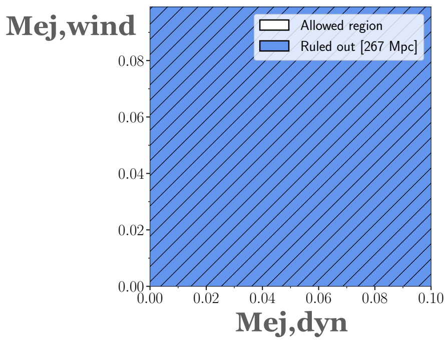
see also [Coughlin..**MB**..+2020ab, MNRAS / Ackley..**MB**..+2020, A&A / Andreoni..**MB**..+2020b, ApJ / Kasliwal..**MB**..,+2020, ApJ / Andreoni..**MB**..,+2021, ApJ]

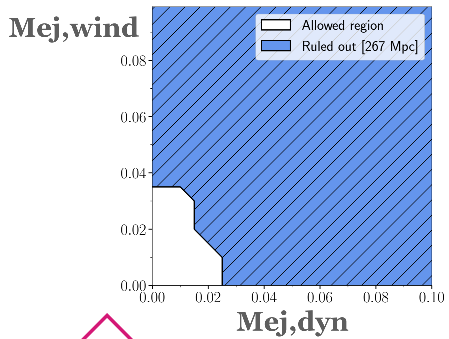
Hunting for kilonovae in O3

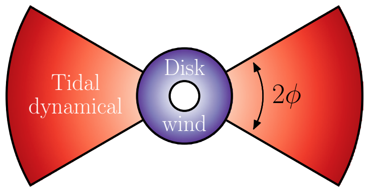
Constraining the parameter space of models from non-detections

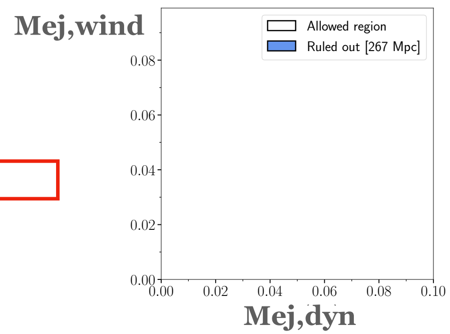
[Anand, Coughlin, Kasliwal, **MB** +2020, Nature Astronomy] [Andreoni..**MB**..+2020a, ApJ]









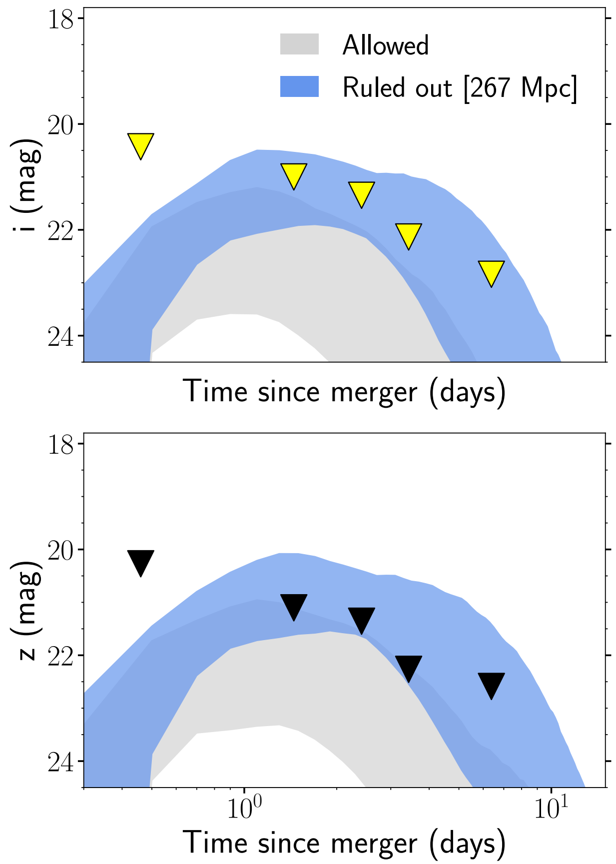


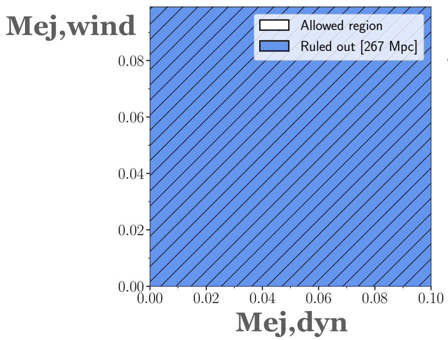
see also [Coughlin..**MB**..+2020ab, MNRAS / Ackley..**MB**..+2020, A&A / Andreoni..**MB**..+2020b, ApJ / Kasliwal..**MB**..,+2020, ApJ / Andreoni..**MB**..,+2021, ApJ]

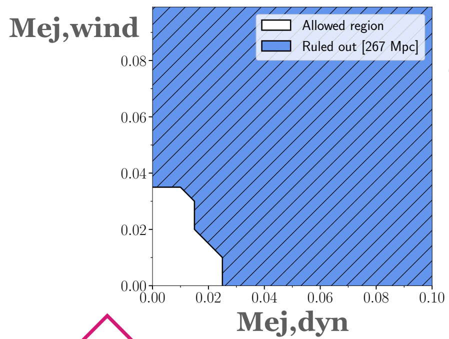
Hunting for kilonovae in O3

Constraining the parameter space of models from non-detections

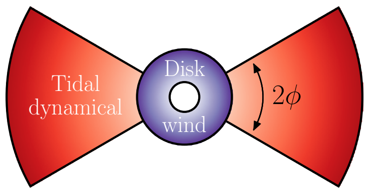
[Anand, Coughlin, Kasliwal, **MB** +2020, Nature Astronomy] [Andreoni..**MB**..+2020a, ApJ]

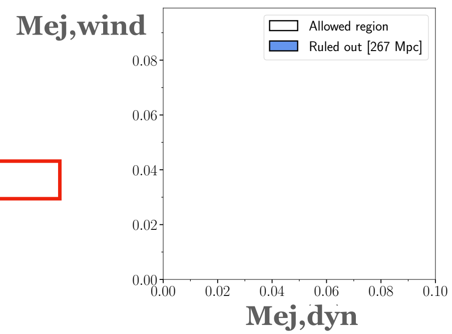






Rule out systems with **low mass ratio**, **high BH spin** or **large NS radii**

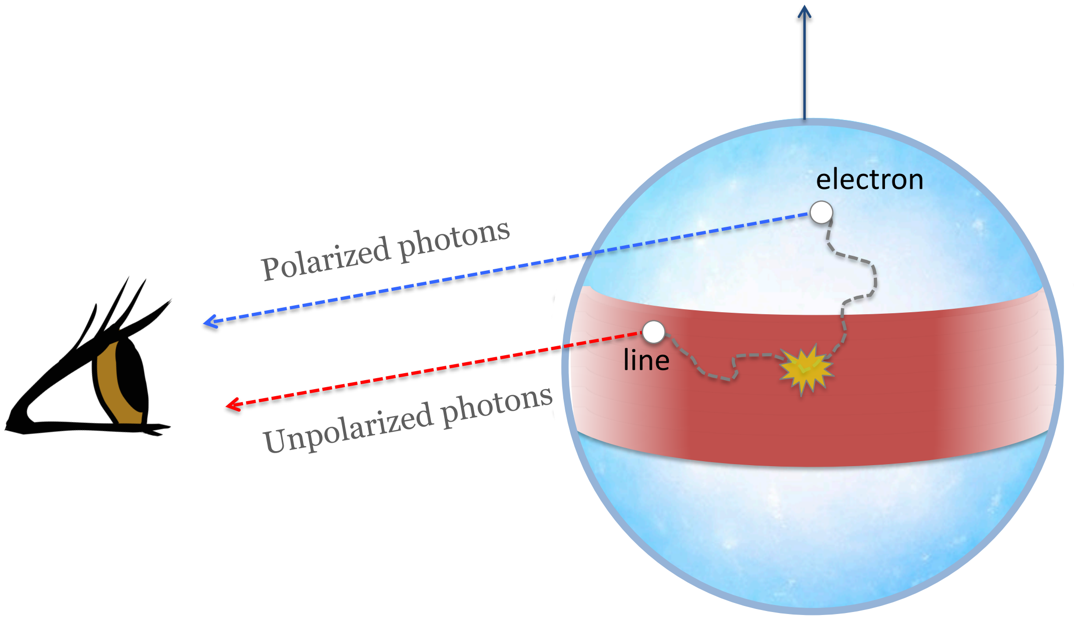


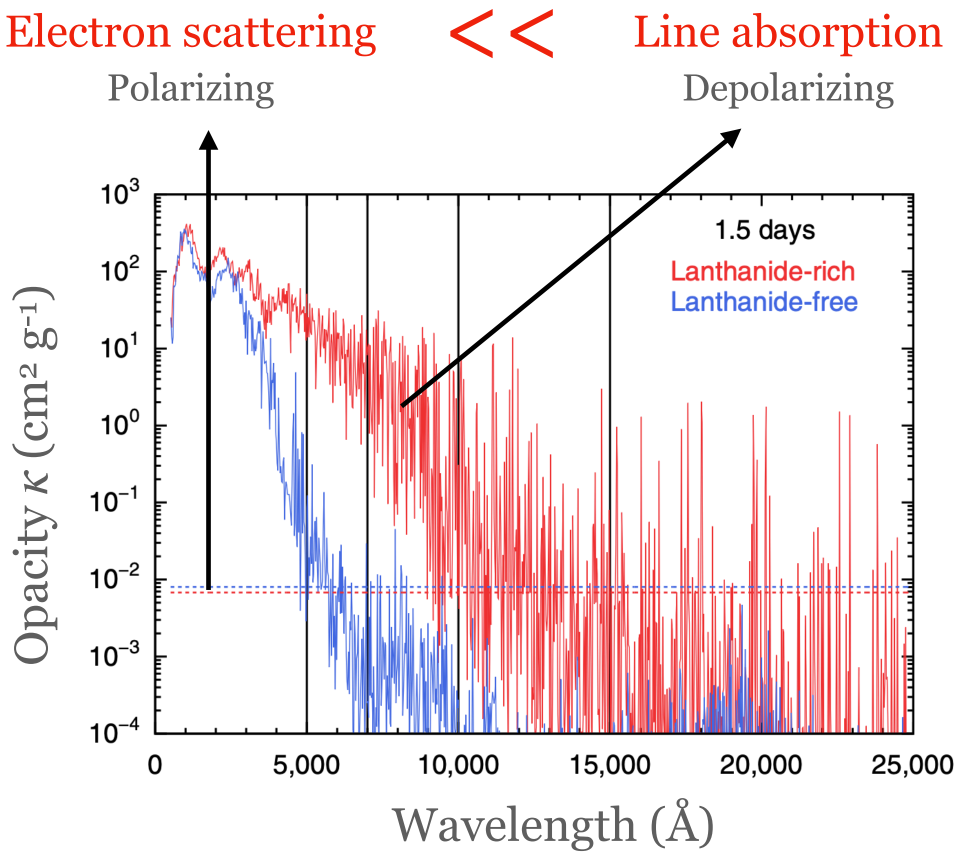


see also [Coughlin..**MB**..+2020ab, MNRAS / Ackley..**MB**..+2020, A&A / Andreoni..**MB**..+2020b, ApJ / Kasliwal..**MB**..,+2020, ApJ / Andreoni..**MB**..,+2021, ApJ]

Polarization signal

Constraining the viewing angle and the presence of a lanthanide-free component

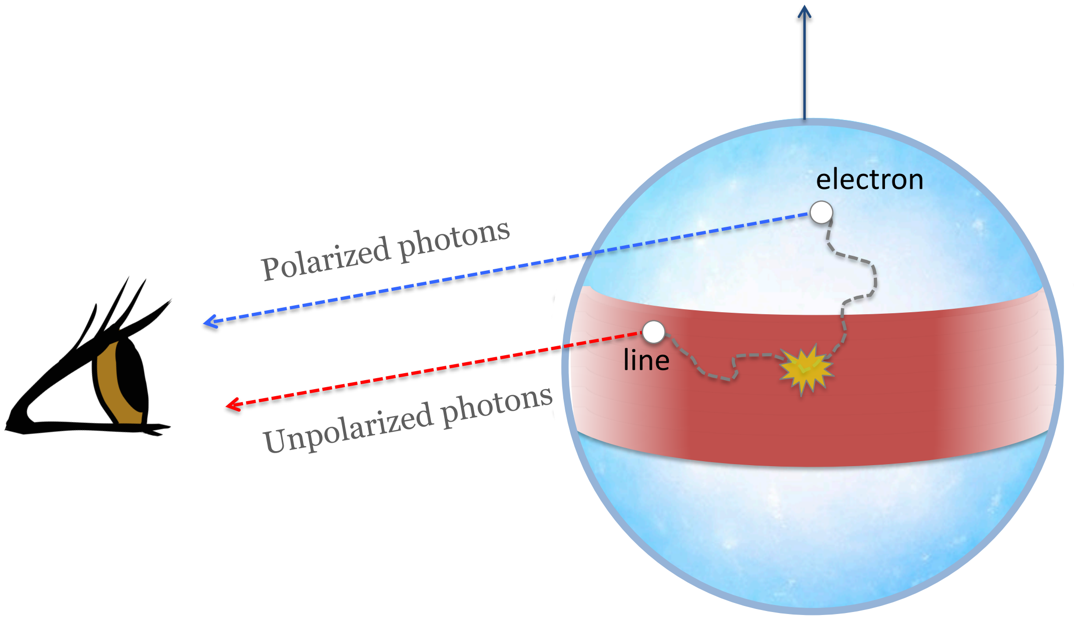


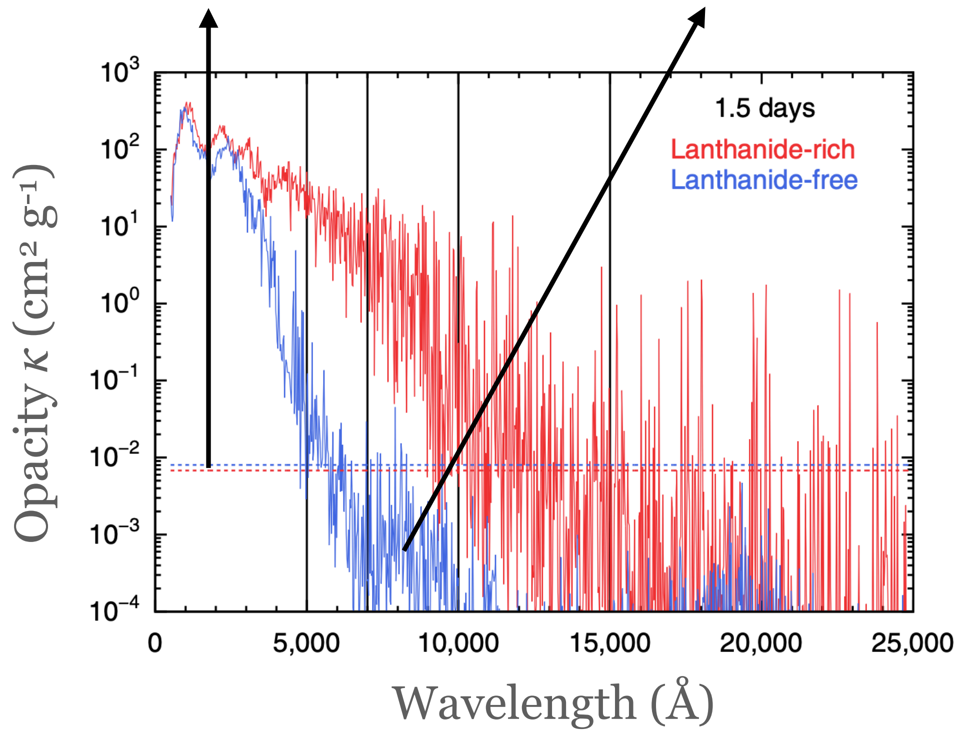


[Tanaka+2018, ApJ]

Polarization signal

Constraining the viewing angle and the presence of a lanthanide-free component





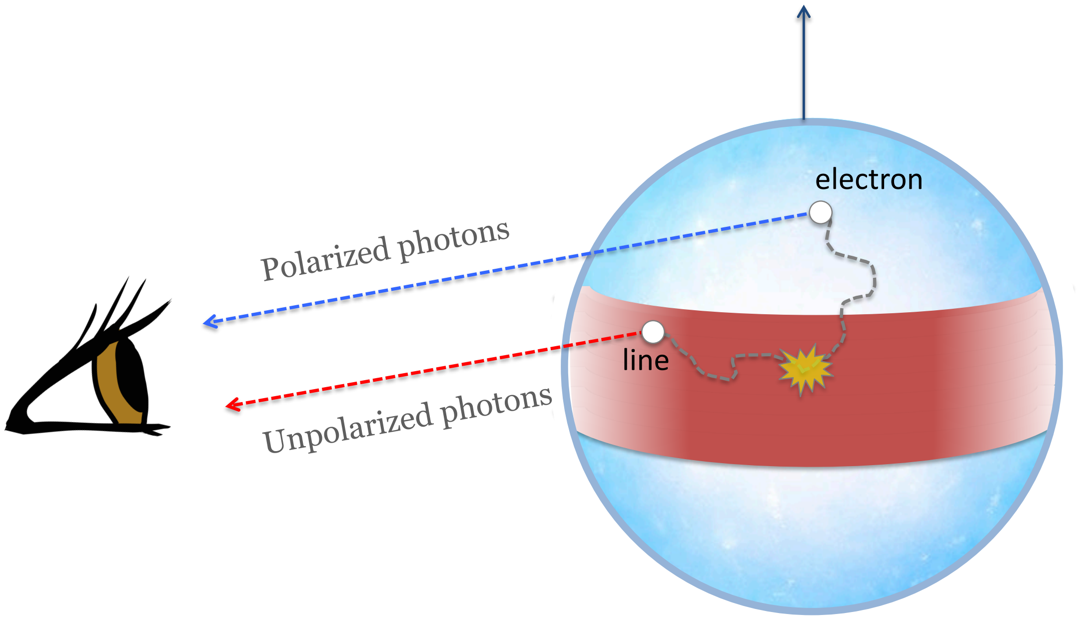
[Tanaka+2018, ApJ]

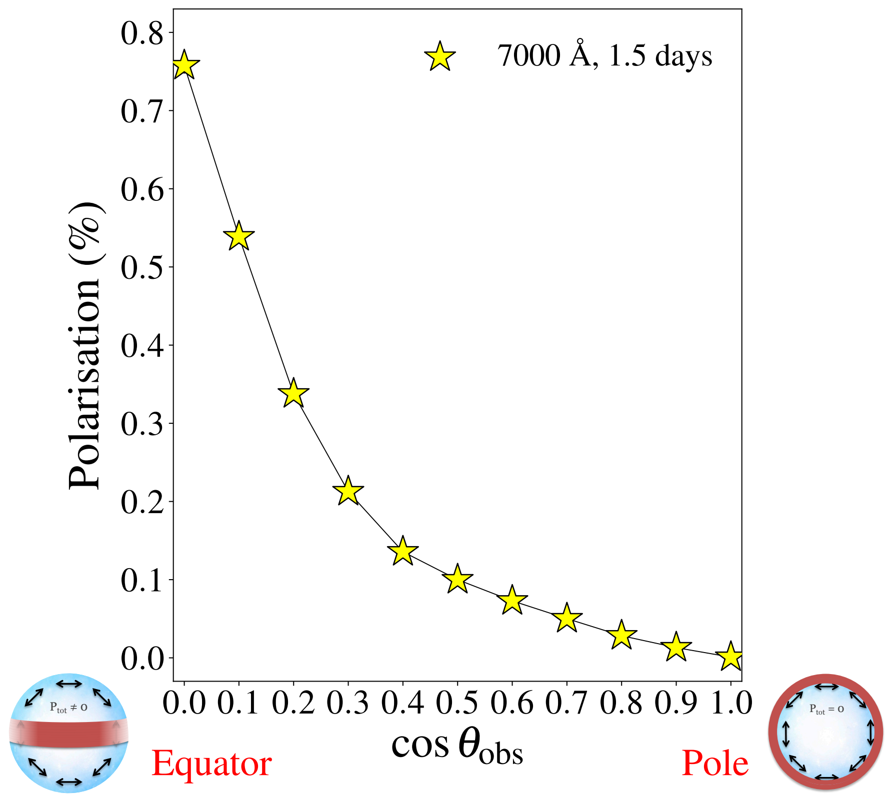
Polarization signal

Constraining the viewing angle and the presence of a lanthanide-free component

NS+NS

[**MB**+2019, Nature Astronomy]



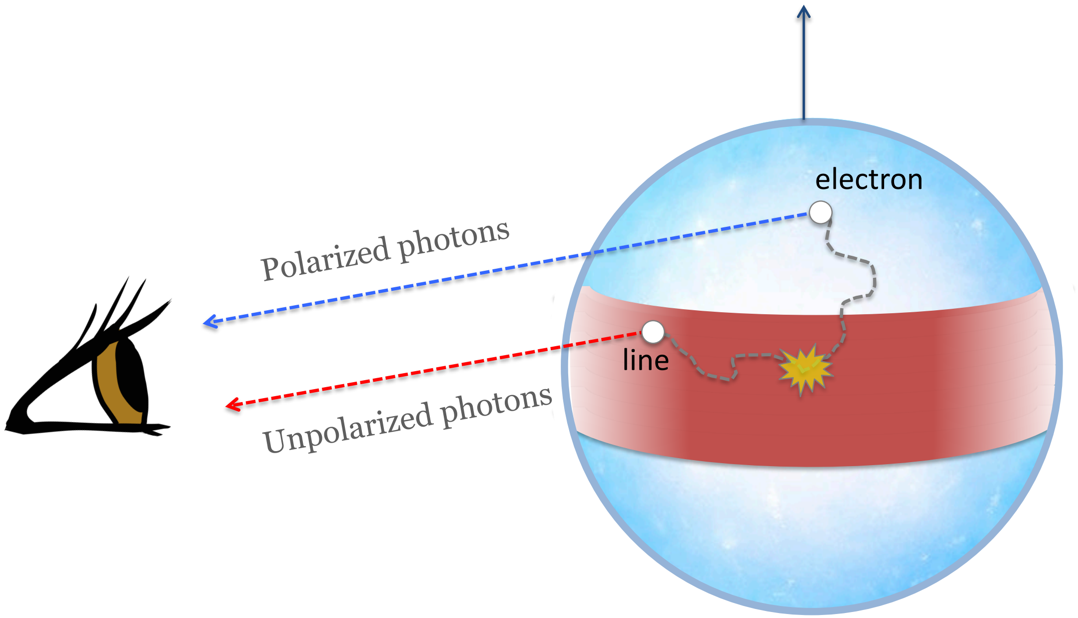


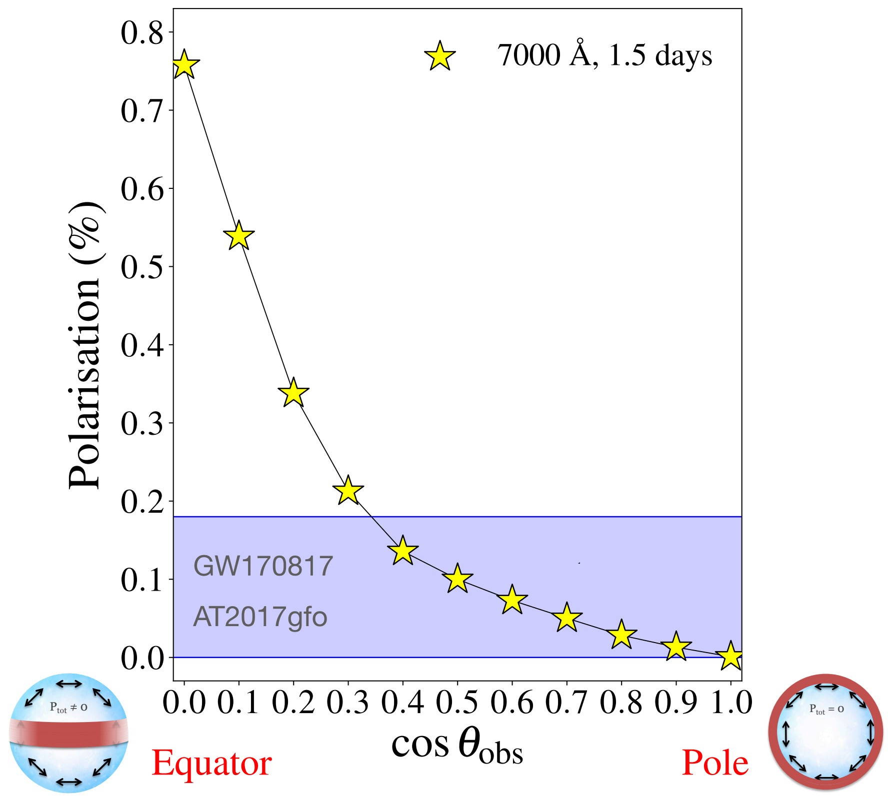
Polarization signal

Constraining the viewing angle and the presence of a lanthanide-free component

NS+NS

[**MB**+2019, Nature Astronomy]



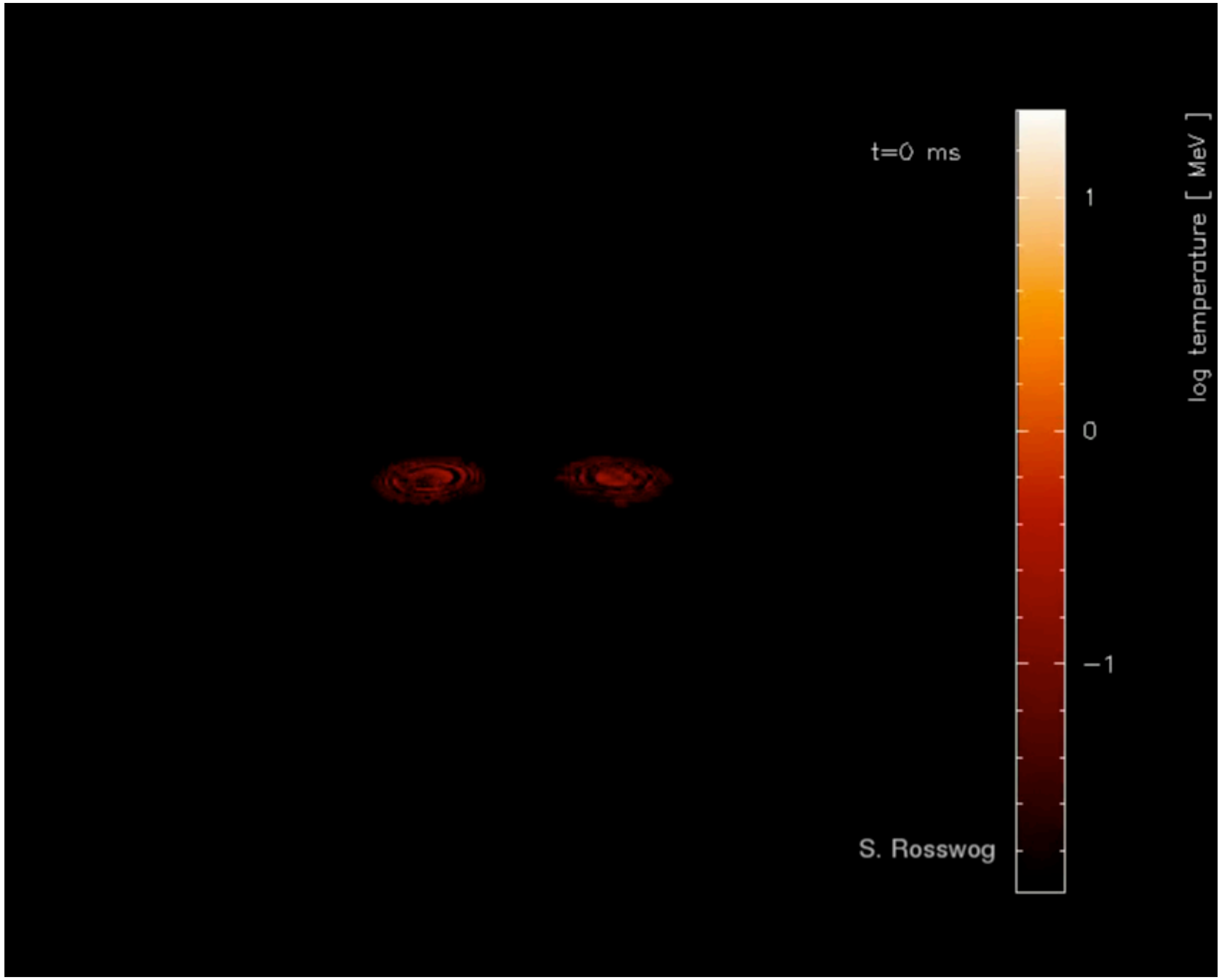


Kilonovae from NR simulations


Tim Dietrich


Vivek Chaurasia


Stephan Rosswog

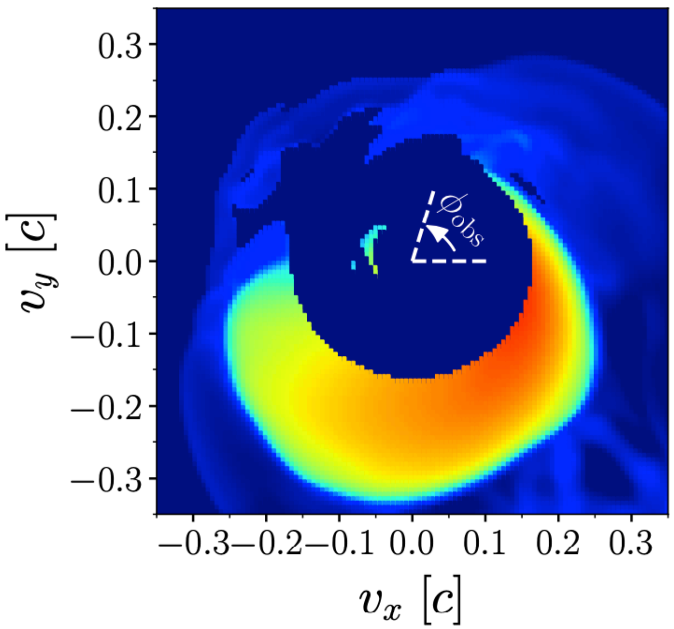


Polarization signal

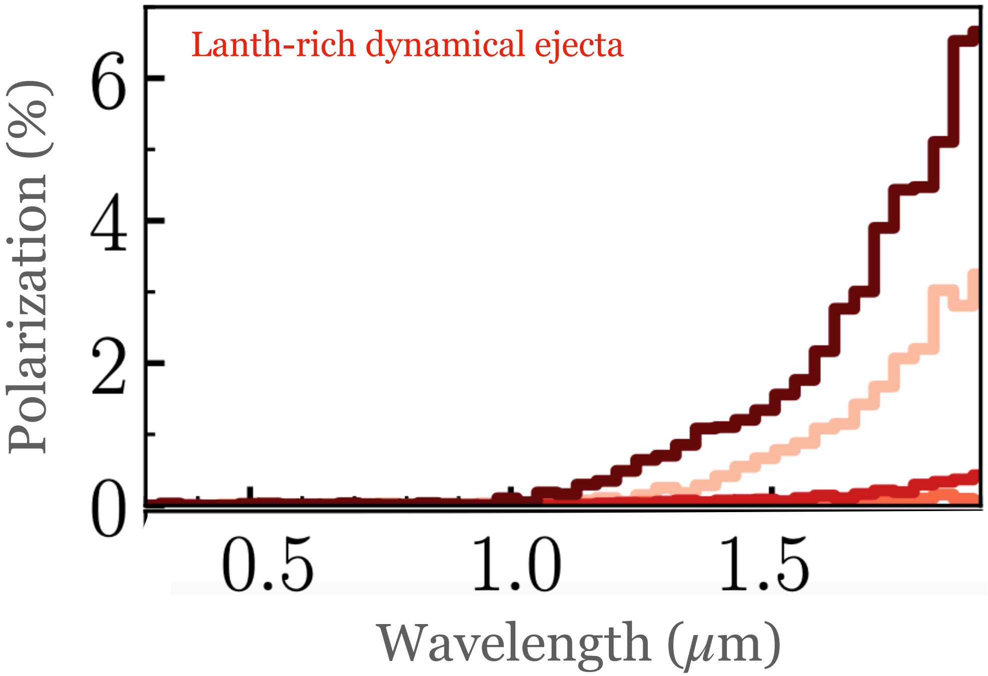
Constraining the presence of a lanthanide-free component

BH+NS

[**MB**, Kyutoku+2021, MNRAS]



Model for dynamical ejecta from [Kyutoku+2015, PRD]

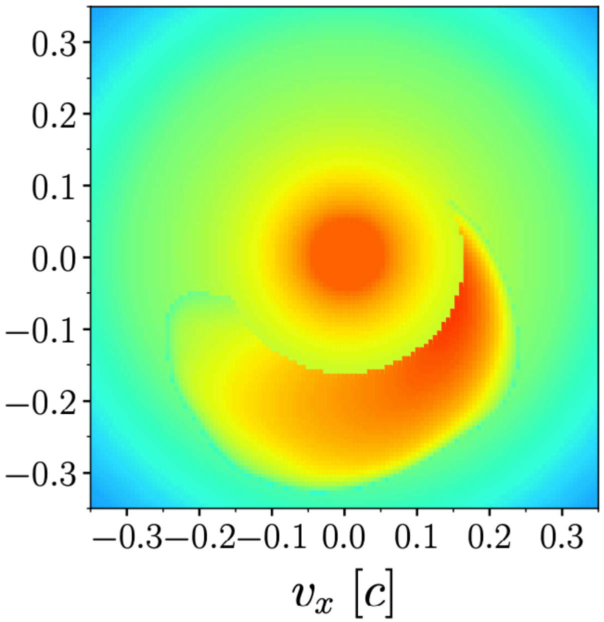


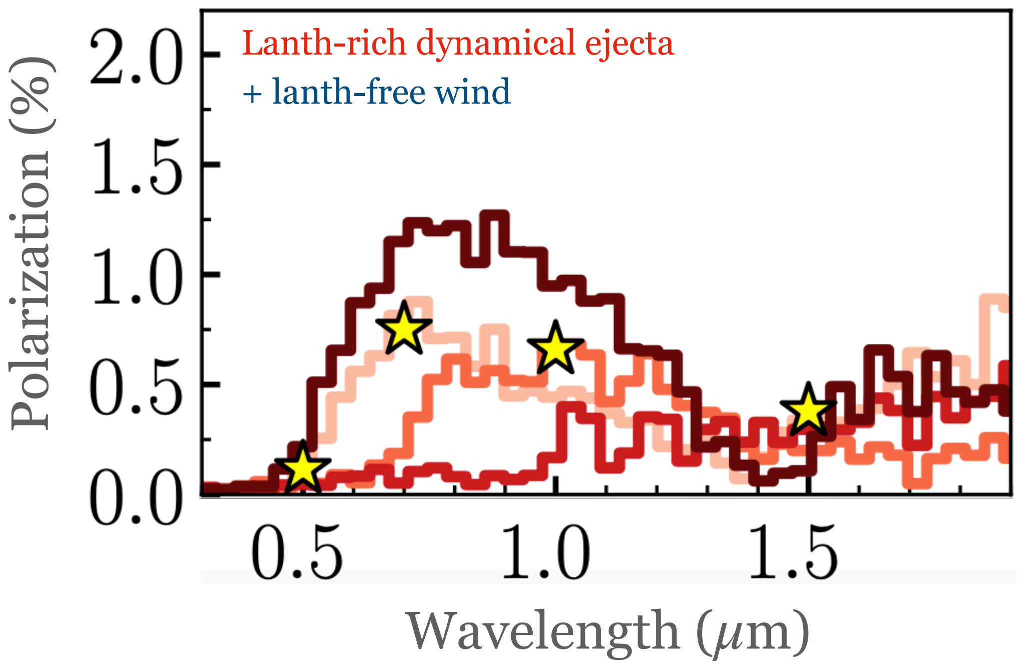
Polarization signal

Constraining the presence of a lanthanide-free component

BH+NS

[**MB**+2021, MNRAS]



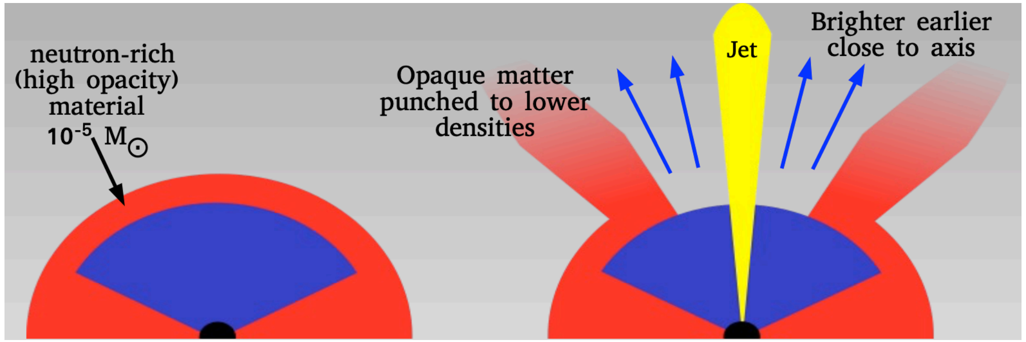


Model for dynamical ejecta from [Kyutoku+2015, PRD]

Jet-ejecta interaction

Making kilonovae brighter and bluer

[Nativi, **MB**, Lundman, Rosswog+2021, MNRAS]

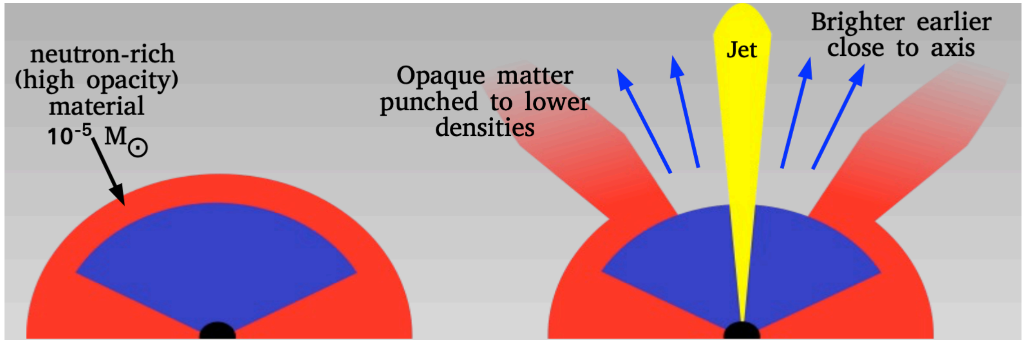


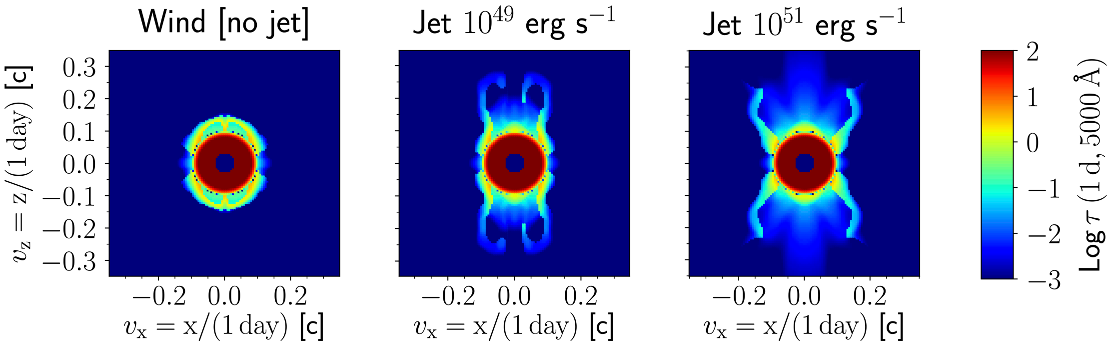
see also [Klion, Duffert, Kasen & Quataert 2021, MNRAS]

Jet-ejecta interaction

Making kilonovae brighter and bluer

[Nativi, **MB**, Lundman, Rosswog+2021, MNRAS]





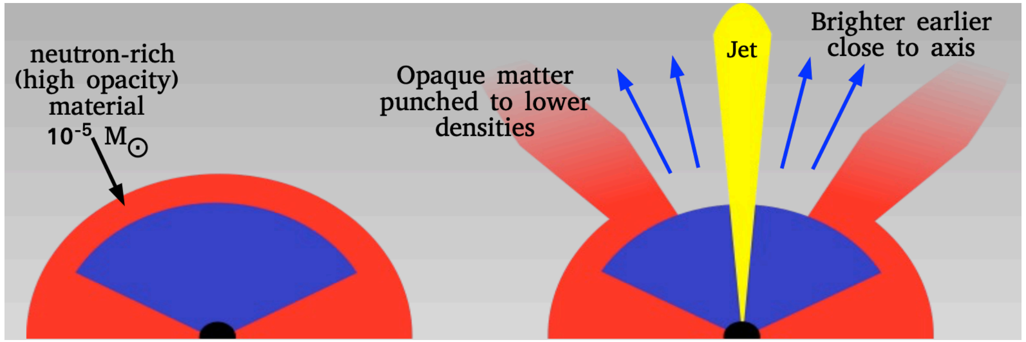
Wind models from [Perego+2014, MNRAS]

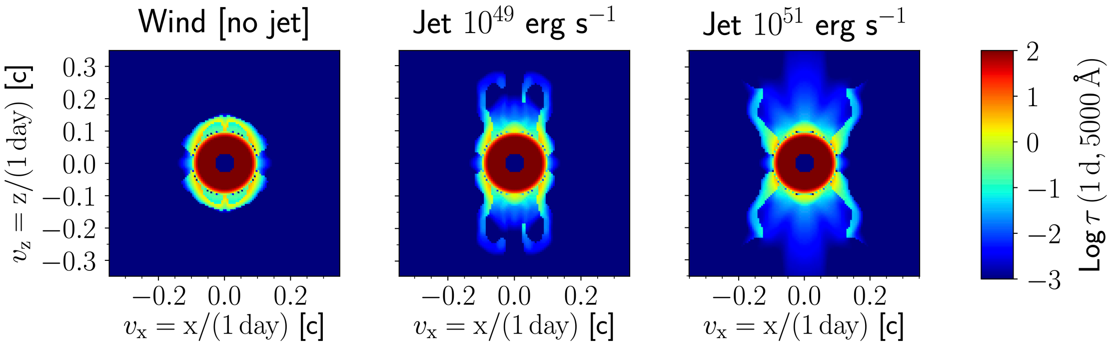
see also [Klion, Duffert, Kasen & Quataert 2021, MNRAS]

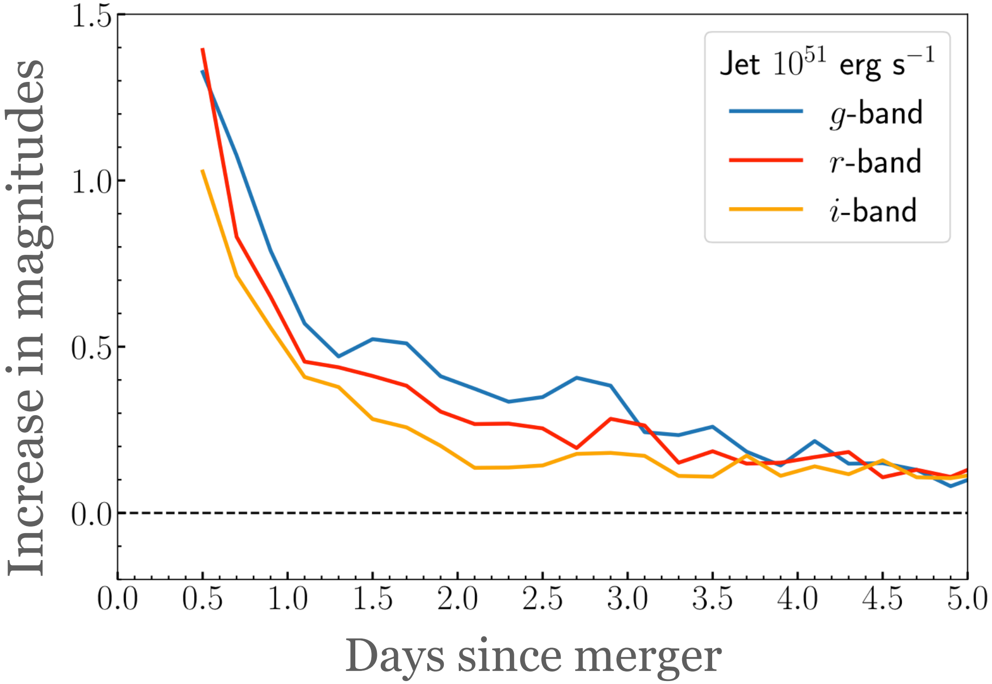
Jet-ejecta interaction

Making kilonovae brighter and bluer

[Nativi, **MB**, Lundman, Rosswog+2021, MNRAS]







Wind models from [Perego+2014, MNRAS]

see also [Klion, Duffert, Kasen & Quataert 2021, MNRAS]

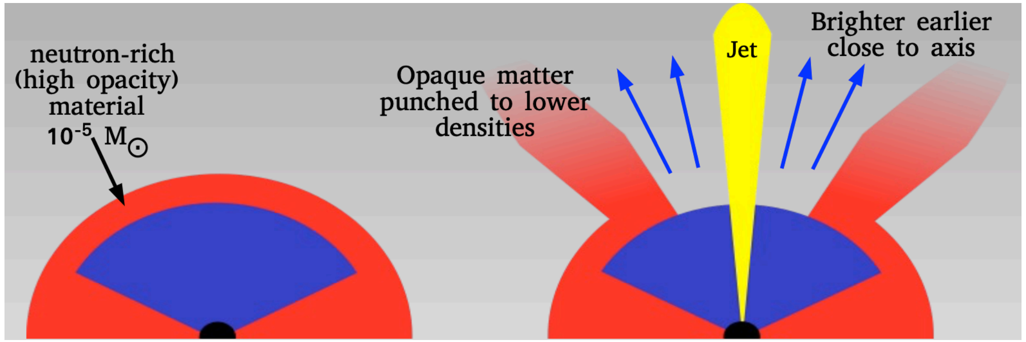
Jet-ejecta interaction

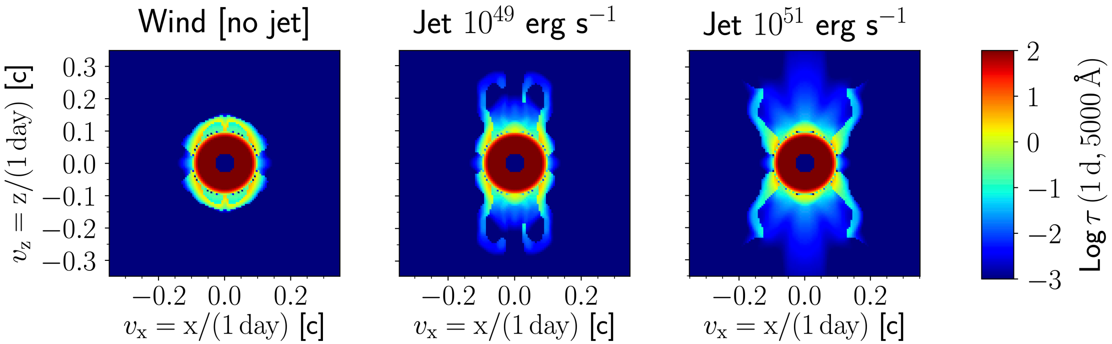
Making kilonovae brighter and bluer

[Shrestha & **MB**, in prep]

[Nativi, **MB**, Lundman, Rosswog+2021, MNRAS]

Impact on polarization?





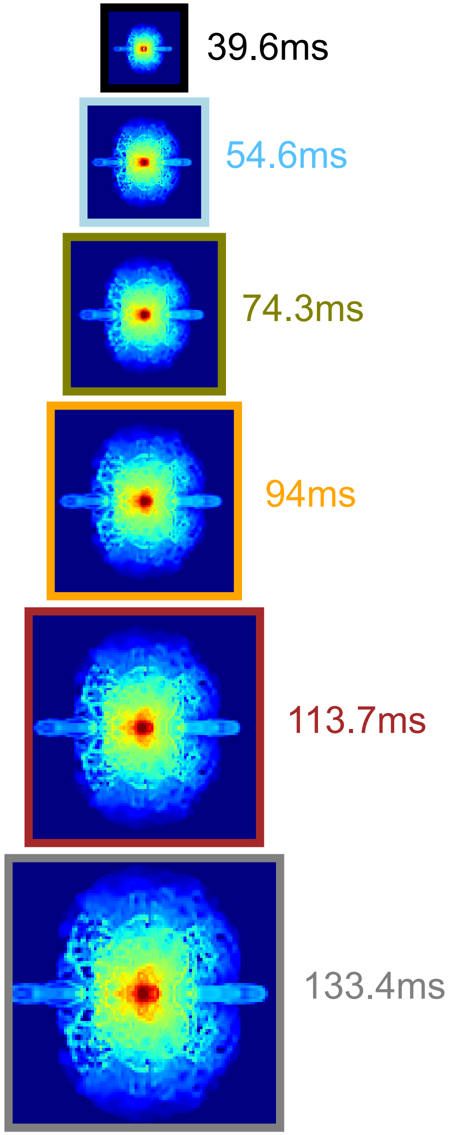
Wind models from [Perego+2014, MNRAS]

see also [Klion, Duffert, Kasen & Quataert 2021, MNRAS]

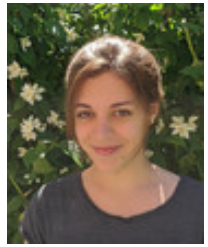
Homologous expansion

Reached from ~100ms [dynamical ejecta only]

[Neuweiler, Dietrich, **MB**+, arXiv:2208.13460]



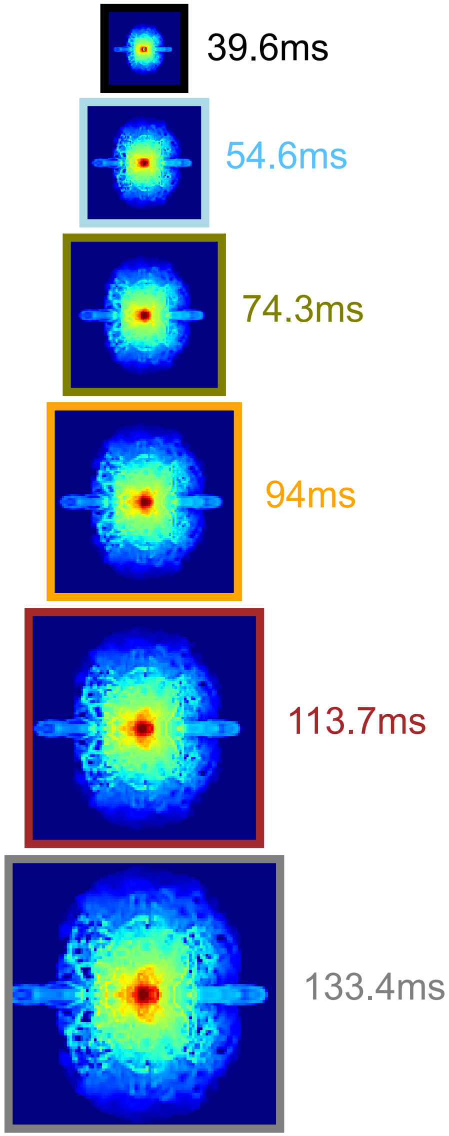
BAM code, q=1, EoS: H4



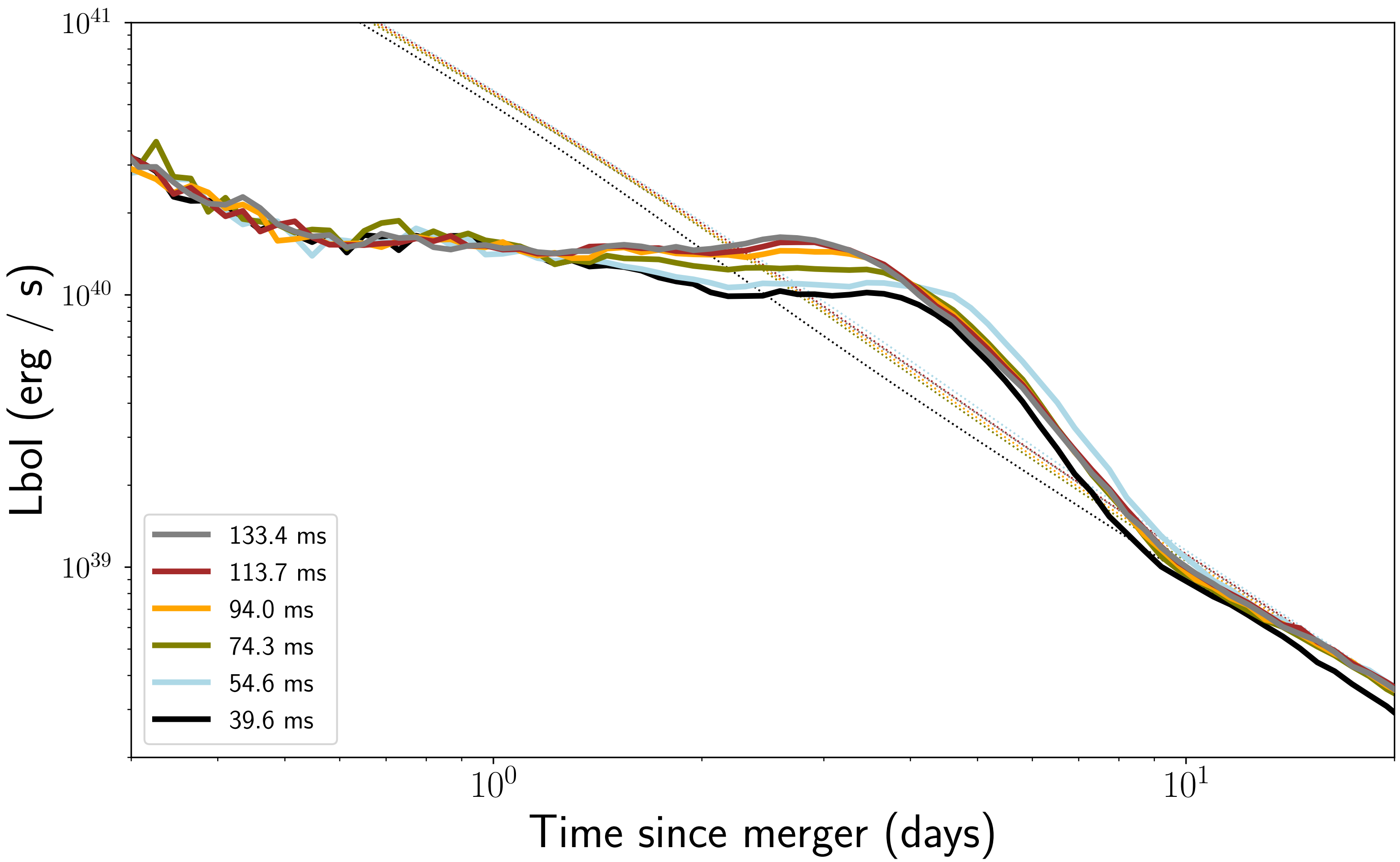
Homologous expansion

Reached from ~100ms [dynamical ejecta only]

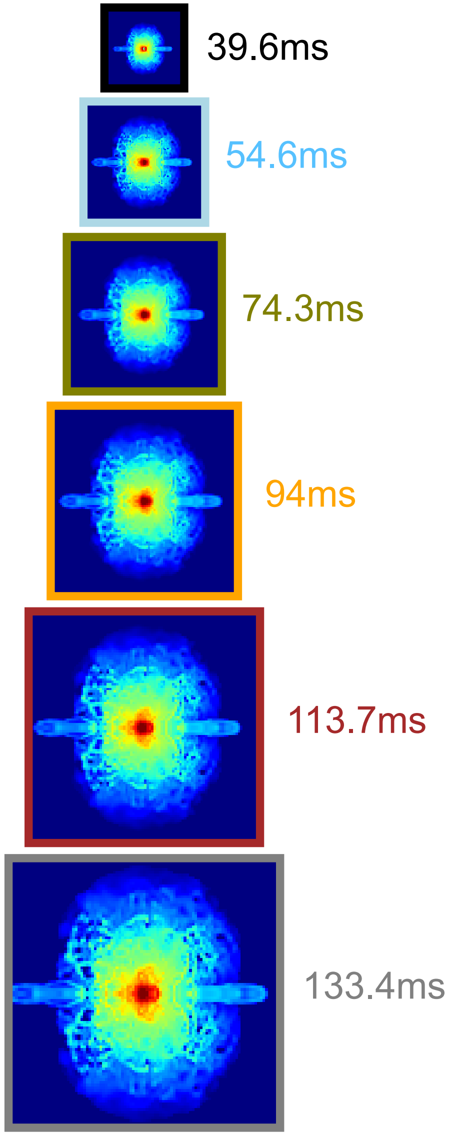
[Neuweiler, Dietrich, **MB**+, arXiv:2208.13460]



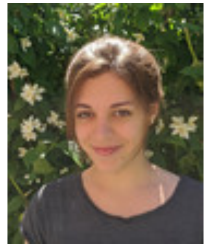
BAM code, q=1, EoS: H4



Homologous expansion

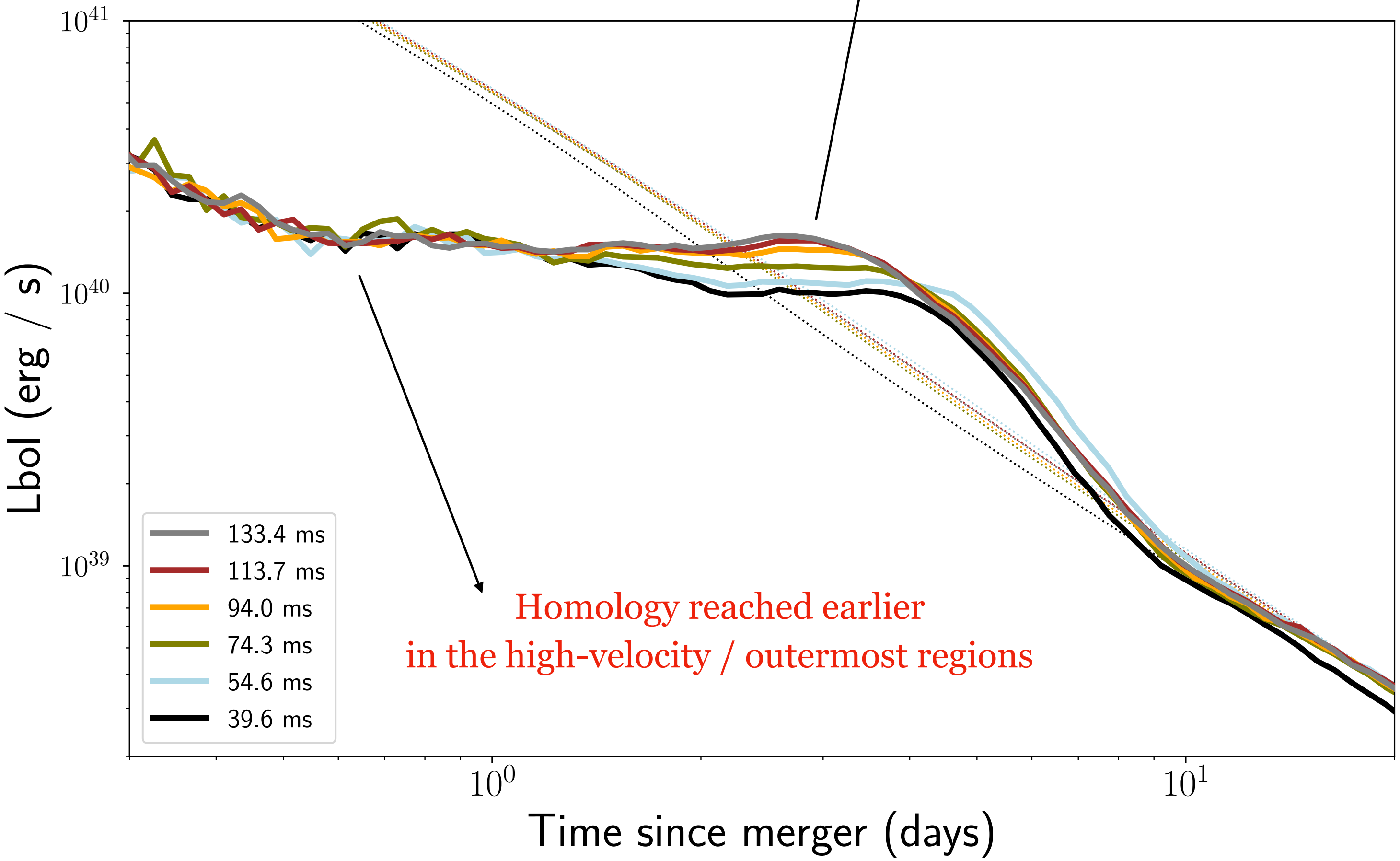


BAM code, q=1, EoS: H4

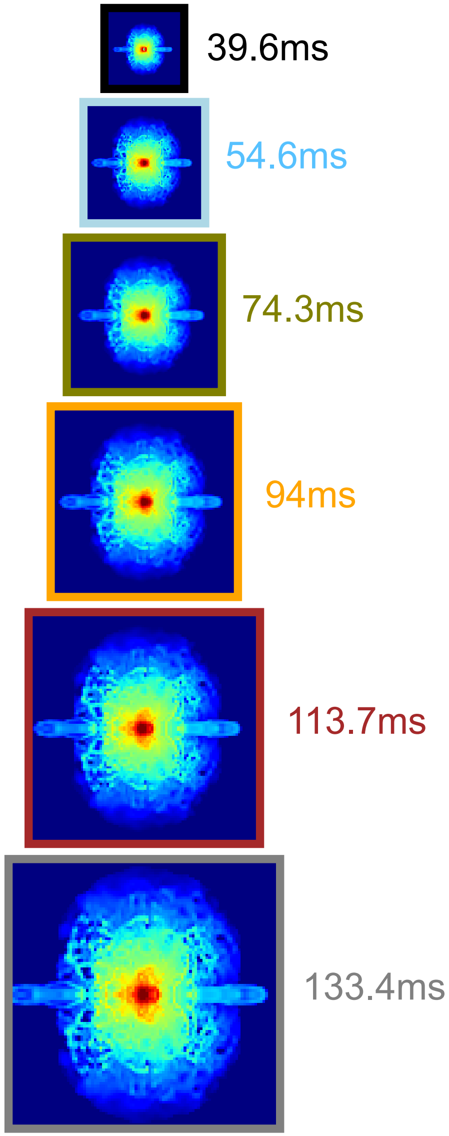


Reached from ~100ms [dynamical ejecta only]

[Neuweiler, Dietrich, **MB**+, arXiv:2208.13460]



Homologous expansion



BAM code, q=1, EoS: H4

Reached from ~100ms [dynamical ejecta only]

[Neuweiler, Dietrich, **MB**+, arXiv:2208.13460]

JHK near-infrared light-curves