

Measurement of the Coulomb dissociation cross section of the neutron rich nitrogen isotopes $^{20,21}\text{N}$

Marko Röder

NAVI Meeting
February 2015, Darmstadt

Supported by
GSI F&E (DR-ZUBE),
NupNET NEDENSAA (05 P09 CRFN5),
Plattform für Detektortechnologie und -systeme



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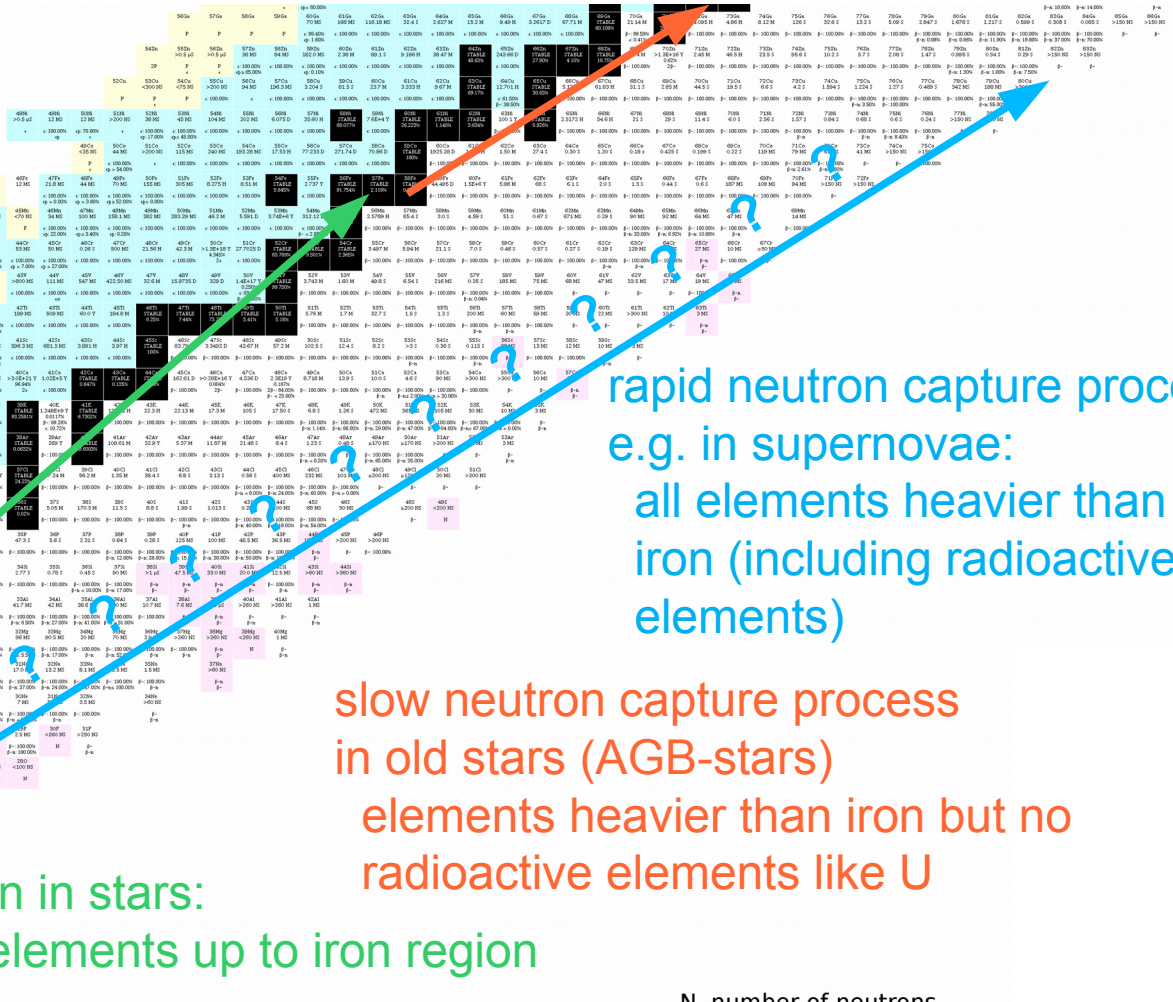


Member of the Helmholtz Association

Nucleosynthesis

Z, Number of protons

■ stable isotopes
 others are unstable isotopes



rapid neutron capture process
 e.g. in supernovae:
 all elements heavier than iron
 (including radioactive elements)

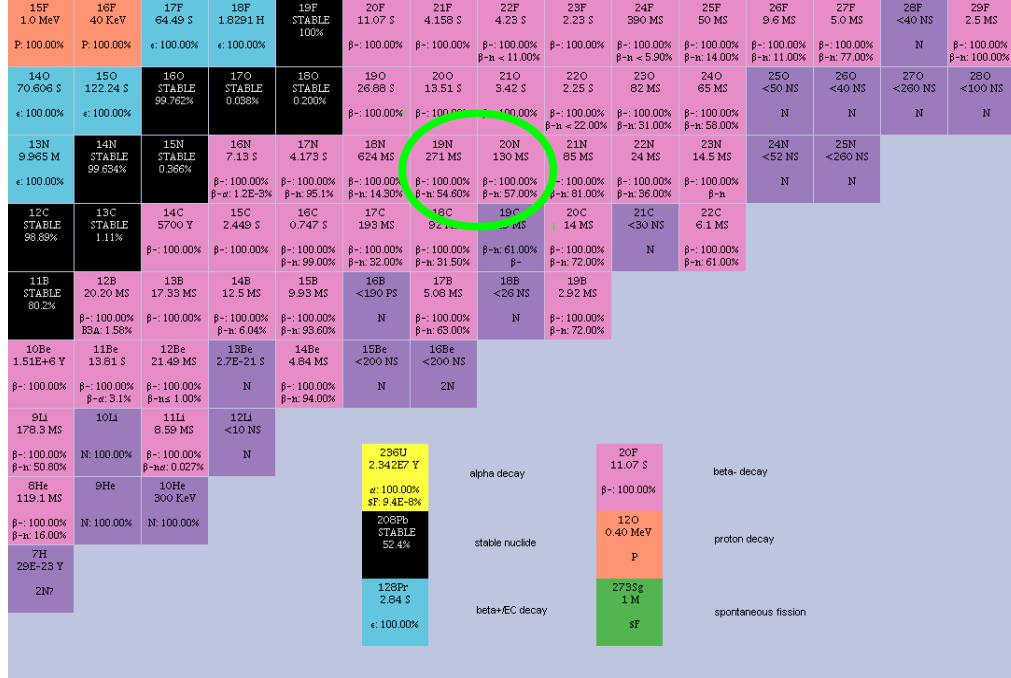
slow neutron capture process
 in old stars (AGB-stars)
 elements heavier than iron but no
 radioactive elements like U

fusion in stars:
 all elements up to iron region

Big Bang nucleosynthesis
 75% H, 25% He

N, number of neutrons

Astrophysical Motivation



How were the chemical elements created?

- elements up to iron by fusion in Big Bang and stars
- heavier elements by neutron capture processes
 - slow (s-process): discrepancies with observations
 - rapid (r-process): its path involves nuclei far away from the valley of stability (¹⁴N, ¹⁵N are stable)
- ²⁰N(γ ,n)¹⁹N studied to extract cross sections of ¹⁹N(n, γ)²⁰N via time inversion
- core-collapse supernova with neutrino driven wind scenario (Takahashi et al. 1994)
- Big Bang nucleosynthesis (Rauscher et al. 1994)

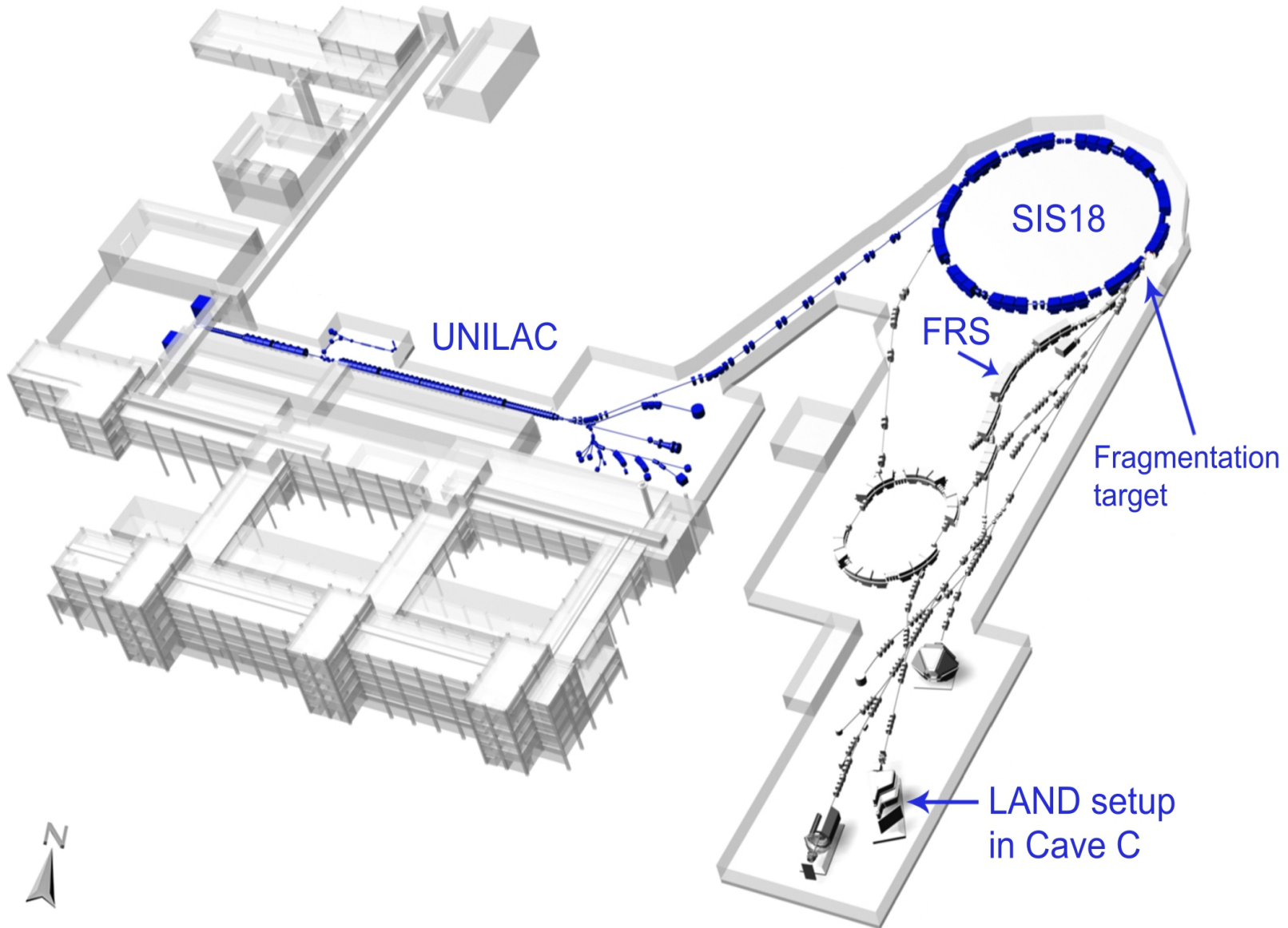
200 13.51 \$ P: 100.00%	210 3.42 \$ P: 100.00%	220 2.25 \$ P: 100.00% P-n < 22.00%
19N 271 MS P: 100.00% P-n: 54.60%	20N 130 MS P: 100.00% P-n: 57.00%	21N 85 MS P: 100.00% P-n: 81.00%
18C 92 MS P: 100.00% P-n: 31.50%	19C 49 MS P: 100.00% P-n: 61.00% P-	20C 14 MS P: 100.00% P-n: 72.00%

Impact of Light Neutron Rich Nuclei to r-process Abundance

M. Terasawa et al. (2001):

“[...] We find that a new nuclear reaction flow path opens in the very **light, neutron-rich** region. This new nuclear reaction flow can change the **final heavy-element abundances** by as much as an **order of magnitude**. [...]”

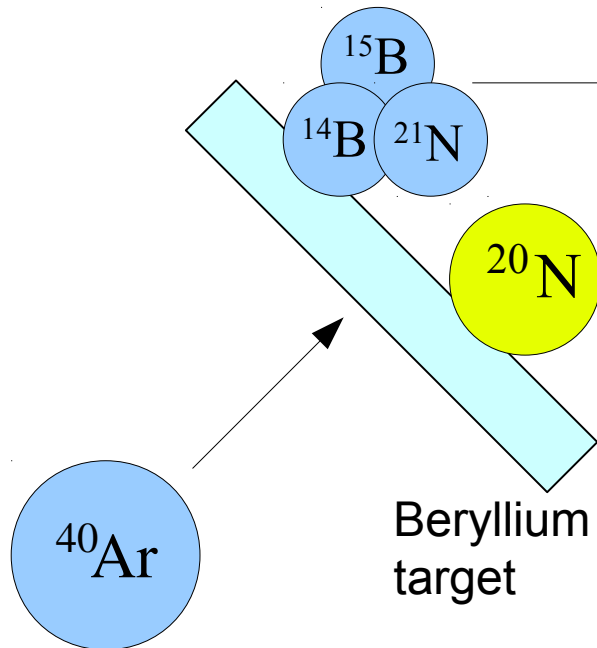
GSI Accelerator Facility



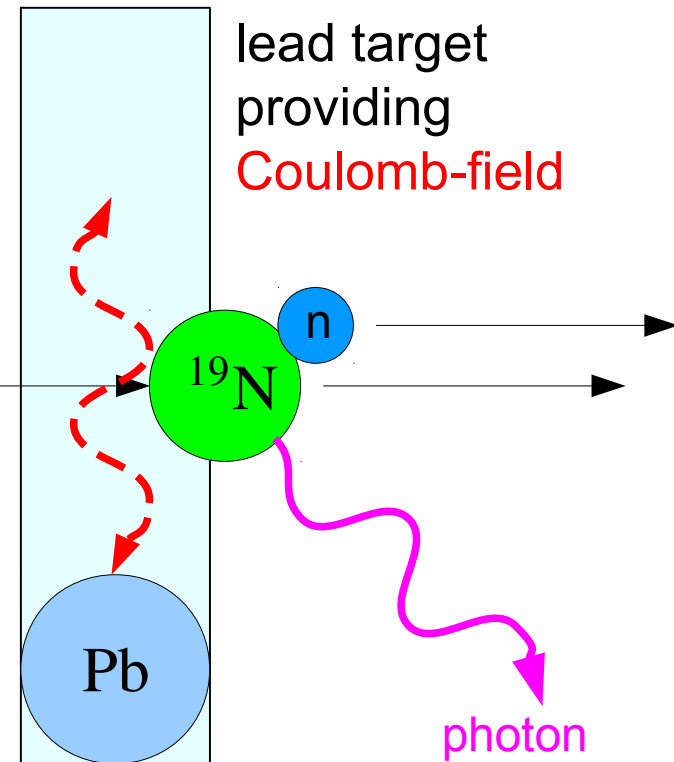
Coulomb Dissociation

$$\sigma_{CoulEx} = p_{Pb} \cdot \frac{M_{Pb}}{d_{Pb} \cdot N_A} - p_C \cdot \alpha \cdot \frac{M_C}{d_C \cdot N_A} - p_{empty} \cdot \left(\frac{M_{Pb}}{d_{Pb} \cdot N_A} - \alpha \cdot \frac{M_C}{d_C \cdot N_A} \right)$$

Radioactive Ion Beam Production

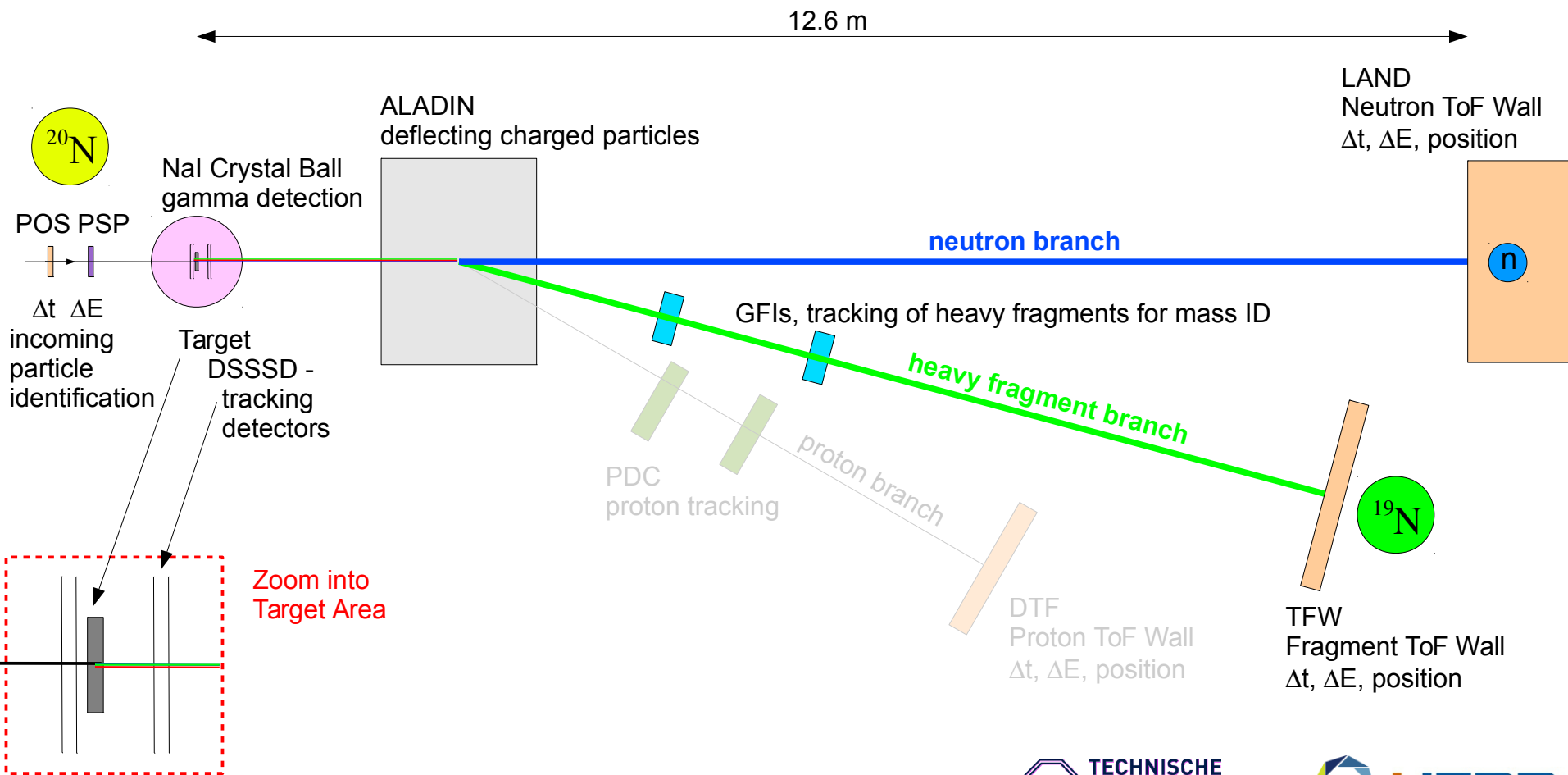
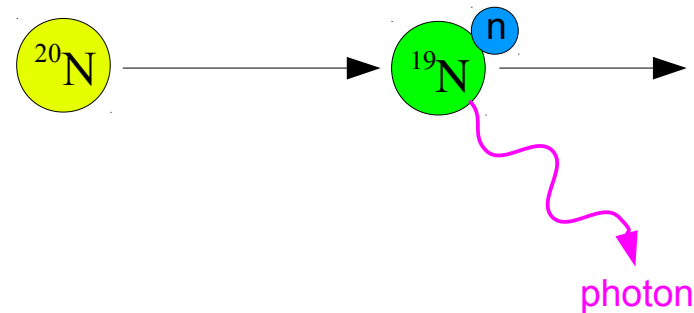


Reaction at R3B Setup

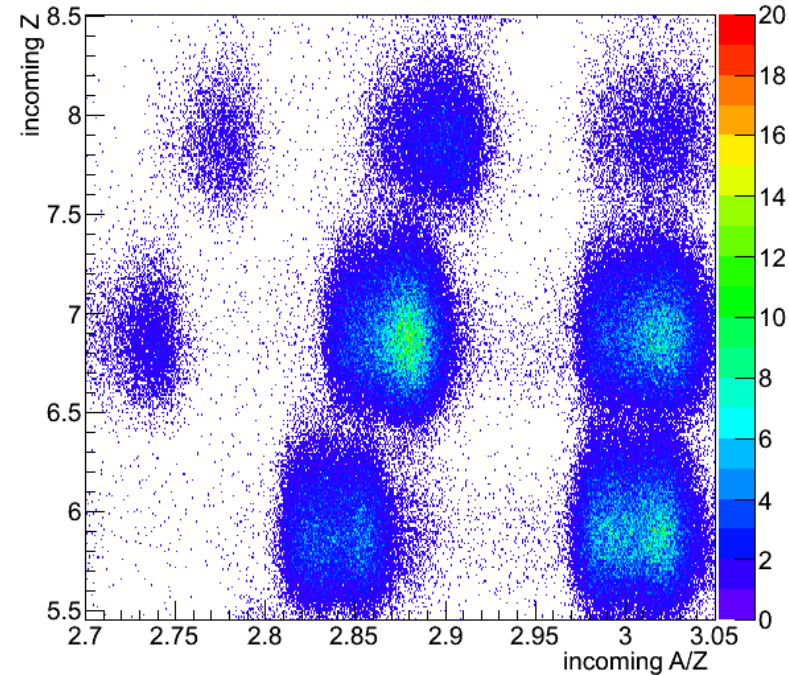
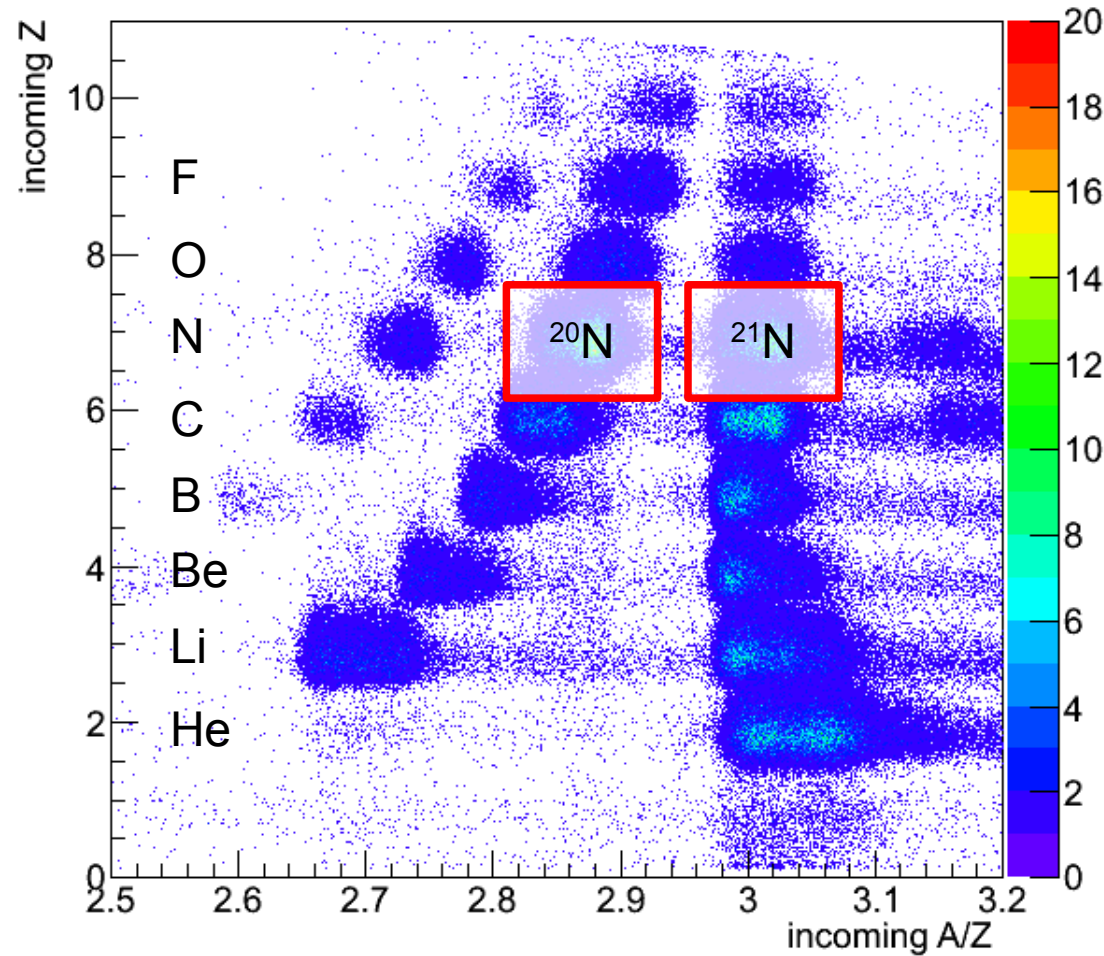


LAND/R3B Setup

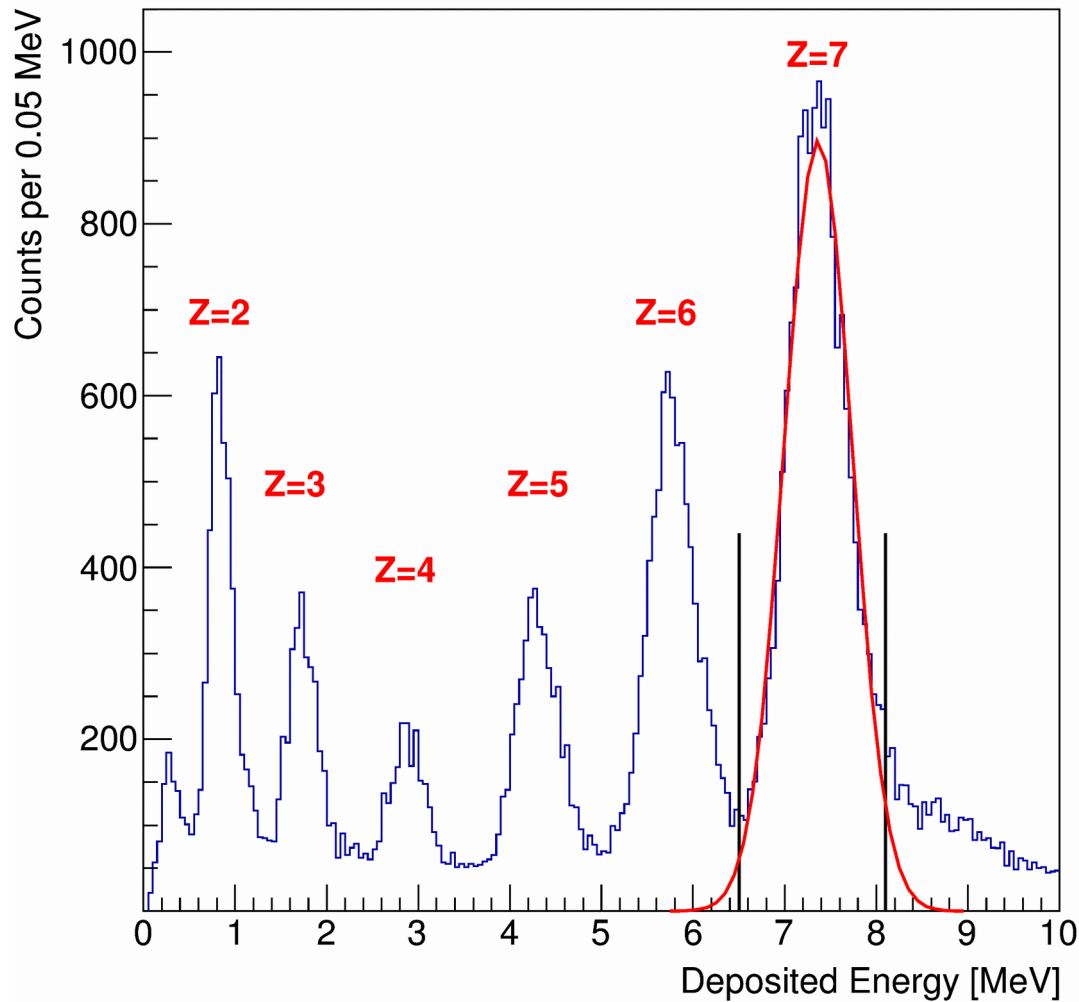
- kinematically complete measurement
- measure all outgoing particles



Incoming Particle Identification

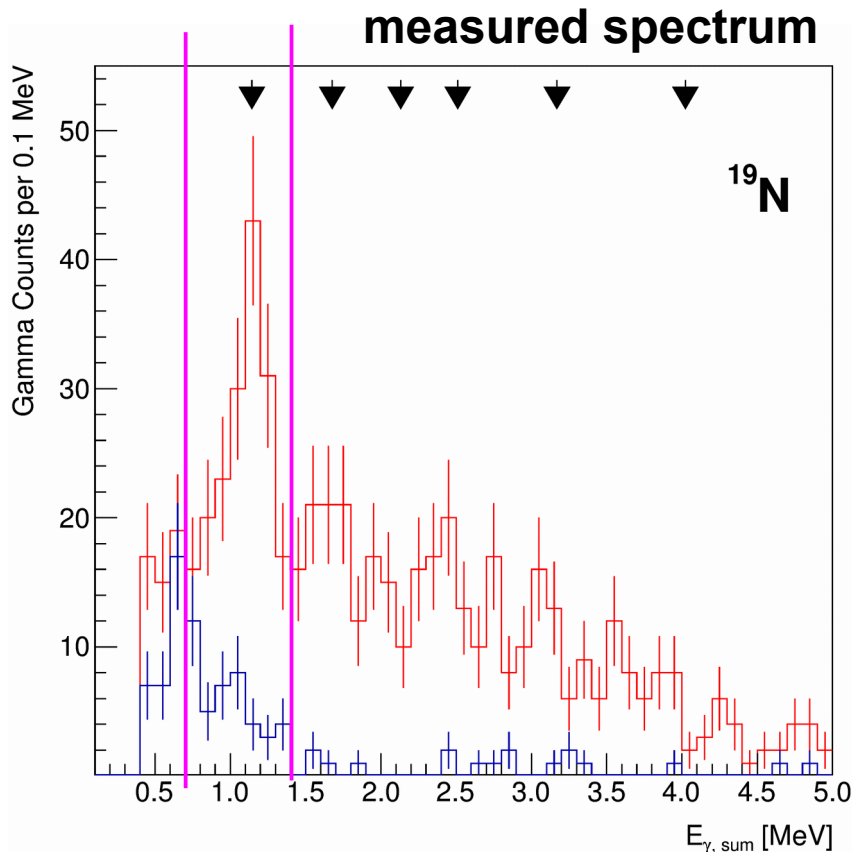


Identifying Charge of outgoing Fragment

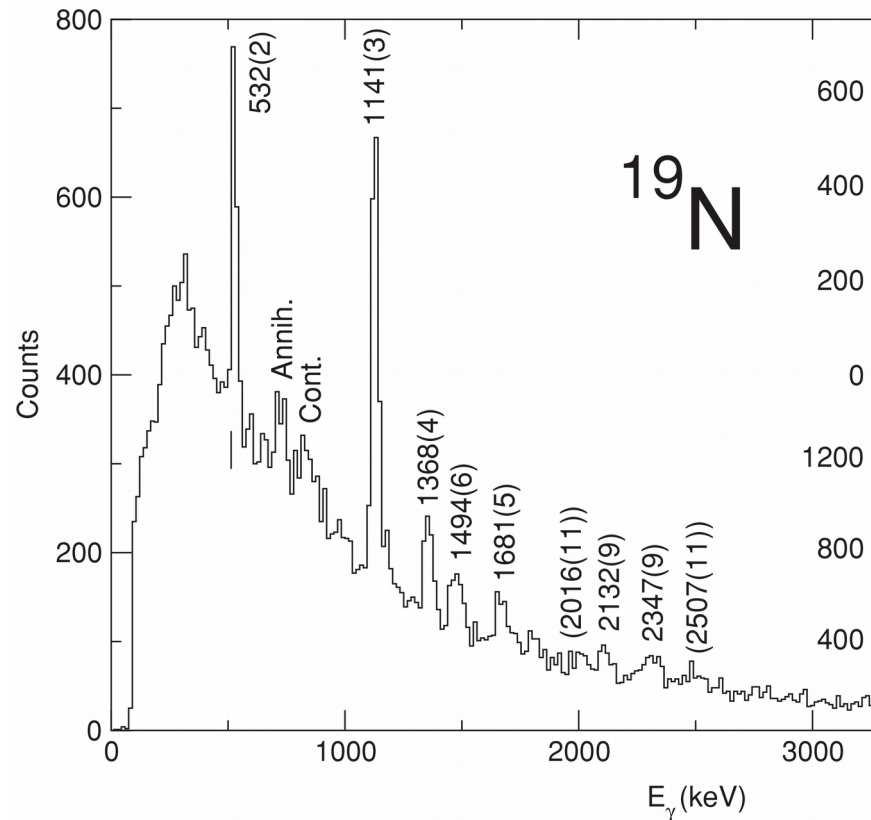


- Incoming ^{20}N
- Requiring one hit in LAND
- Accepting only one hit in Fragment ToF Wall (TFW)

Measured Gamma Spectrum



in ^{20}N , Pb-target, out ^{19}N , Hit in LAND
 Summed energy of all detected γ
 Threshold for Addback: 0.3 MeV
 Threshold for cluster summing: 0.4 MeV

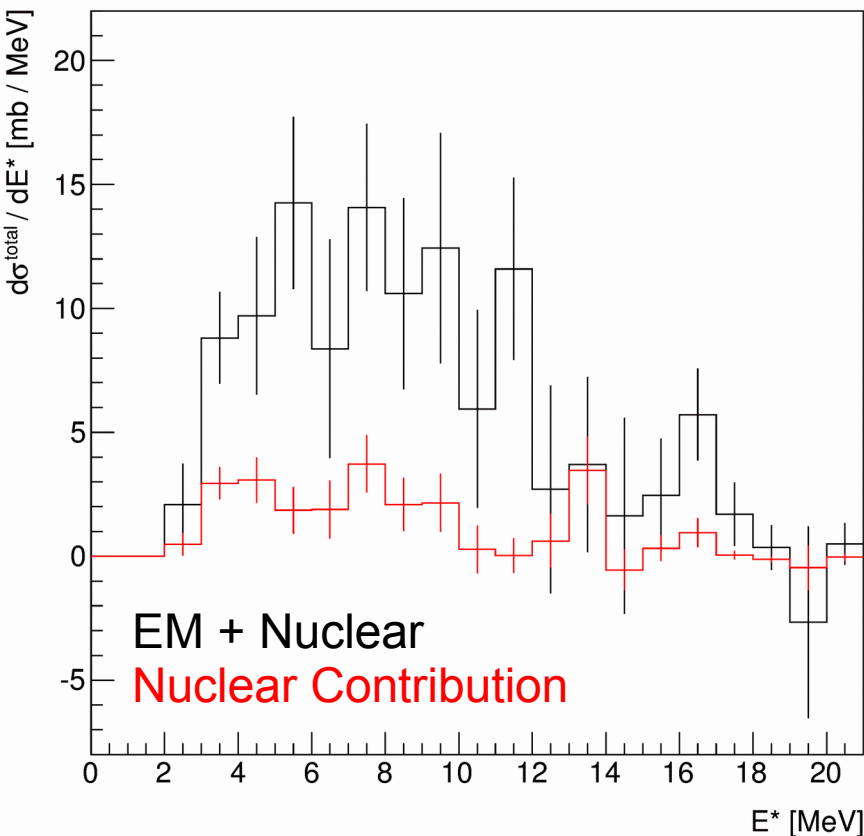


D. Sohler, et al., 2008, PRC 77, 044303
 at GANIL, in-flight spectra with Doppler-correction, Addback, Ge-detector,
 $E(^{36}\text{S})=77.5\text{MeV}$

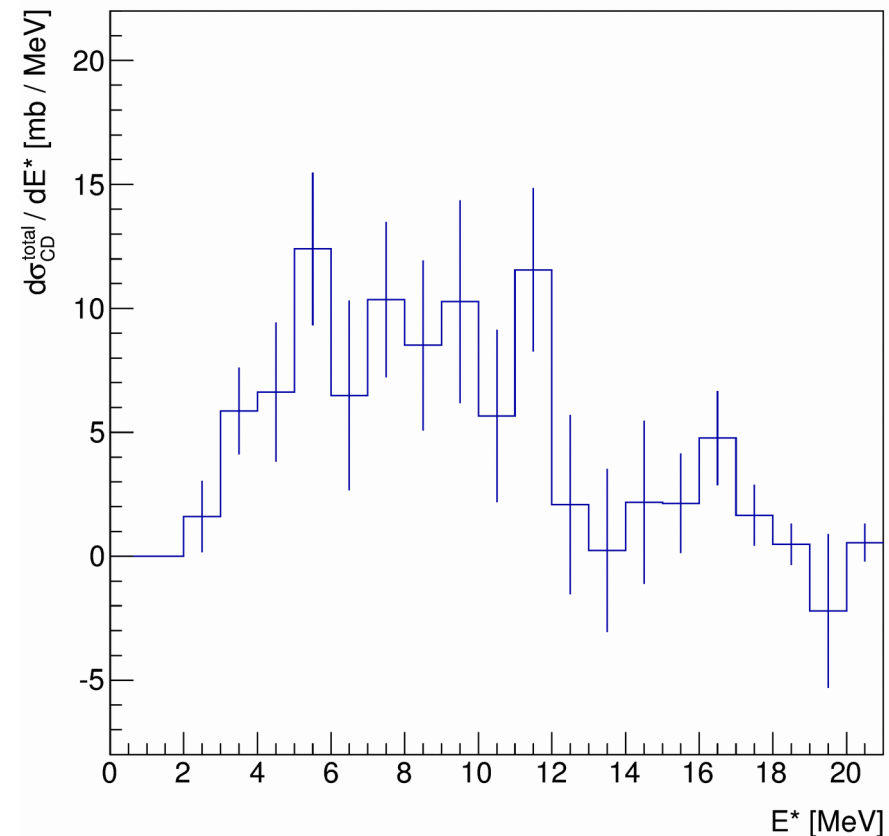
Total Excitation Energy Spectrum



Partial Contributions



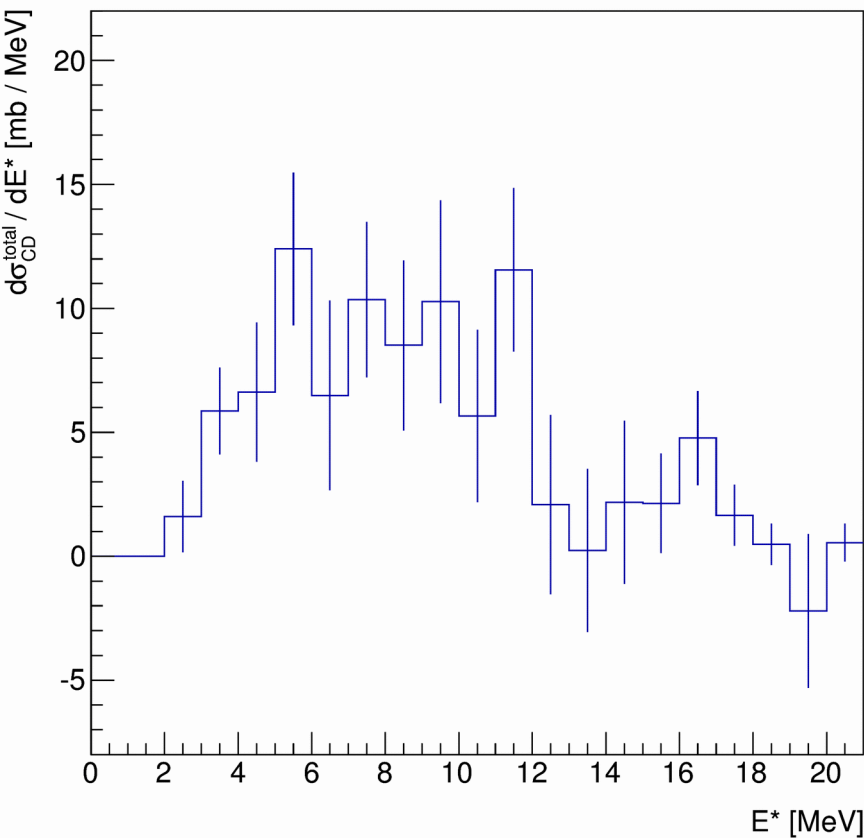
Coulomb Dissociation Cross Section



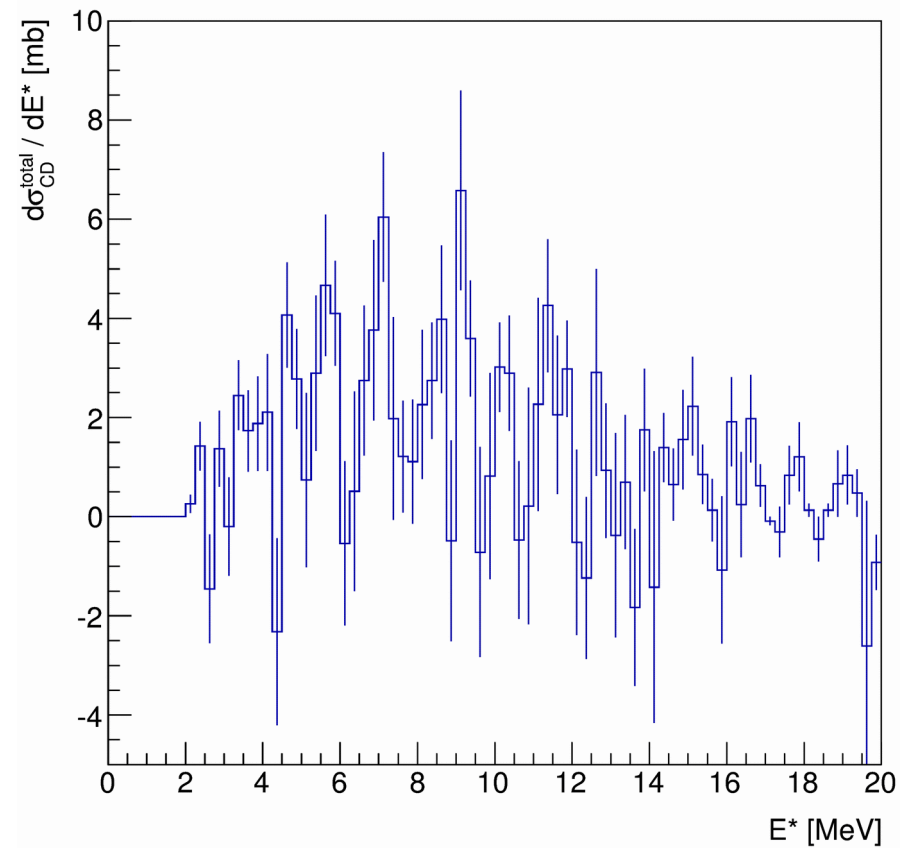
Total Excitation Energy Spectrum



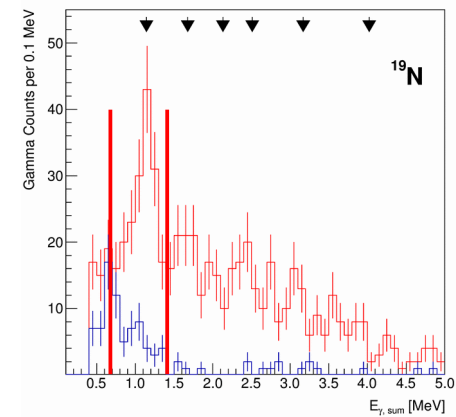
Coulomb Dissociation Cross Section



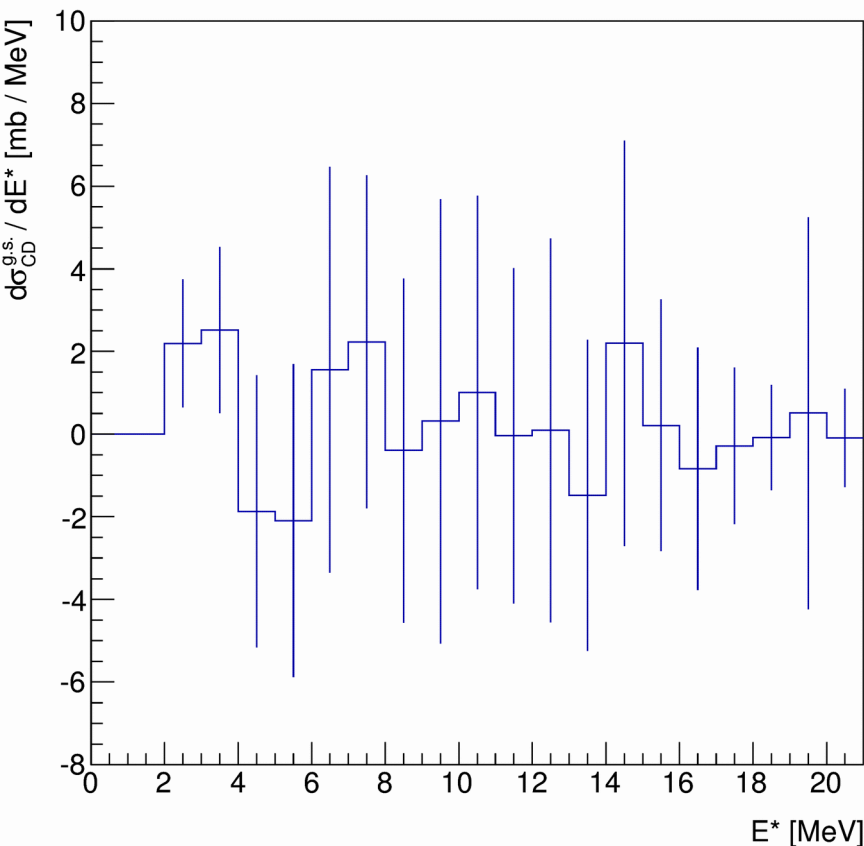
Finer Binning



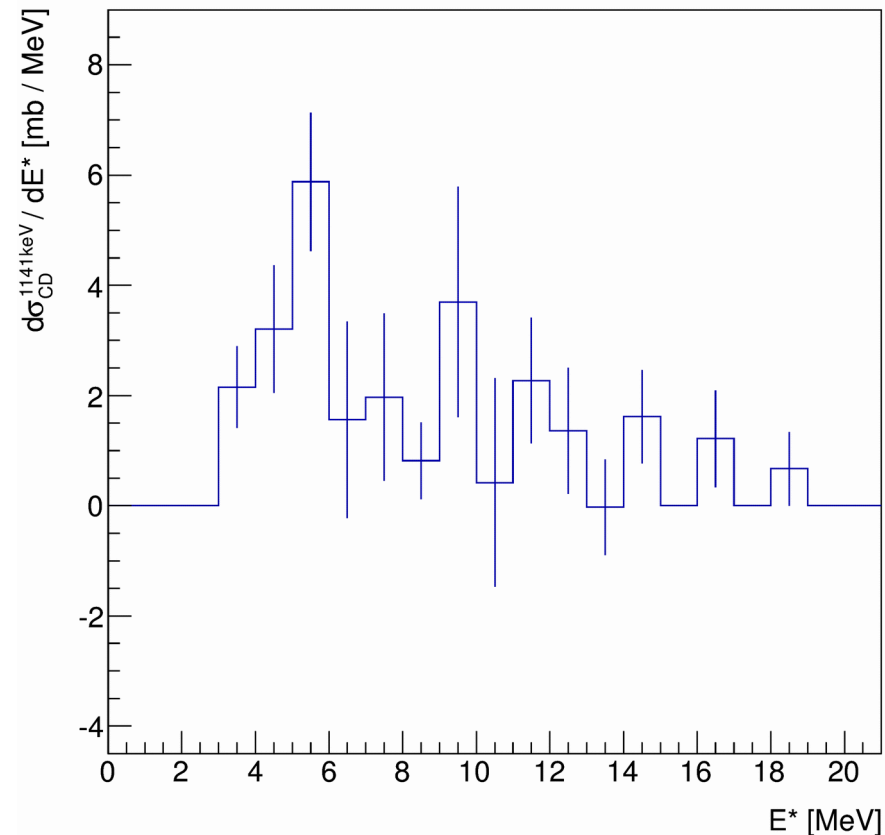
Excitation Energy Spectrum



Ground State $E(\gamma, \text{sum}) < 0.7 \text{ MeV}$



First Excited State (1141keV)



Cross Section $^{20}\text{N}(\gamma, n)^{19}\text{N}$

Uncertainties

- statistical uncertainty
- LAND efficiency 6%
- Crystal Ball eff. 5%
- incoming particle identification

$$\sigma_{\text{CD}}(\text{total}) = (96 \pm 12^{\text{stat}} \pm 4^{\text{InPID}} \pm 6^{\text{LAND}}) \text{ mb}$$

$$\sigma_{\text{CD}}(\text{all excited}) = (80 \pm 10^{\text{stat}} \pm 2^{\text{InPID}} \pm 5^{\text{LAND}} \pm 5^{\text{CB}}) \text{ mb}$$

$$\sigma_{\text{CD}}(\text{ground state}) = (16 \pm 15^{\text{stat}} \pm 1^{\text{InPID}} \pm 1^{\text{LAND}} \pm 1^{\text{CB}}) \text{ mb}$$

$$\sigma_{\text{CD}}(\text{1st excited state}) = (28 \pm 5^{\text{stat}} \pm 2^{\text{InPID}} \pm 2^{\text{LAND}} \pm 2^{\text{CB}}) \text{ mb}$$

Photo Absorption Cross Section

$$\sigma_{\gamma,n} \equiv \sigma_{E1}^{\text{photo}} = \frac{d\sigma_{CD}}{dE^*} \frac{1}{n_{E1}(E^*)} E^* \quad E_\gamma = E^*$$

$\sigma[{}^{20}\text{N}(\gamma,n){}^{19}\text{N}^{\text{g.s.}}] < 9.2 \text{ mb}$ for $E^* = (0 \text{ MeV}, 15 \text{ MeV})$; 90% confidence level

Invariance under time reversal via principle of detailed balance

Neutron Capture Cross Section

$$\sigma_{n,\gamma} = \frac{2(2J_A + 1)}{(2J_{A-1} + 1)(2J_n + 1)} \frac{k_\gamma^2}{k_{c.m.}^2} \sigma_{\gamma,n}$$

$$k_\gamma = E^* / (\hbar c)$$

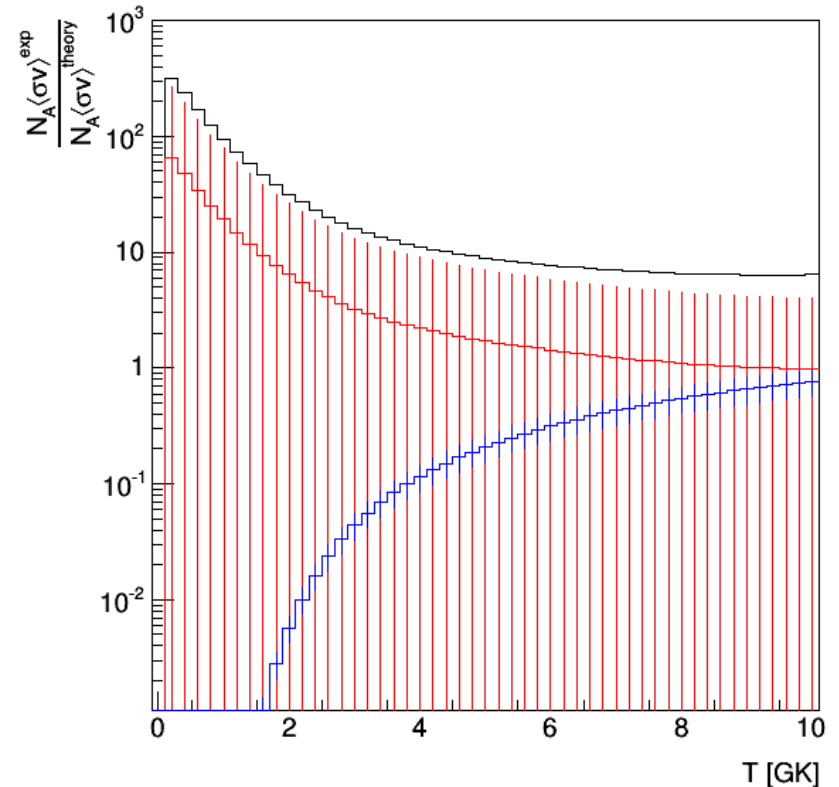
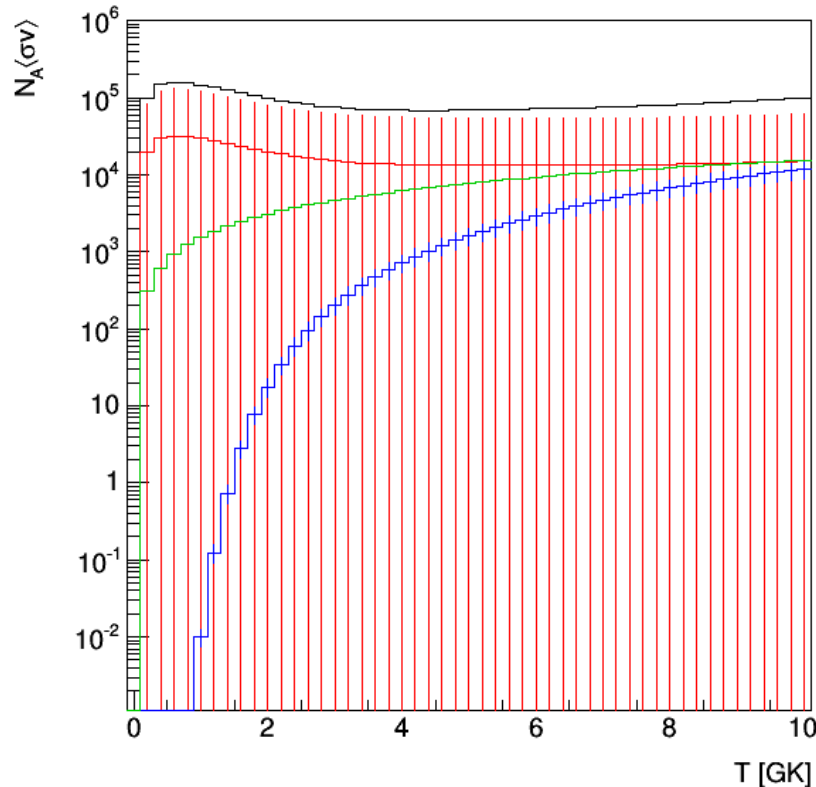
$$k_{c.m.}^2 = 2\mu(E^* - S_{1n}) / \hbar^2$$

$$\mu = (M_{19\text{N}} \cdot M_n) / (M_{19\text{N}} + M_n)$$

$\sigma[{}^{19}\text{N}^{\text{g.s.}}(n,\gamma){}^{20}\text{N}] < 0.24 \text{ mb}$ for $E^* = (0 \text{ MeV}, 15 \text{ MeV})$; 90% confidence level

Maxwellian Average Reaction Rates

Scaled to theoretic data



$^{19}\text{N}[\text{g.s.}](n,\gamma)^{20}\text{N}$

$^{19}\text{N}[\text{1st excited state}](n,\gamma)^{20}\text{N}$ including population probability

$^{19}\text{N}(n,\gamma)^{20}\text{N}$ upper limit of g.s. and 1st excited state transitions

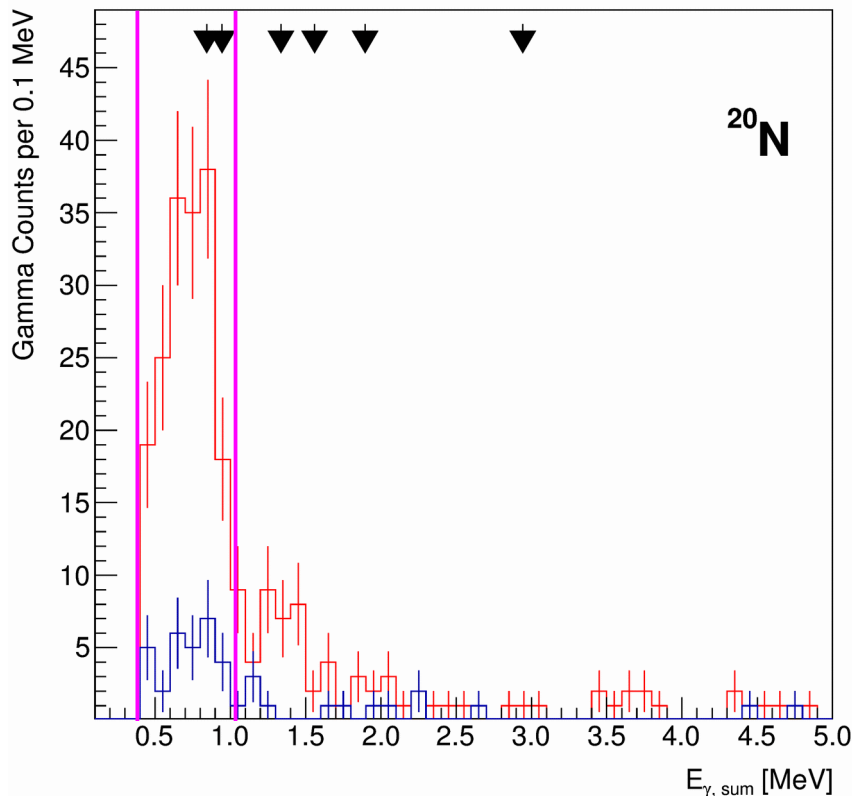
$^{19}\text{N}(n,\gamma)^{20}\text{N}$ Rauscher et al. 1994 (theoretic)



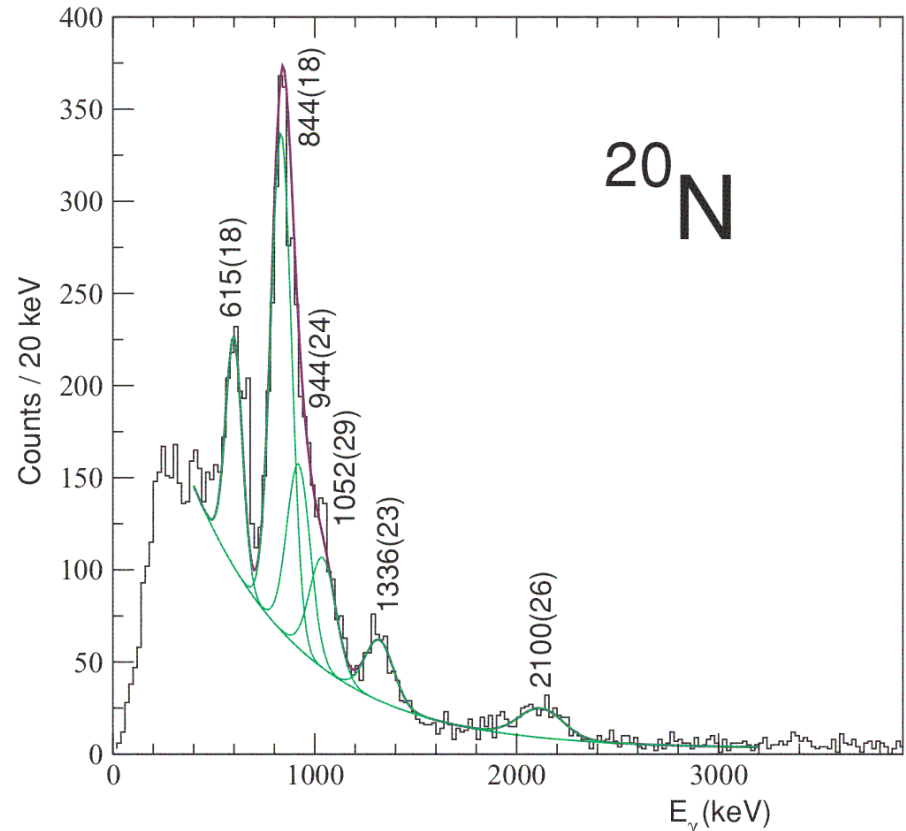
Measured Gamma Spectrum



measured spectrum



in ^{21}N , Pb-target, out ^{20}N , Hit in LAND
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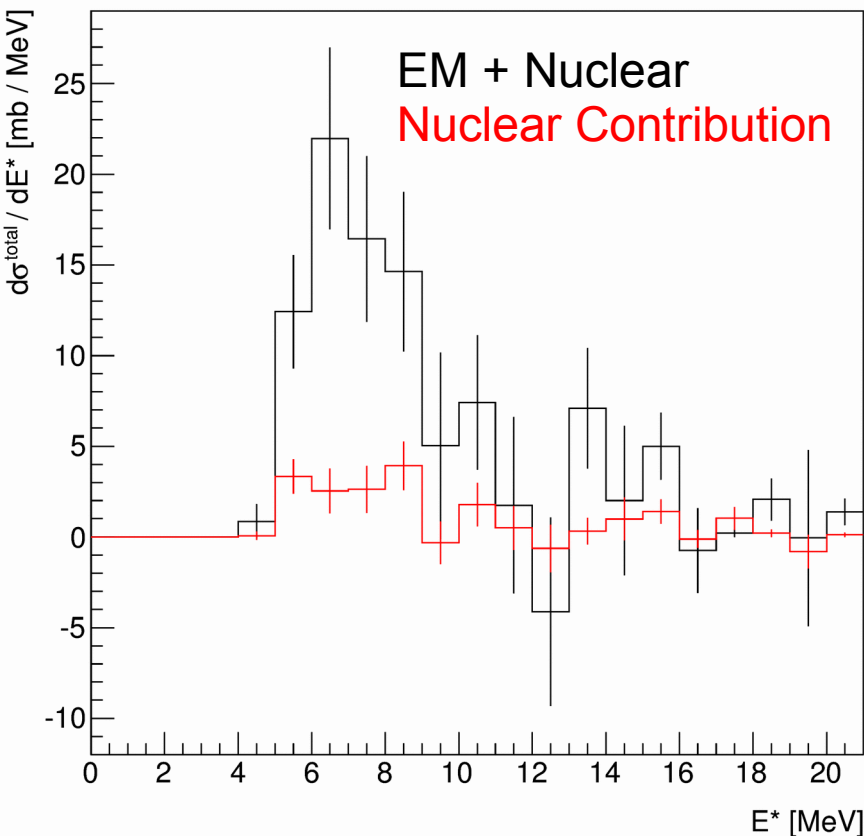


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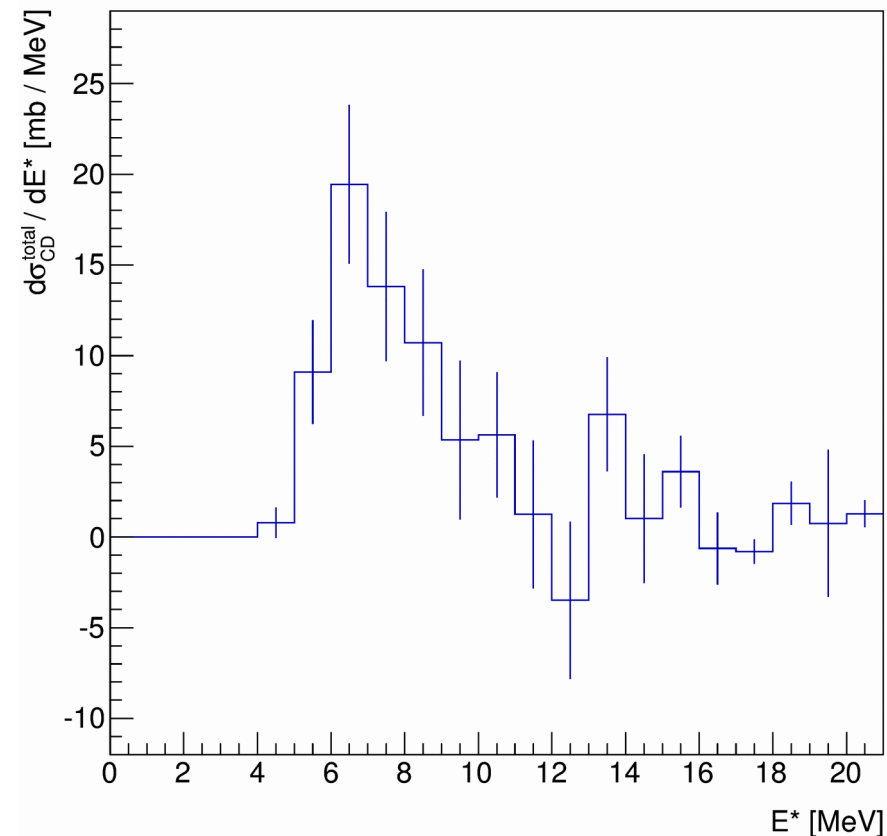
Total Excitation Energy Spectrum



Partial Contributions



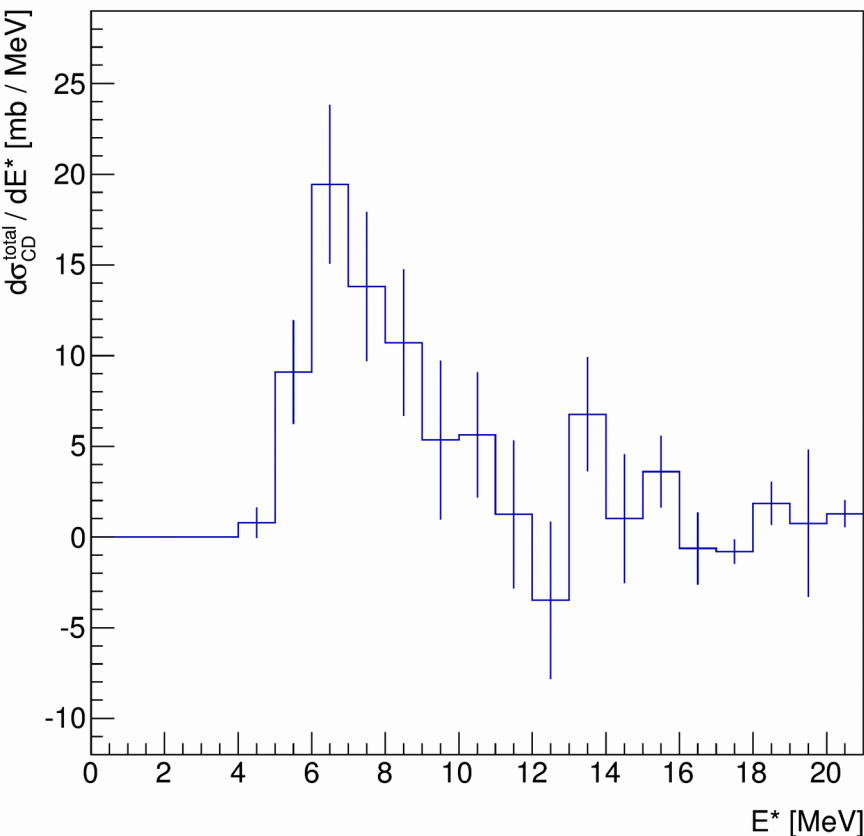
Coulomb Dissociation Cross Section



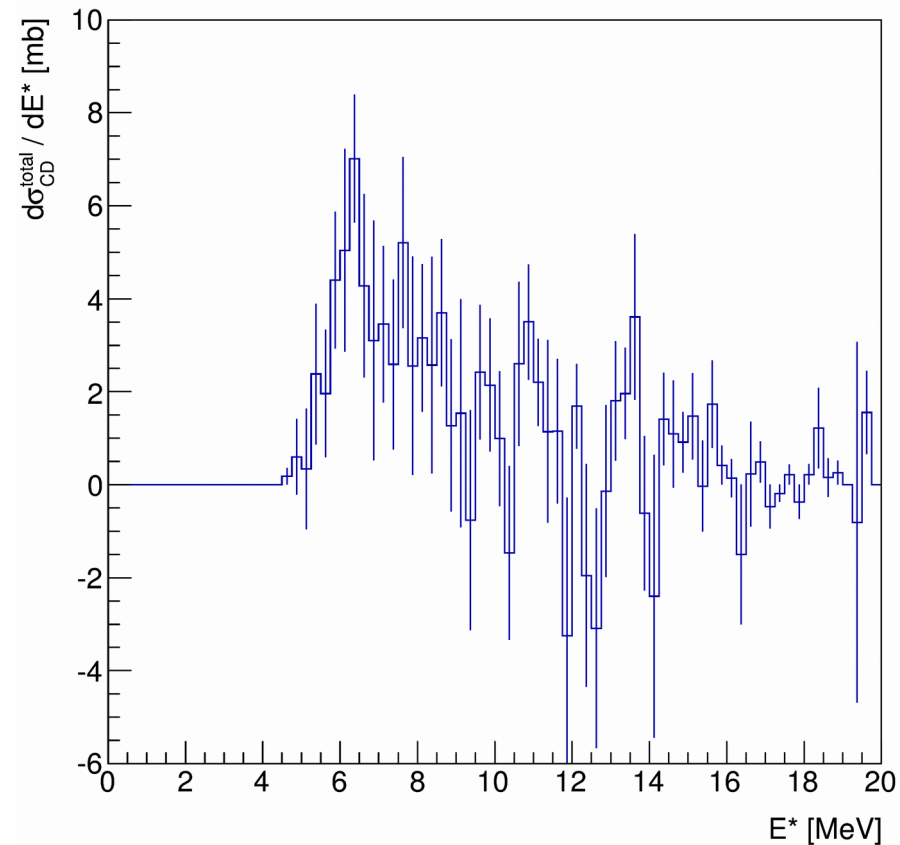
Total Excitation Energy Spectrum



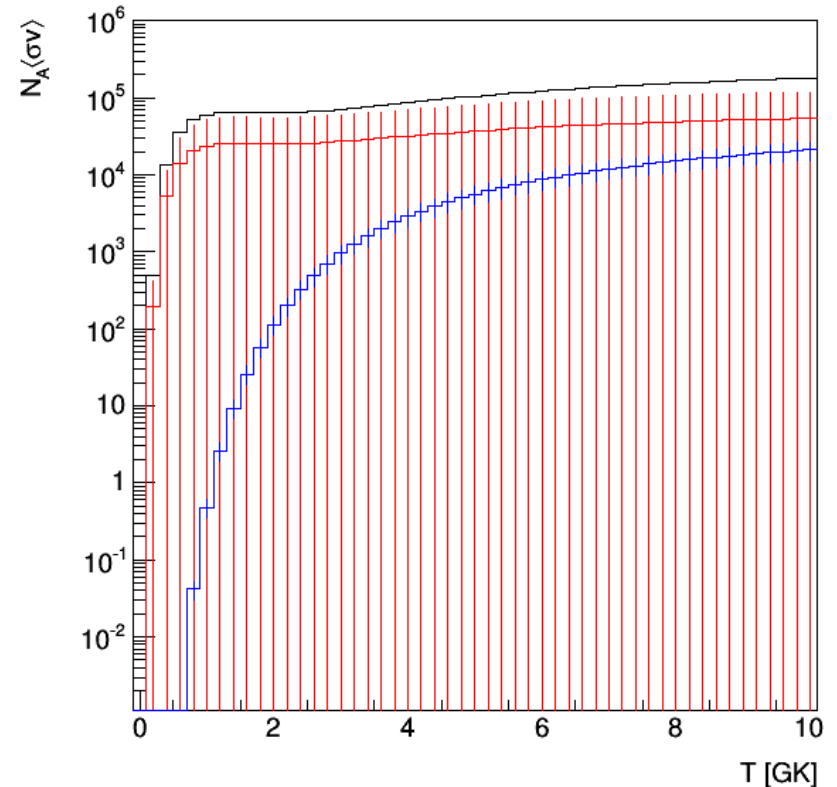
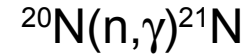
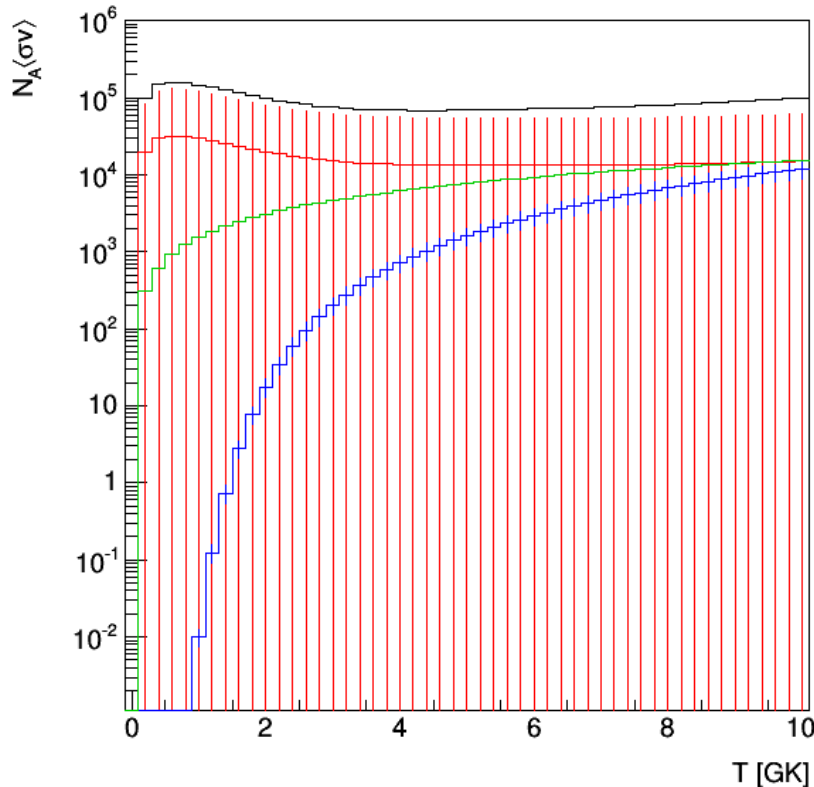
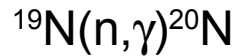
Coulomb Dissociation Cross Section



Finer Binning



Maxwellian Average Reaction Rates



$^{A}\text{N}[\text{g.s.}](n,\gamma)^{A+1}\text{N}$

$^{A}\text{N}[\text{1st excited state}](n,\gamma)^{A+1}\text{N}$ including population probability

$^{A}\text{N}(n,\gamma)^{A+1}\text{N}$ upper limit of g.s. and 1st excited state transitions

$^{A}\text{N}(n,\gamma)^{A+1}\text{N}$ Rauscher et al. 1994 (theoretic)

Cross Section $^{21}\text{N}(\gamma, n)^{20}\text{N}$

$$\sigma_{\text{CD}}(\text{total}) = (84 \pm 12^{\text{stat}} \pm 5^{\text{inPID}} \pm 5^{\text{LAND}}) \text{ mb}$$

$$\sigma_{\text{CD}}(\text{all excited}) = (53 \pm 9^{\text{stat}} \pm 0^{\text{inPID}} \pm 3^{\text{LAND}} \pm 3^{\text{CB}}) \text{ mb}$$

$$\sigma_{\text{CD}}(\text{ground state}) = (31 \pm 15^{\text{stat}} \pm 5^{\text{inPID}} \pm 2^{\text{LAND}} \pm 2^{\text{CB}}) \text{ mb}$$

Photo Absorption Cross Section

$$\sigma[^{21}\text{N}(\gamma, n)^{20}\text{N}^{\text{g.s.}}] < 10.8 \text{ mb} \quad \text{for } E^* = (0 \text{ MeV}, 15 \text{ MeV}); 90\% \text{ confidence level}$$

Neutron Capture Cross Section

$$\sigma[^{20}\text{N}^{\text{g.s.}}(n, \gamma)^{21}\text{N}] < 0.13 \text{ mb} \quad \text{for } E^* = (0 \text{ MeV}, 15 \text{ MeV}); 90\% \text{ confidence level}$$

Summary

- Measured astrophysically relevant reaction via time inversion by principle of detailed balance
- Coulomb Excitation Cross Section for $^{20}\text{N}(\gamma,n)^{19}\text{N}$
 - $\sigma_{\text{CD}}(\text{total}) = 96 \text{ mb}$, $u_{\text{stat}} = \pm 12 \text{ mb}$, $u_{\text{sys,inPID}} = \pm 4 \text{ mb}$, $u_{\text{sys,LAND}} = \pm 6 \text{ mb}$
- Neutron Capture Cross Section: $\sigma[^{20}\text{N}(\gamma,n)^{19}\text{N}^{\text{g.s.}}] < 9.2 \text{ mb}$
- Coulomb Excitation Cross Section for $^{21}\text{N}(\gamma,n)^{20}\text{N}$
 - $\sigma_{\text{CD}}(\text{total}) = 84 \text{ mb}$, $u_{\text{stat}} = \pm 12 \text{ mb}$, $u_{\text{sys,inPID}} = \pm 5 \text{ mb}$, $u_{\text{sys,LAND}} = \pm 5 \text{ mb}$
- Neutron Capture Cross Section: $\sigma[^{20}\text{N}^{\text{g.s.}}(n,\gamma)^{21}\text{N}] < 0.13 \text{ mb}$
- Discussion & Outlook
 - include in Reaction Network Code for impact on r-process abundance
 - repeat with higher statistics (FAIR)

